

Reheating after Inflation, Gauge fields and GWs



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ILLINOIS
Physics

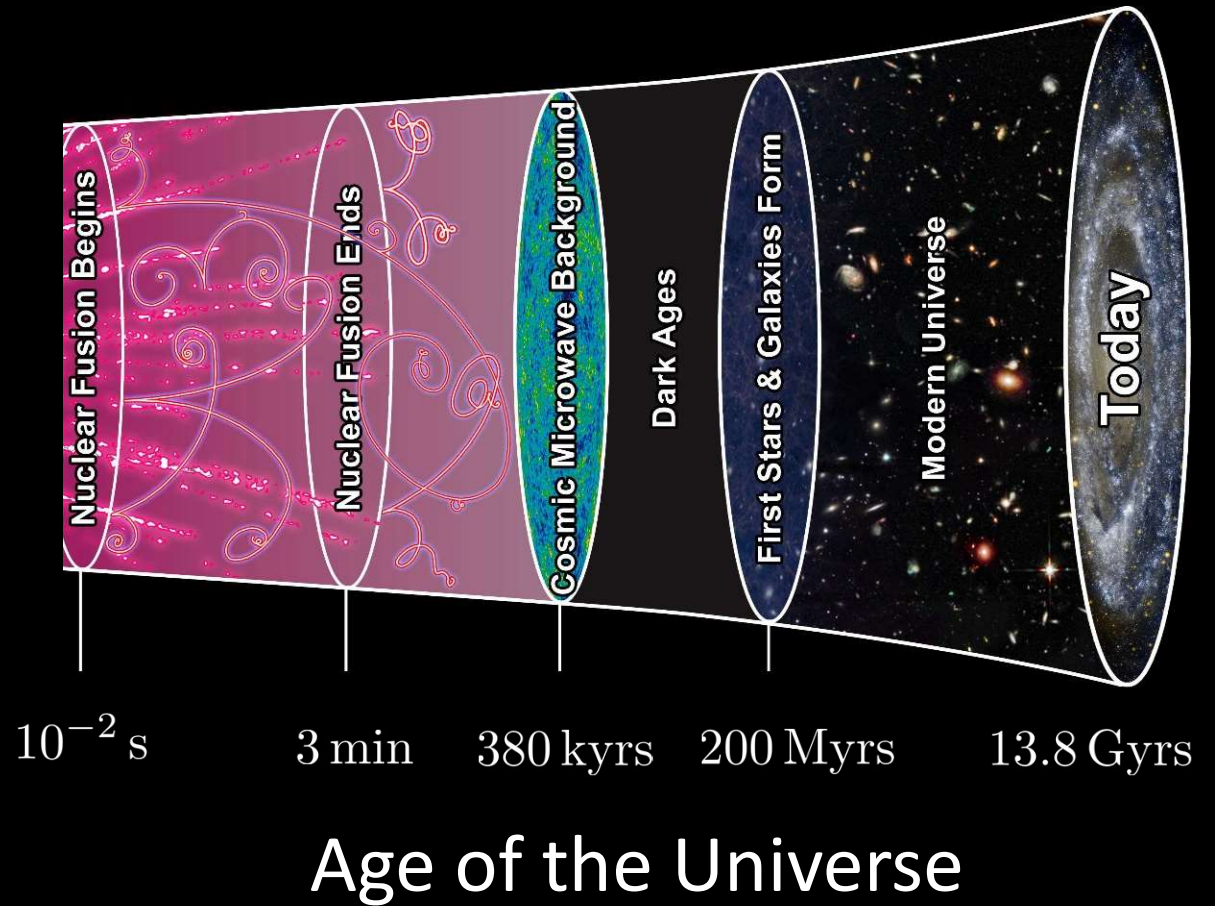
GRAINGER COLLEGE OF ENGINEERING

Collaborators: Peter Adshead, Zachary Weiner, Mustafa Amin, Eiichiro Komatsu, Matthew Reece, Jiji Fan, Leila Mirzagoli, Yuki Watanabe, Ippei Obata...

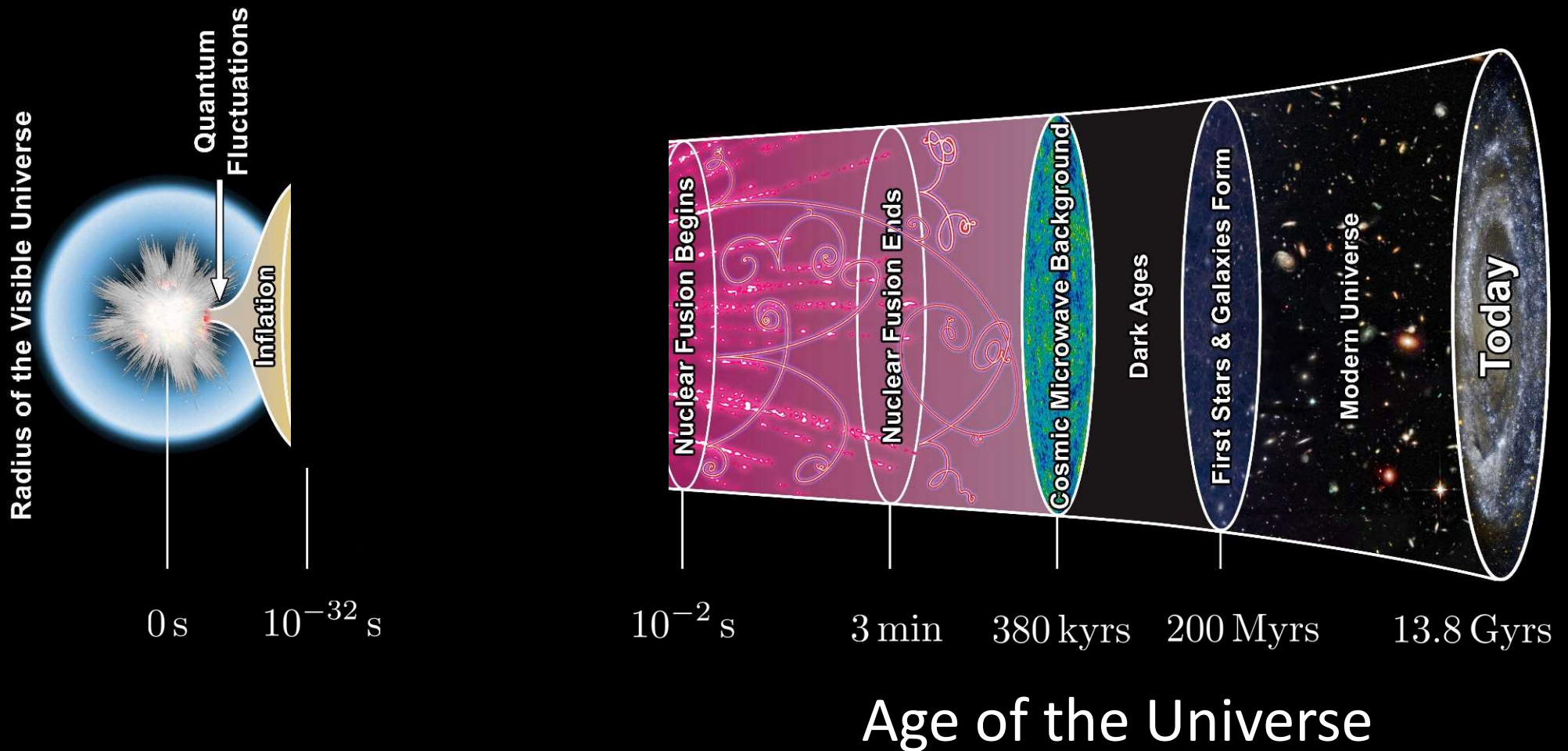
Overview

History of the Universe

Radius of the Visible Universe

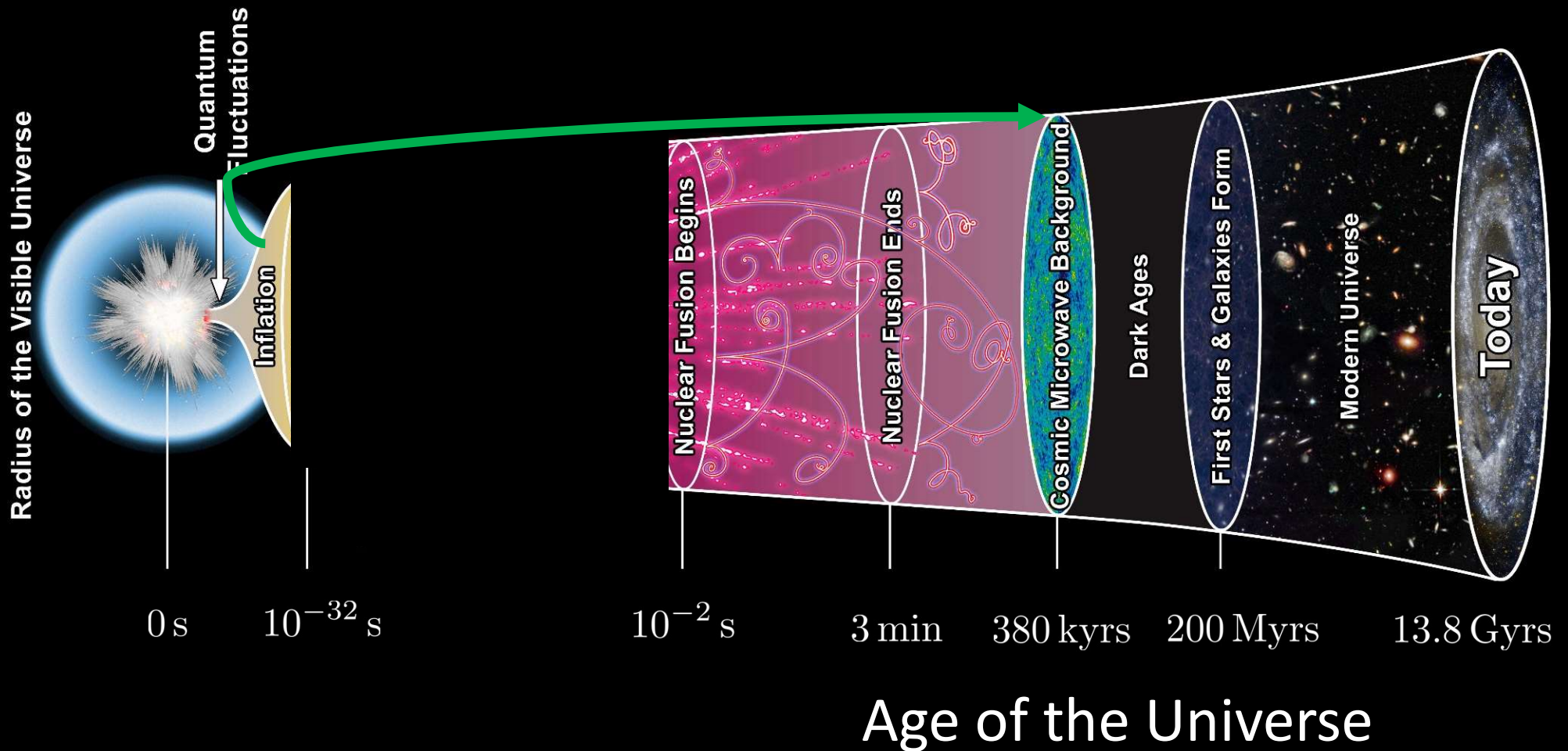


History of the Universe

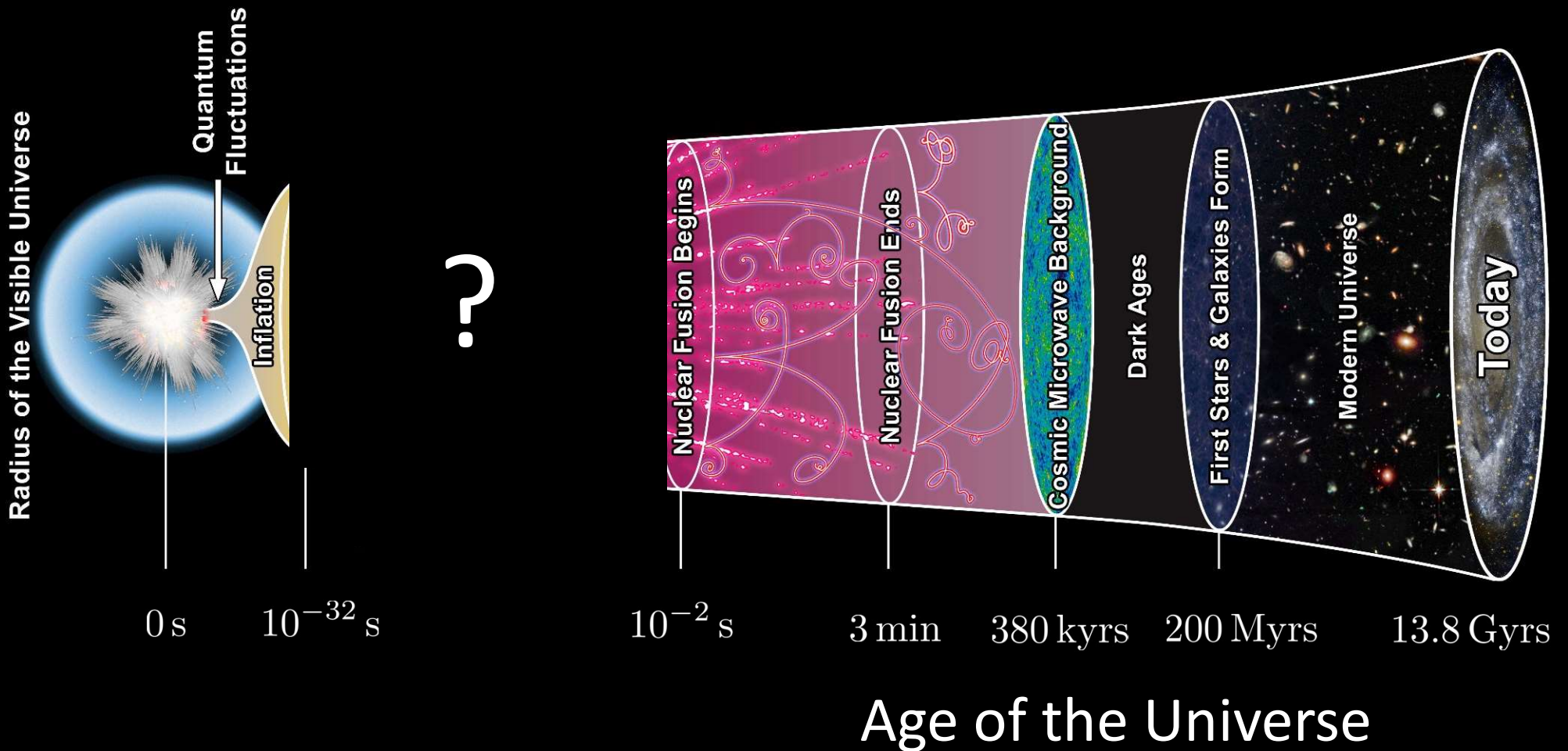


History of the Universe

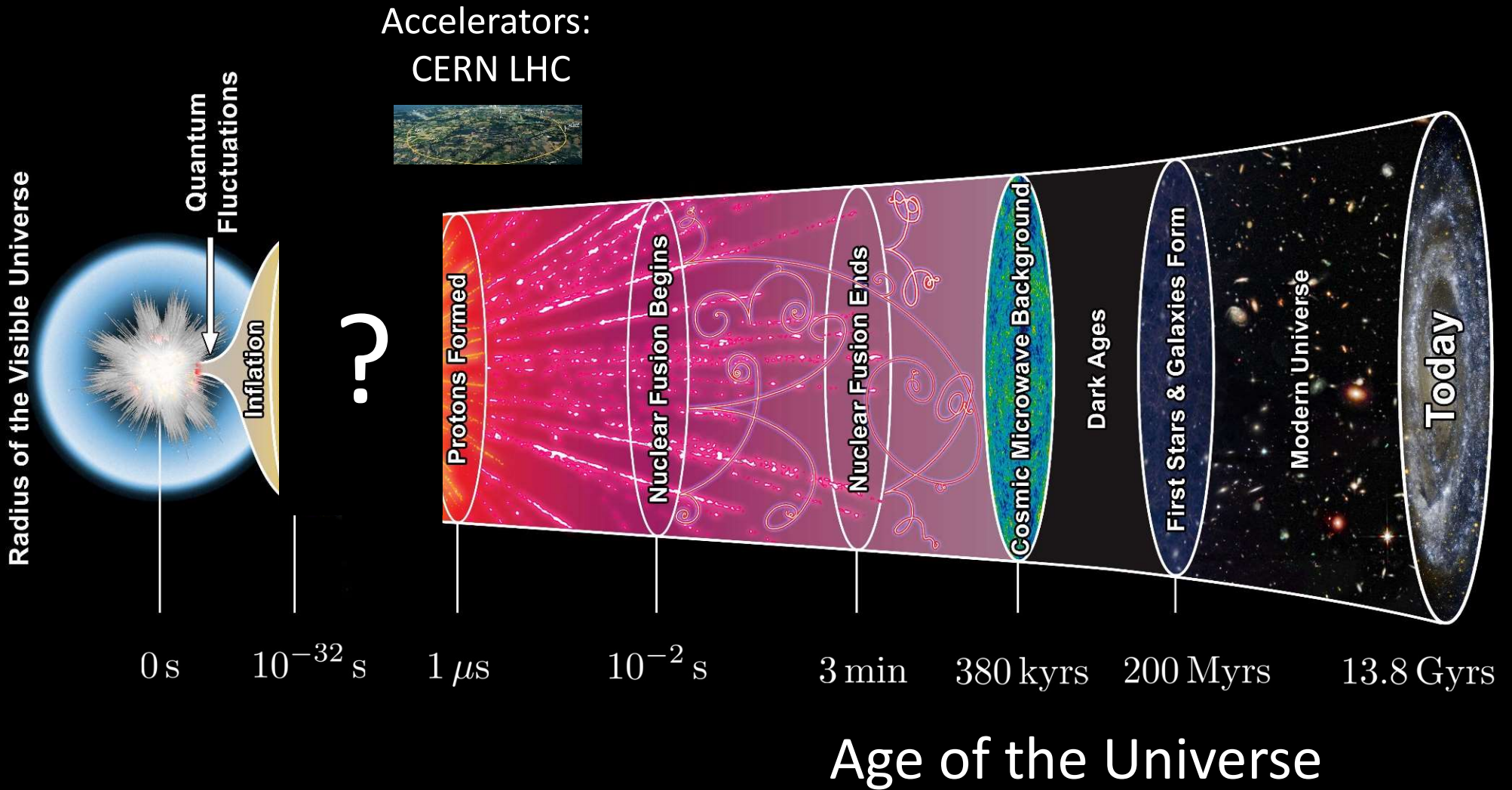
Mukhanov, Chibisov (1981)



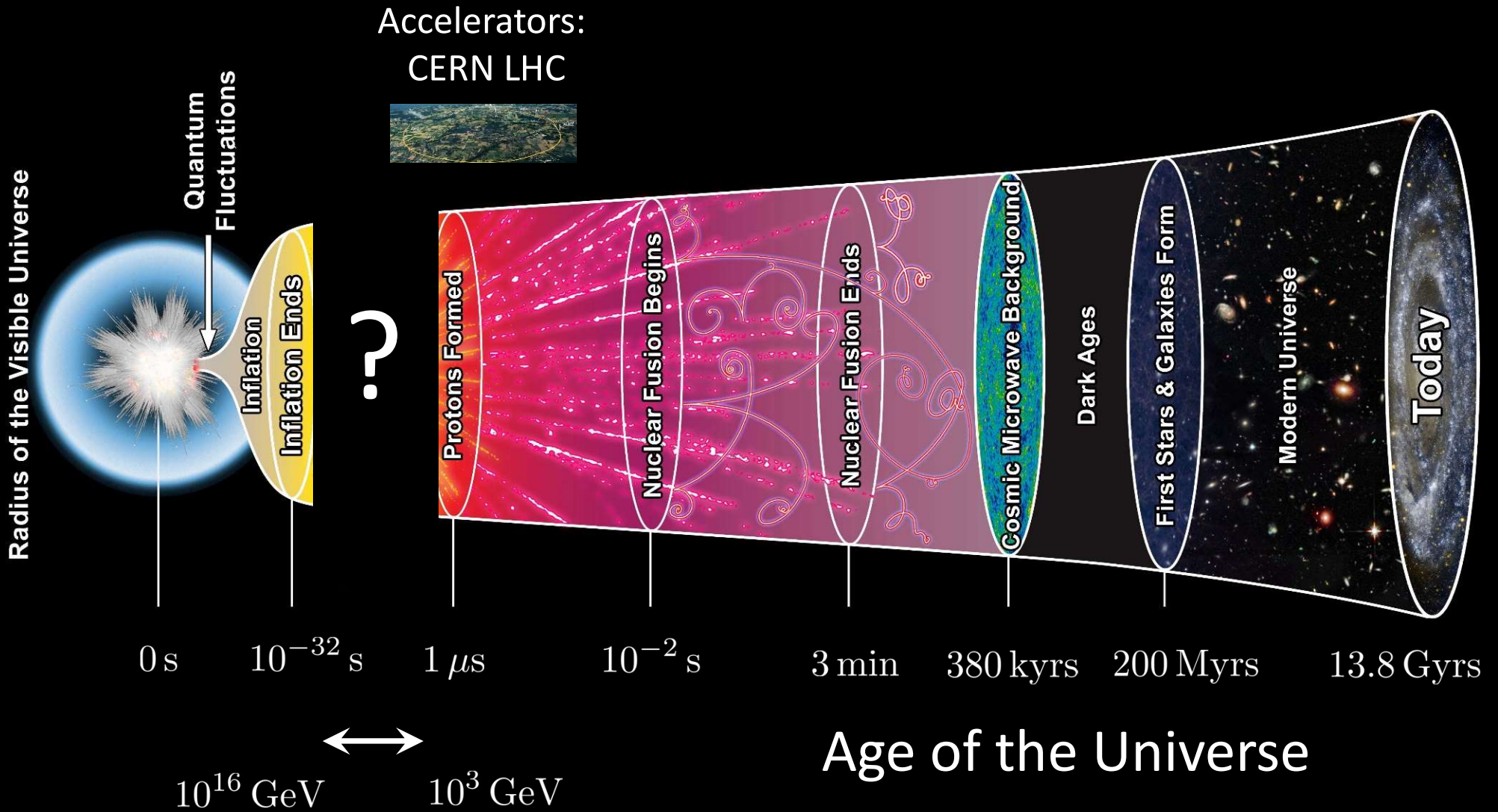
History of the Universe



History of the Universe



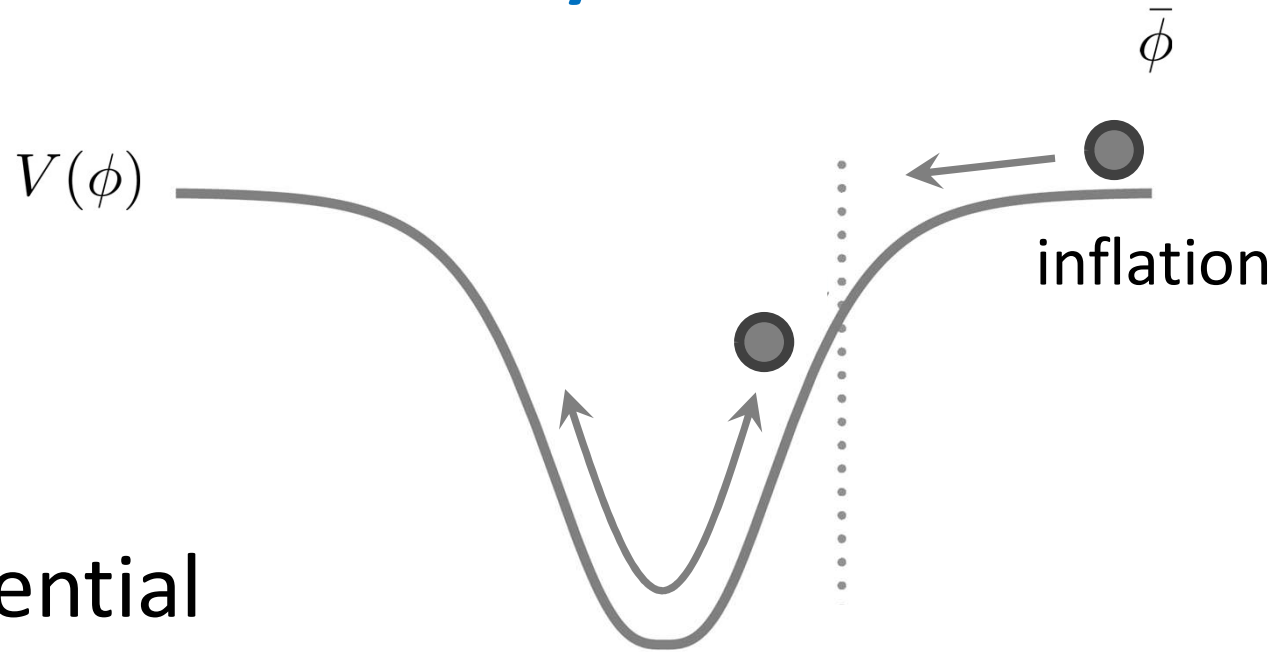
History of the Universe



The Universe at the end of inflation

- Cold, dominated by inflaton $\phi(t)$
- $\phi(t)$ should decay to matter and radiation to recover BBN - *REHEATING*

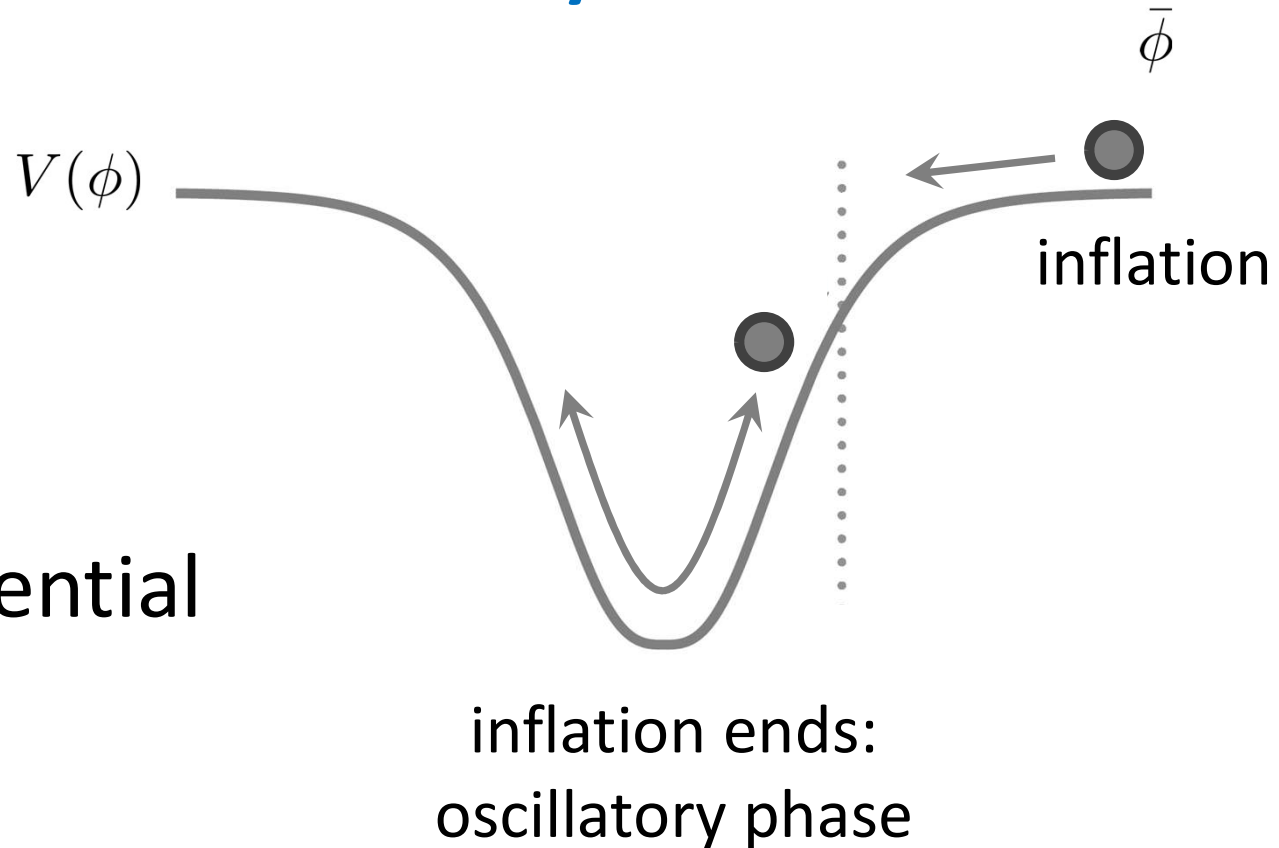
Inflaton decay



- Form of the potential

inflation ends:
oscillatory phase

Inflaton decay



- Form of the potential
- Couplings to other fields

Inflaton decay

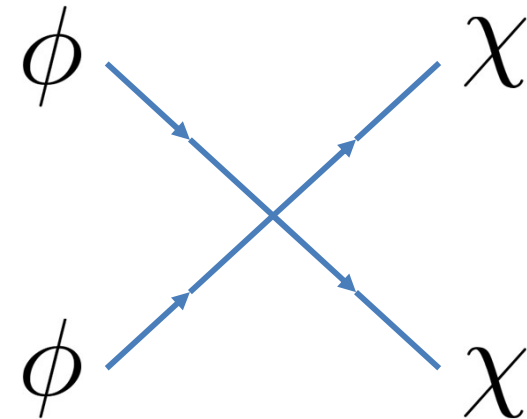
Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

- Perturbative

Inflaton decay

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

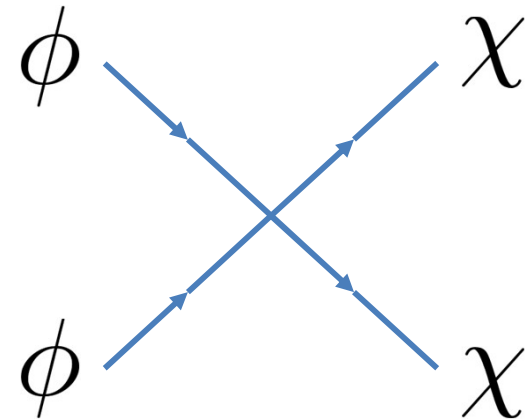
- Perturbative
-Rate and x-section



Inflaton decay

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

- Perturbative
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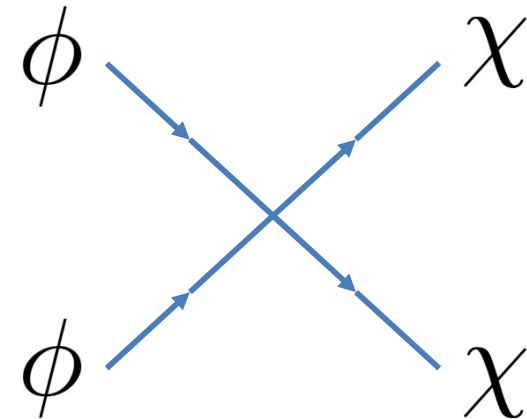


- Non-perturbative (explosive)

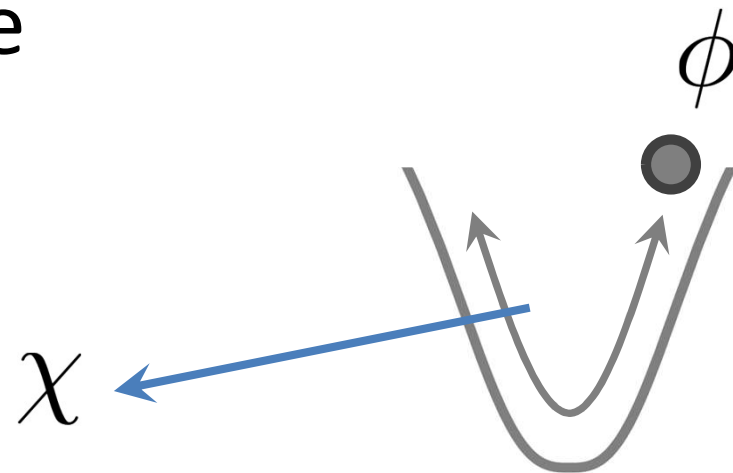
Inflaton decay

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

- Perturbative
-Rate and x-section



- Non-perturbative (explosive)
-parametric resonance



Non-perturbative decay (parametric resonance)

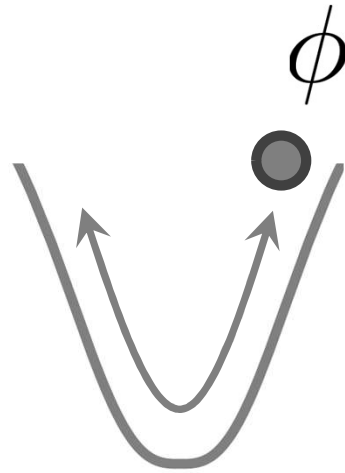
$$V(\phi, \chi) \sim \frac{1}{2}m^2\phi^2 + g^2\phi^2\chi^2 + \dots$$

$$\ddot{\chi} - \Delta\chi + g^2\phi^2\chi = 0$$

Non-perturbative decay (parametric resonance)

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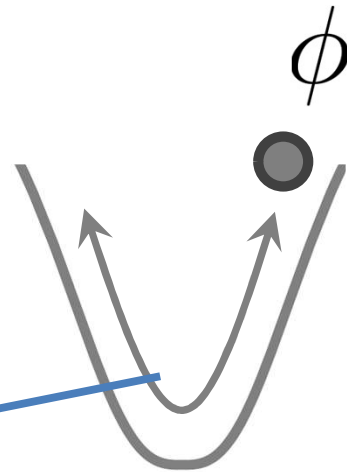


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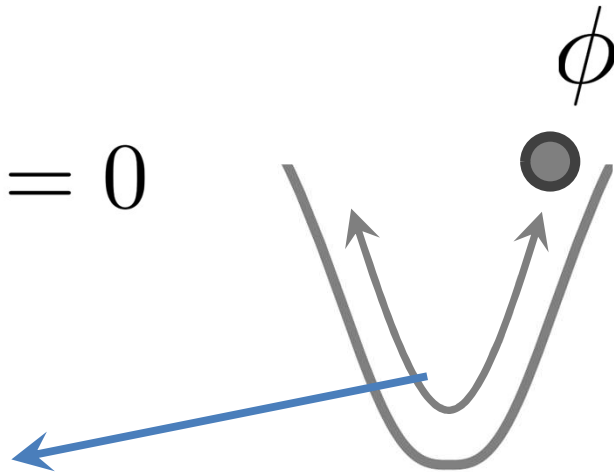


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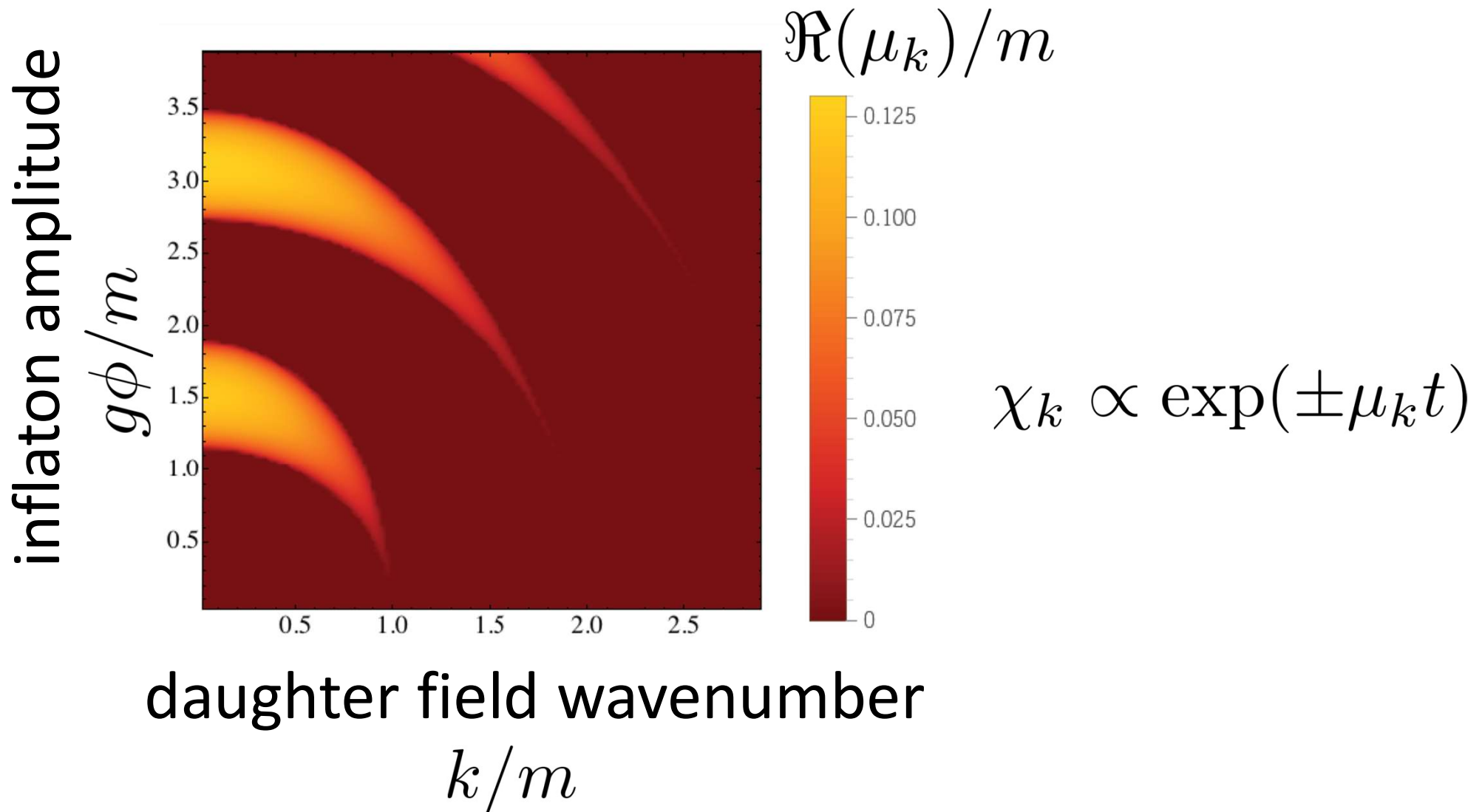
$$\ddot{\chi} - \Delta\chi + g^2\phi^2\chi = 0$$

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if $\Re(\mu_k) \neq 0 \implies$ parametric resonance

Non-perturbative decay (parametric resonance)



Time-line of Reheating

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

time



Time-line of Reheating

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

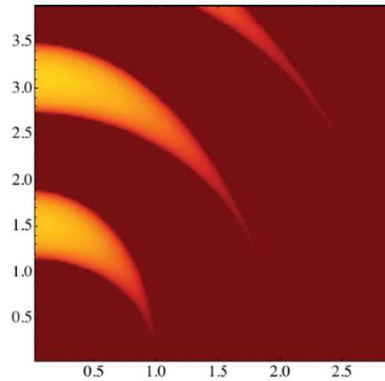
time



inflation

Time-line of Reheating

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)



time

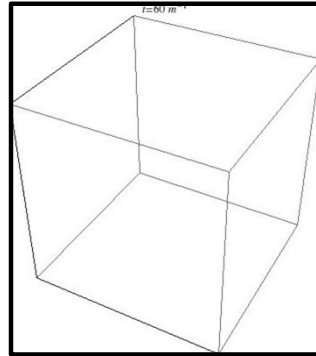
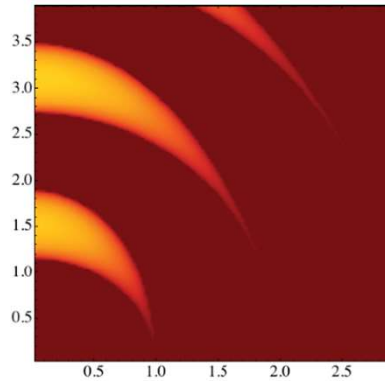


linear regime
(preheating)

inflation

Time-line of Reheating

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)



time

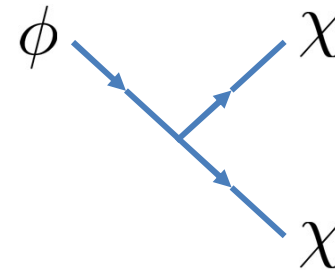
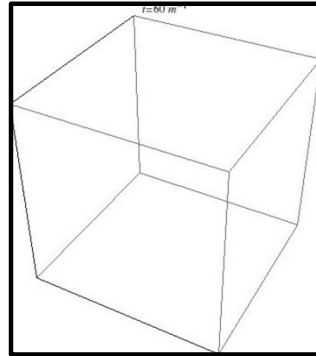
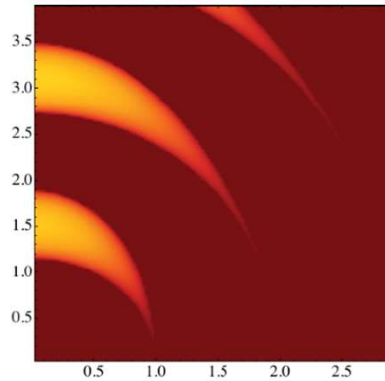
linear regime
(preheating)

non-linear
regime

inflation

Time-line of Reheating

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)



time



linear regime
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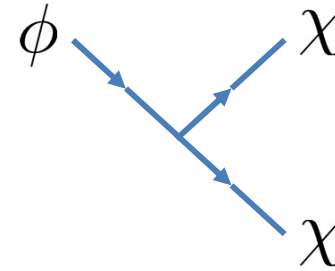
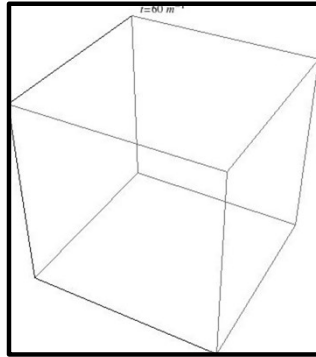
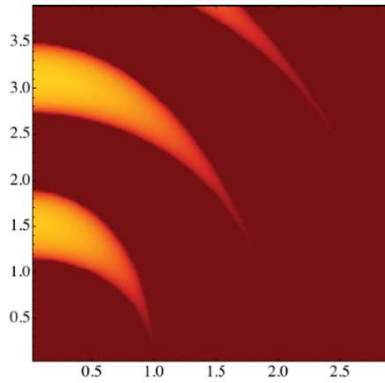
non-linear
regime

perturbative
regime

inflation

Time-line of Reheating

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)



time

inflation

linear regime
(preheating)

non-linear
regime

perturbative
regime

thermalization

nucleosynthesis

Observational signatures

- Difficult
 - very small length scales
 - thermalization

Observational consequences

Allahverdi, Brandenberger, et al. (2010), Amin, et al. (2015), KL (2019) (reviews)

- Expansion **history** effects
- Potentially observable **remnants**:
 - Stochastic GWs
 - Non-gaussianities
 - Solitons
- Constraints on **high energy** physics models:
 - Overproduction of defects
 - Complete decay of $\bar{\phi}(t)$
 - Primordial BHs

The way forward

- **Theory**
 - more realistic models (couplings)
- **Phenomenology**
 - detailed (numerical) analysis
 - observational signatures
- **Observations**
 - non-Gaussianities, high-freq GWs
 - expansion history effects

Our works

The equation of state after inflation

KL and M. Amin, PRD 99 123504 (2019)
KL and M. Amin, PRD 97 023533 (2018)
KL and M. Amin, PRL 119 061301 (2017)

What is w after inflation?

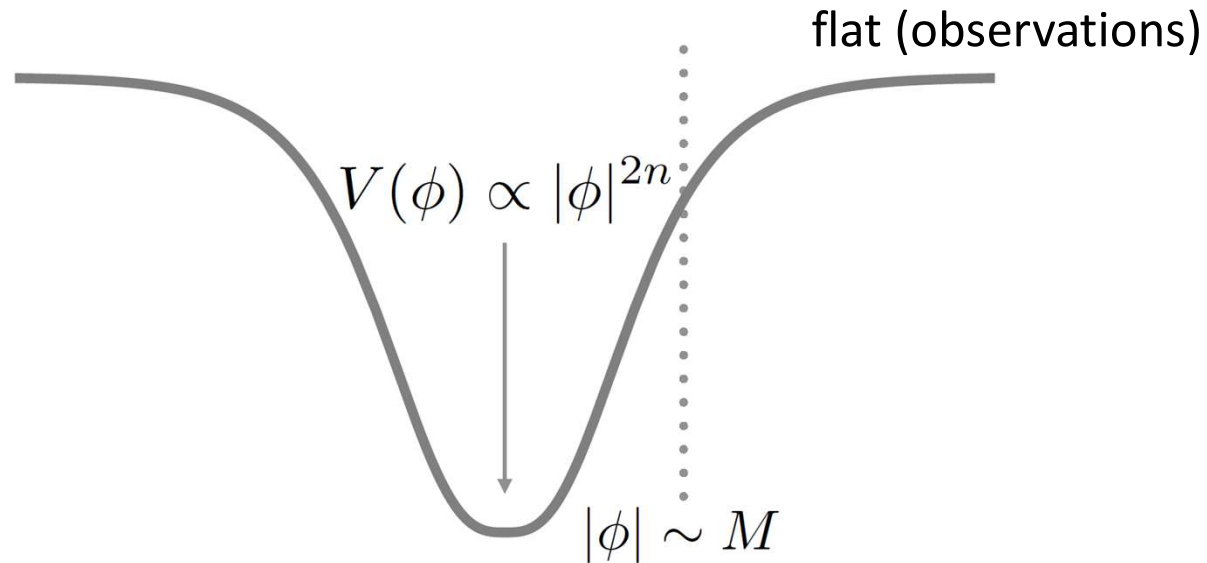
$$w = \frac{\text{pressure}}{\text{energy density}}$$

$$S = \int d^4x \sqrt{-g} \left[\frac{1}{16\pi G} \mathcal{R} - \frac{1}{2} \partial_\mu \phi \partial^\mu \phi - V(\phi) \right]$$

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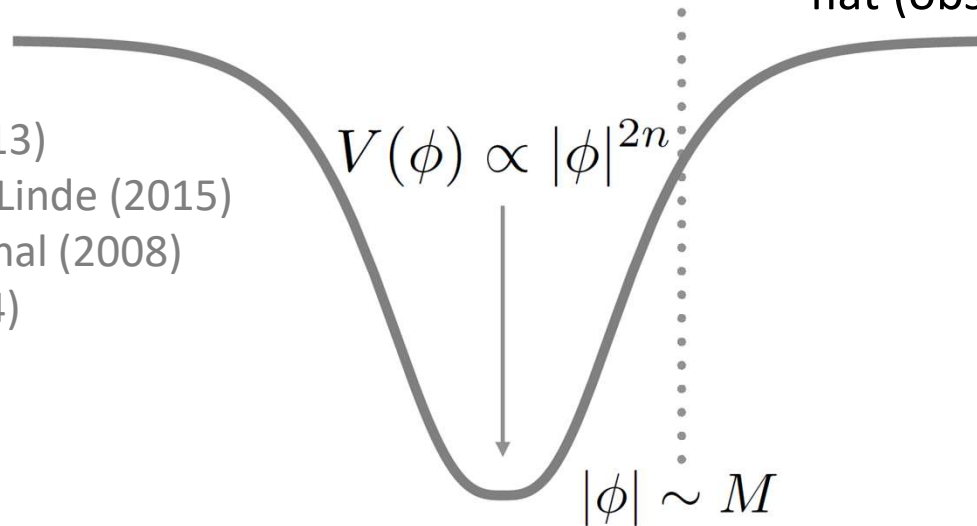
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flat (observations)

Kallosh and Linde (2013)
Carrasco, Kallosh and Linde (2015)
Silverstein and Westphal (2008)
McAllister, et al. (2014)

$$V(\phi) \propto |\phi|^{2n}$$

$$|\phi| \sim M$$

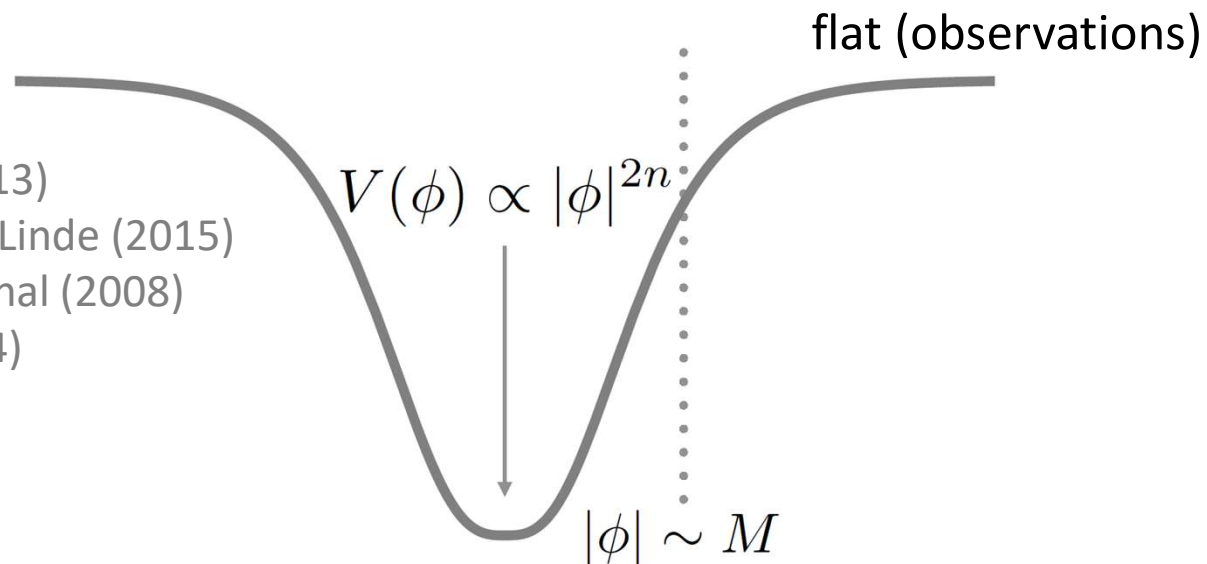


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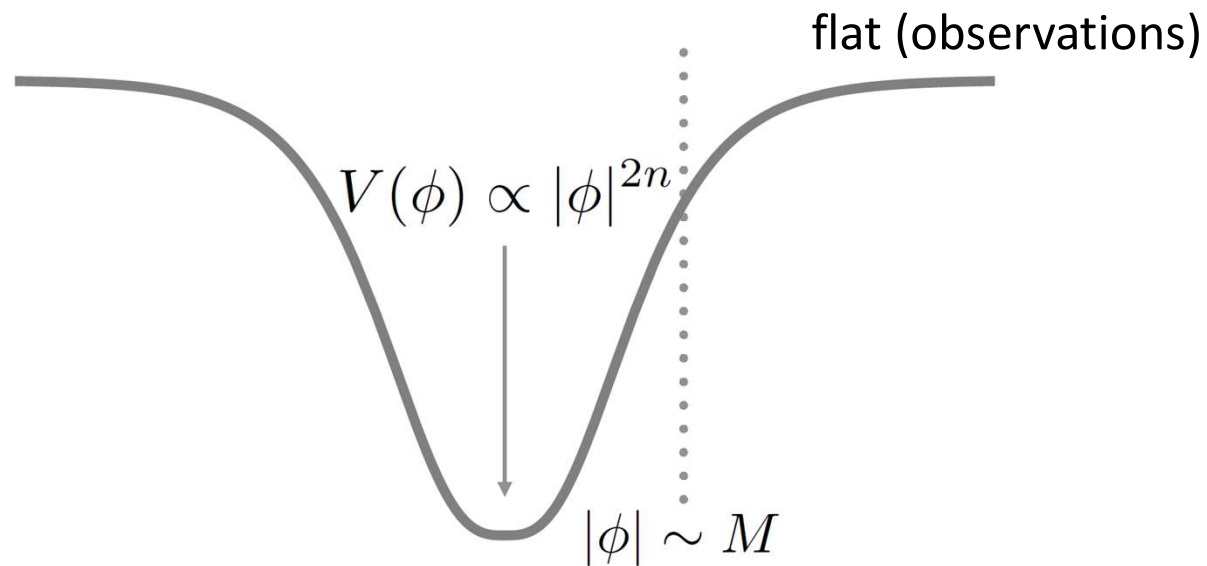
assumption: self-couplings dominate over others

What is w after inflation?

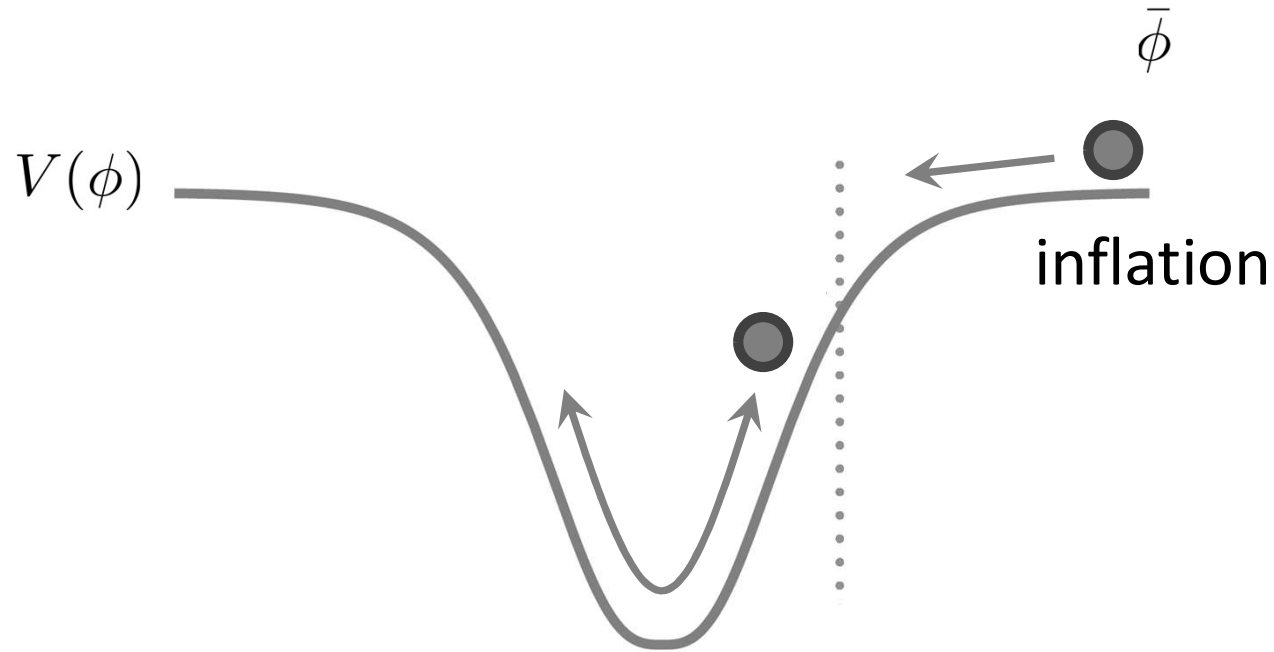
at sufficiently late times:

$$w = \begin{cases} 0 & \text{if } n = 1 \\ 1/3 & \text{if } n > 1 \end{cases}$$

(even without couplings to other fields!)

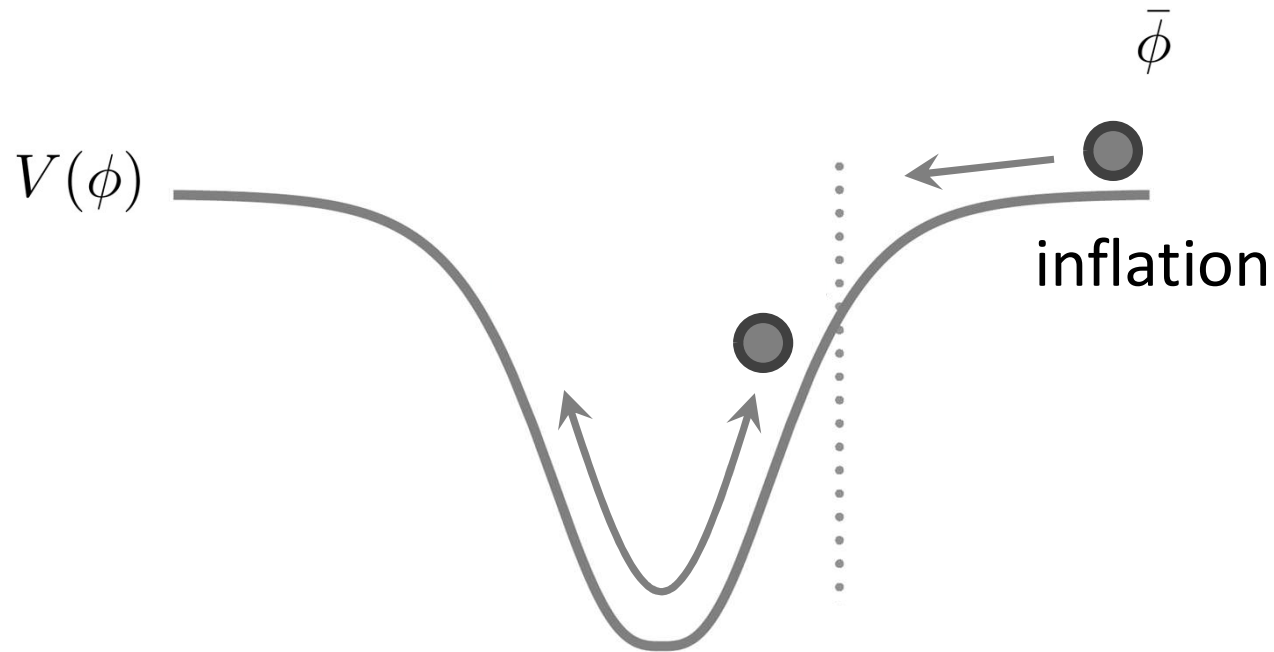


Inflaton dynamics



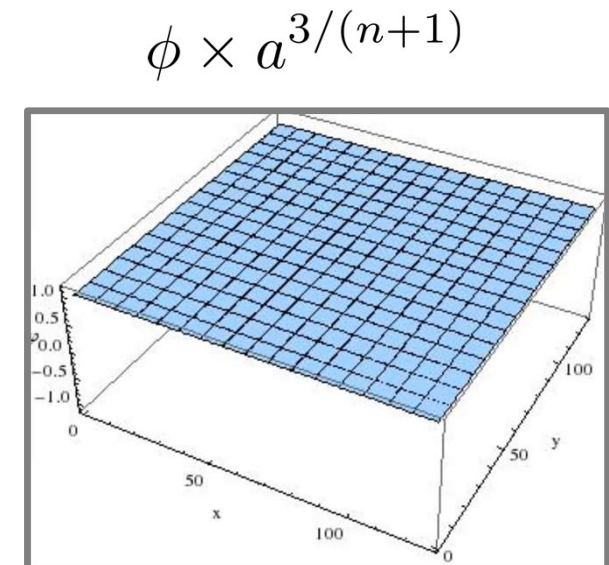
inflation ends:
oscillatory phase

Inflaton (homogeneous) dynamics

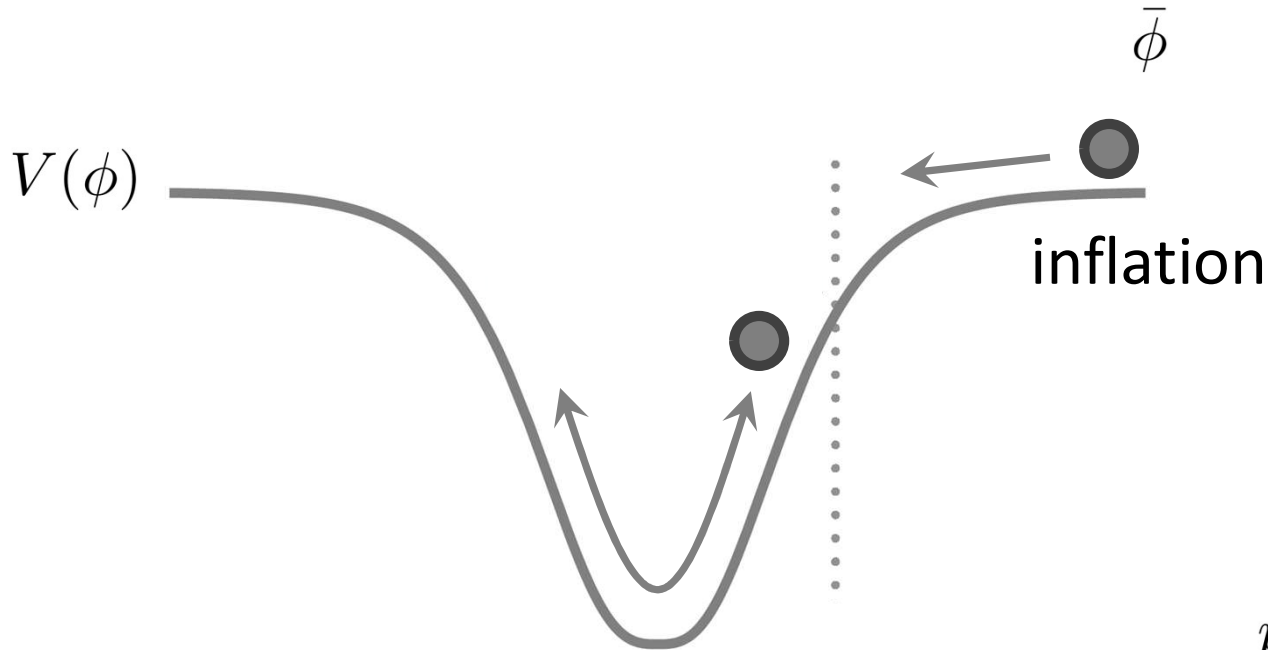


inflation ends:
oscillatory phase

$$\bar{\phi}(t) \propto a(t)^{-3/(n+1)}$$



Inflaton (homogeneous) dynamics

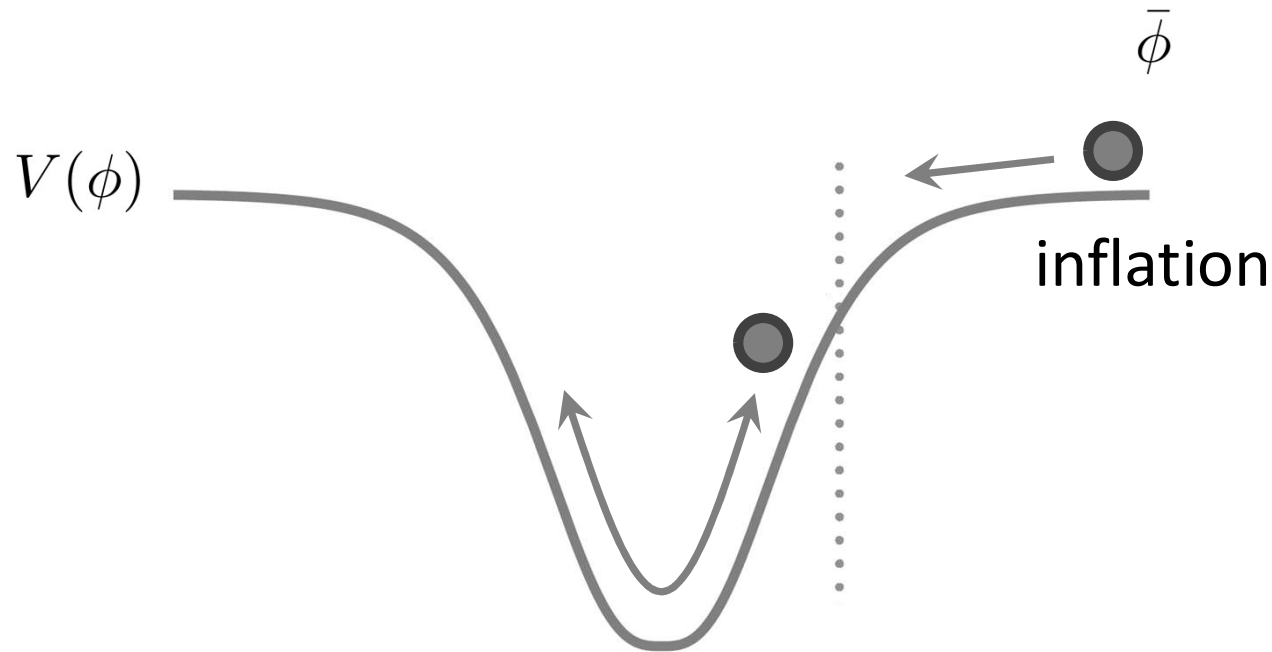


inflation ends:
oscillatory phase

$$\bar{\phi}(t) \propto a(t)^{-3/(n+1)}$$

$$w \equiv \frac{p}{\rho} = \frac{\dot{\phi}^2/2 - (\nabla\phi)^2/6 - V(\phi)}{\dot{\phi}^2/2 + (\nabla\phi)^2/2 + V(\phi)}$$

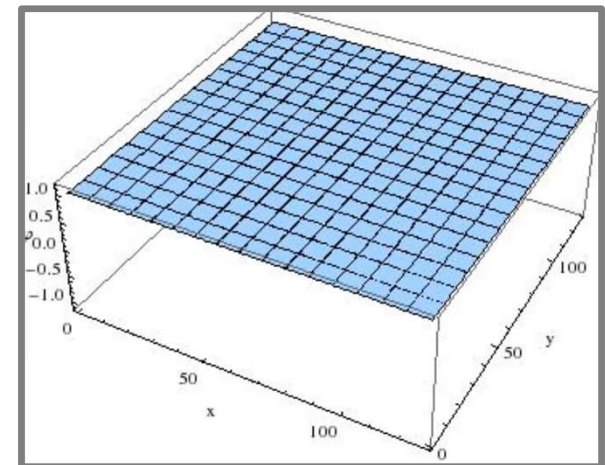
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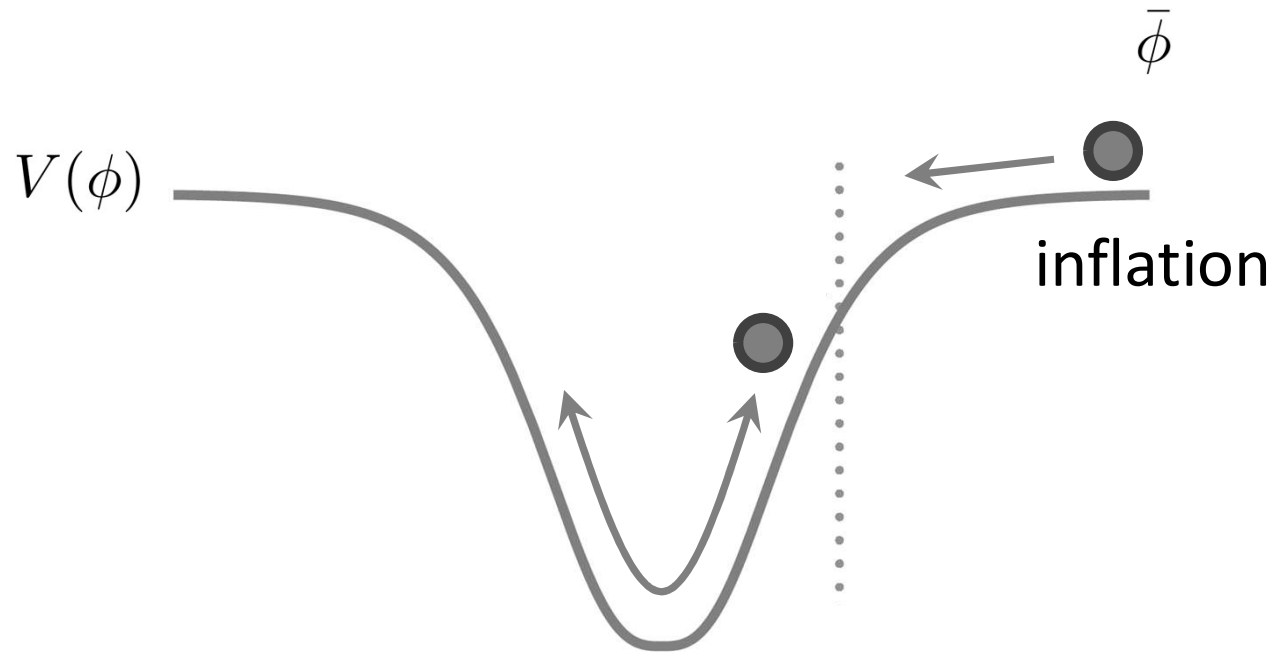
$$w_{\text{hom}} = \frac{n-1}{n+1}$$

$$\phi \times a^{3/(n+1)}$$



Turner (1983)

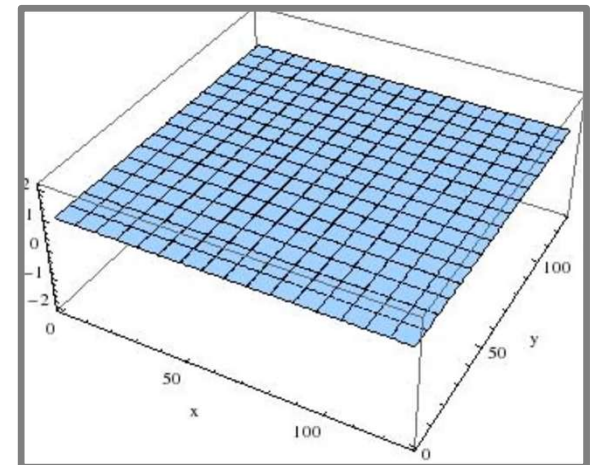
Inflaton (actual) dynamics



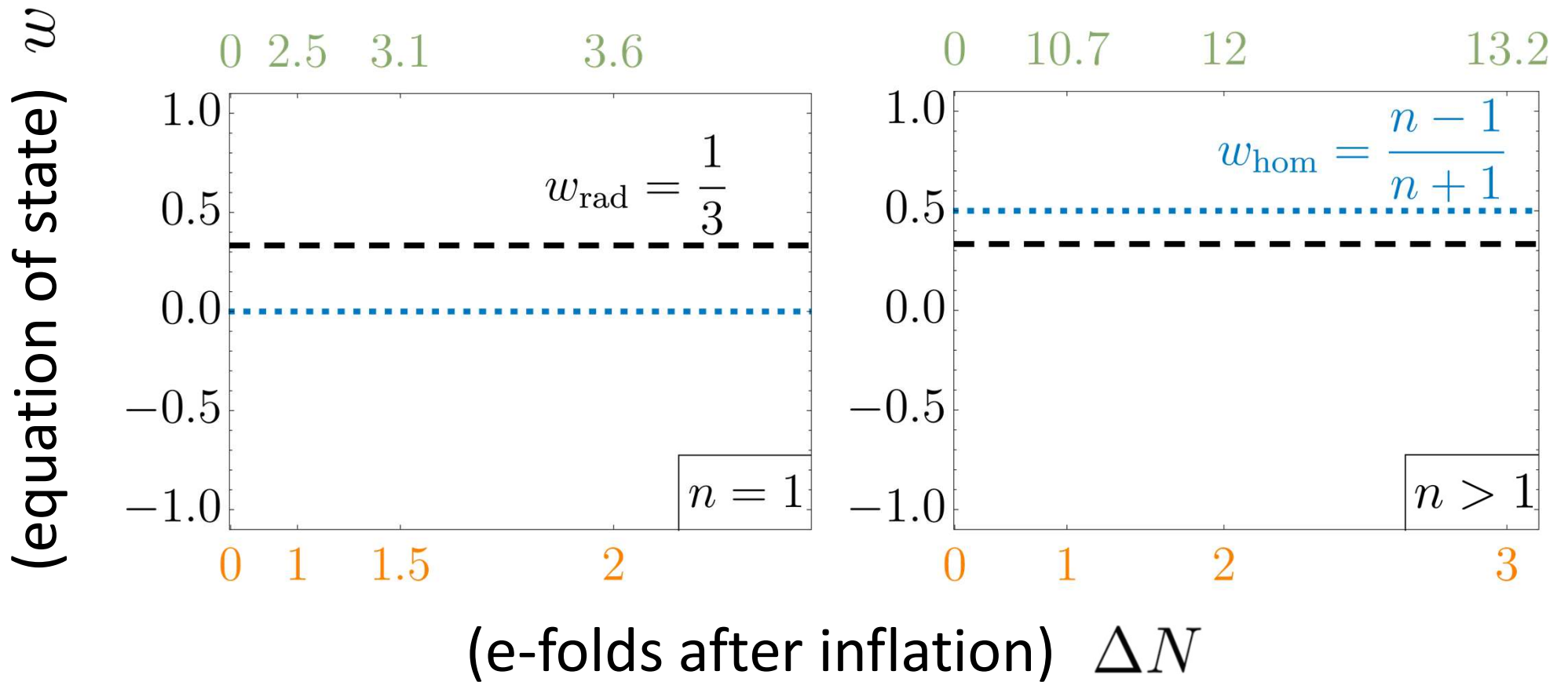
inflation ends:
oscillatory phase

- parametric resonance of $\delta\phi(t, \mathbf{x})$
- $\bar{\phi}$ fragments

$$\phi \times a^{3/(n+1)}$$



Equation of state

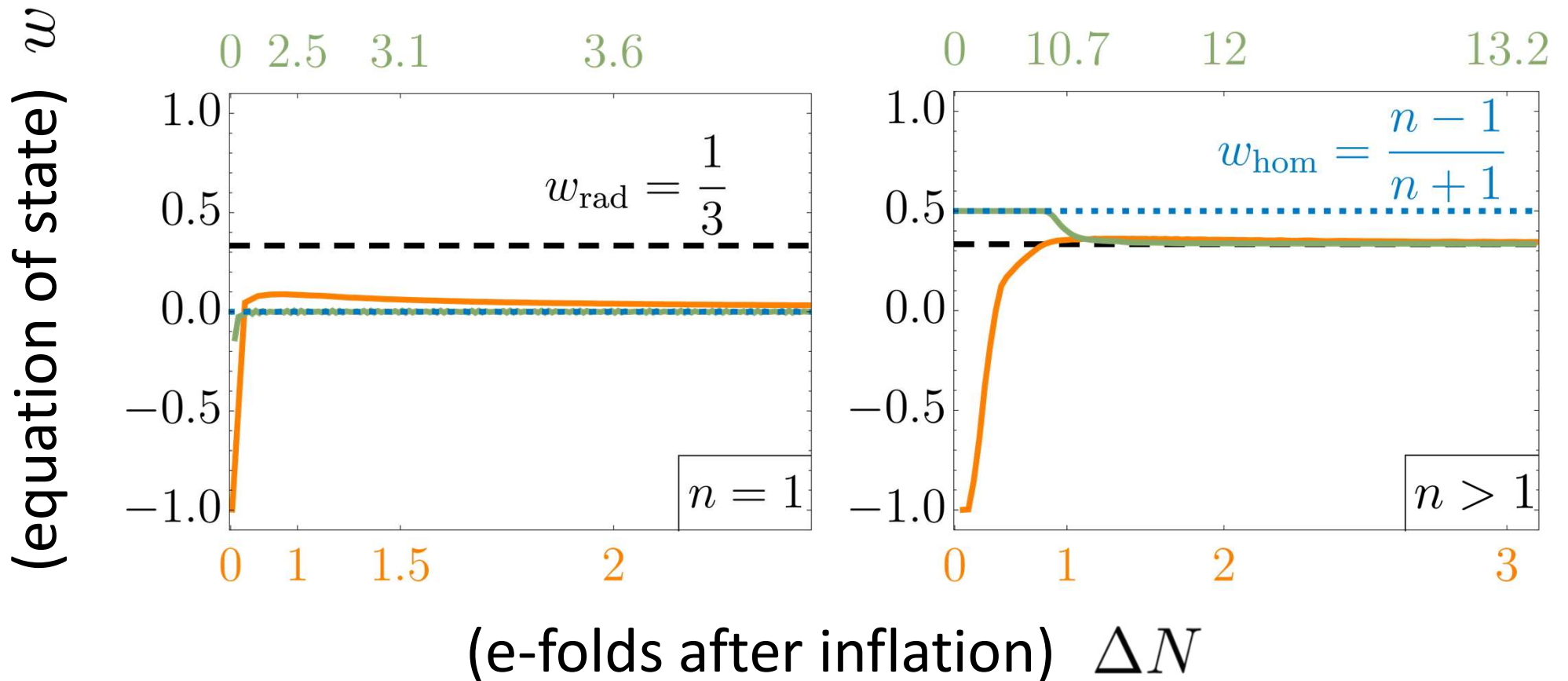


$$\Delta N \equiv \int_{a_{\text{end}}}^a d \ln a$$

Equation of state

— $M \ll m_{\text{pl}}$ (efficient resonance)

— $M \sim m_{\text{pl}}$ (inefficient resonance)



$$\Delta N \equiv \int_{a_{\text{end}}}^a d \ln a$$

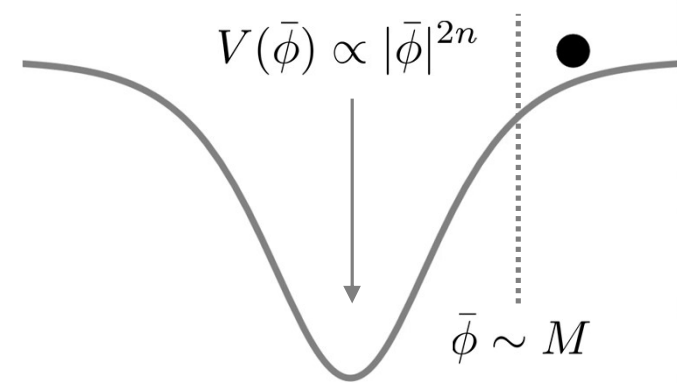
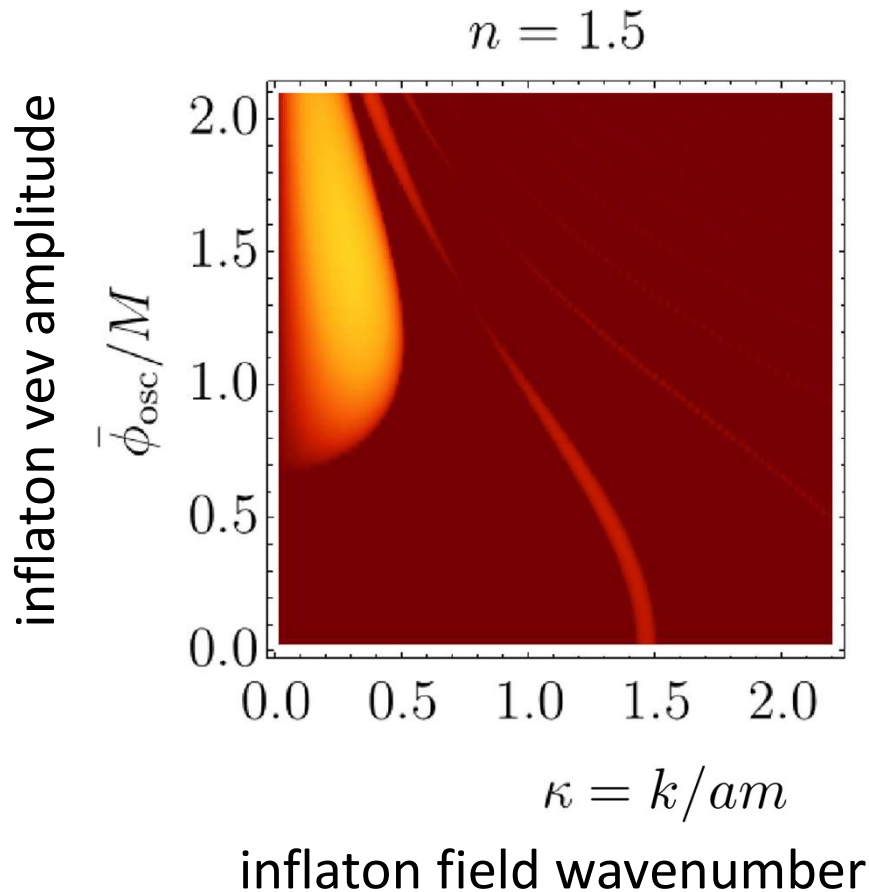
Towards radiation domination

$$n > 1$$

Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)



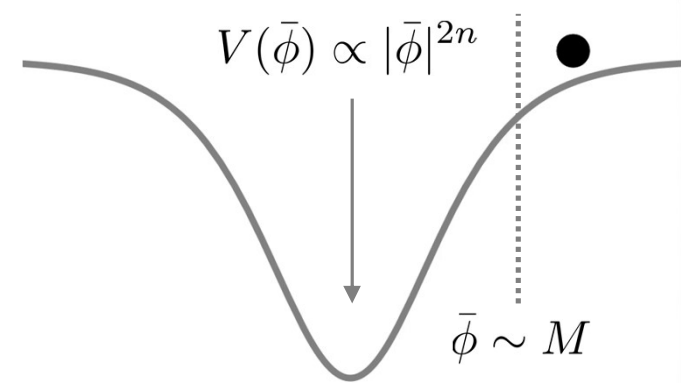
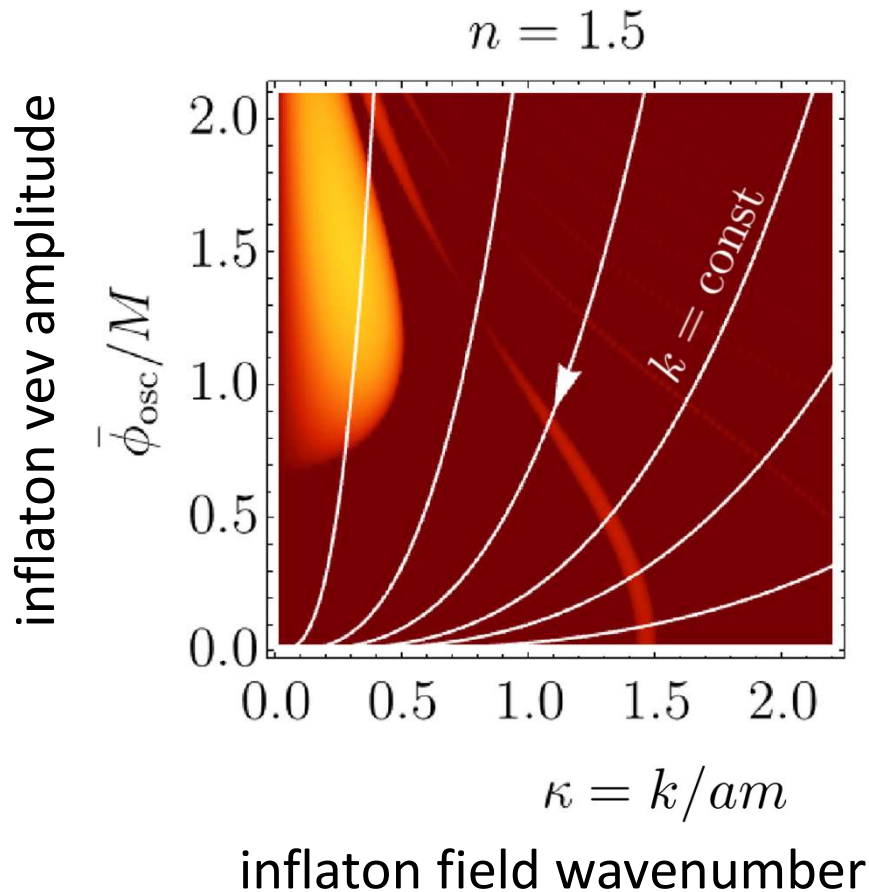
$$\delta\phi_k \propto \exp(\pm\mu_k t)$$

$$m^2 \equiv V'(\bar{\phi}_{\text{osc}})/\bar{\phi}_{\text{osc}}$$

Towards radiation domination

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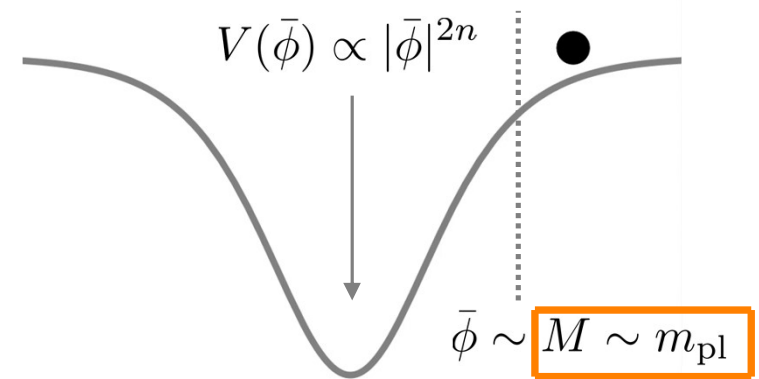
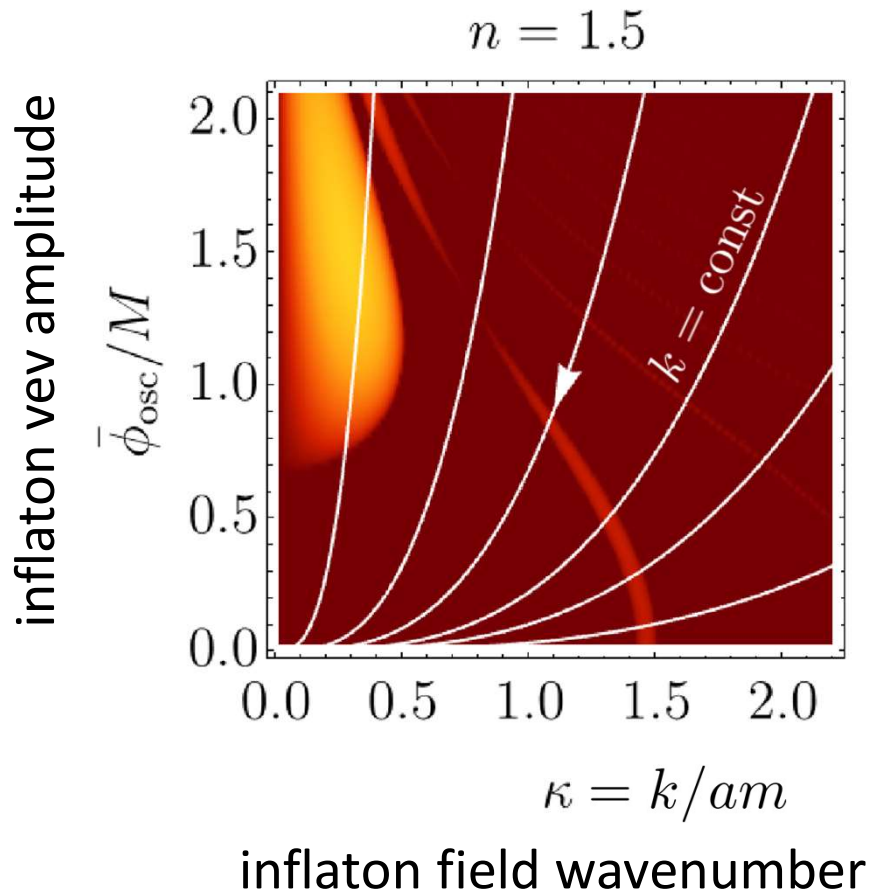
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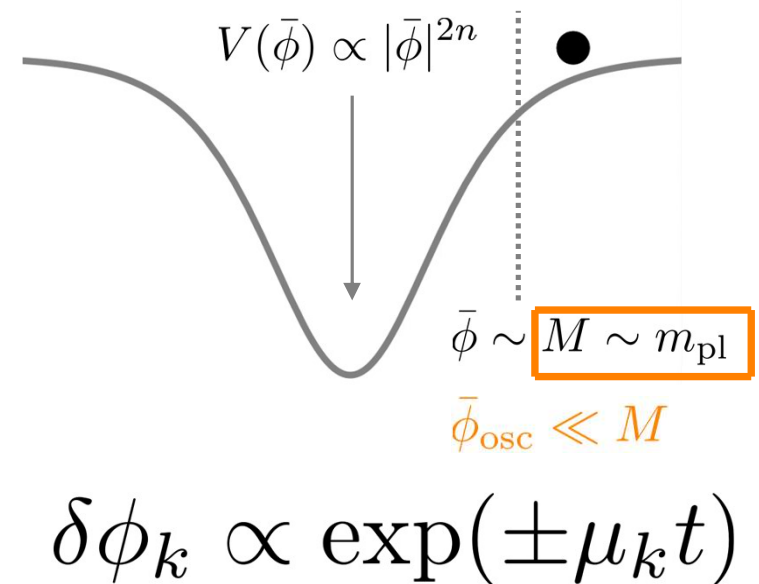
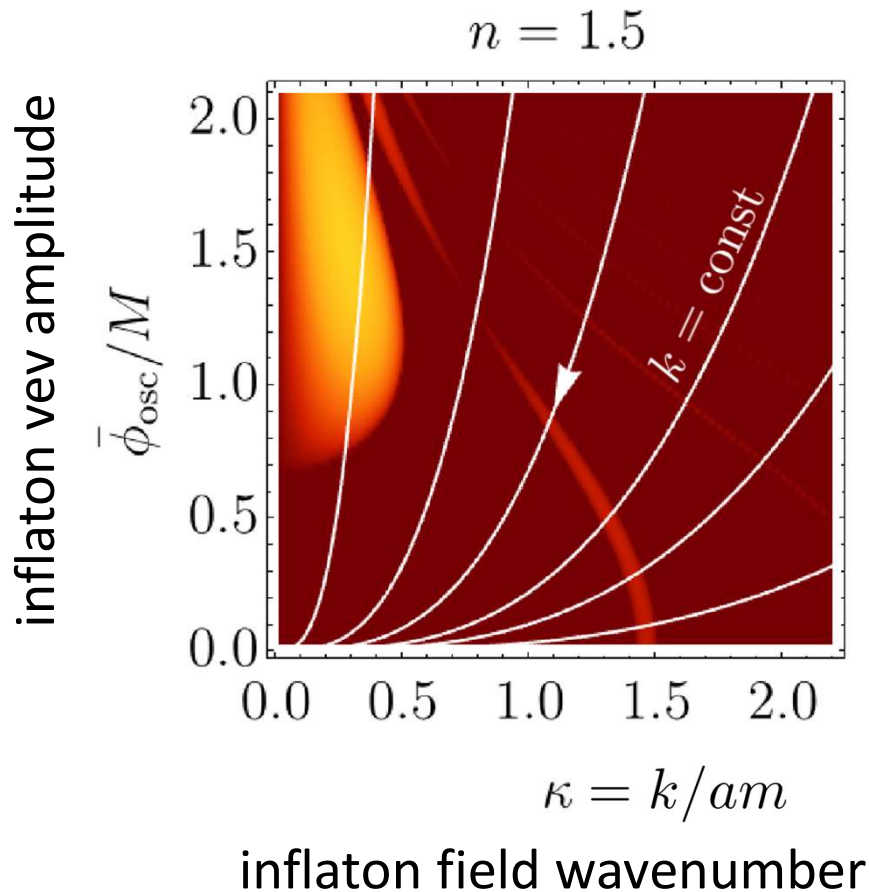
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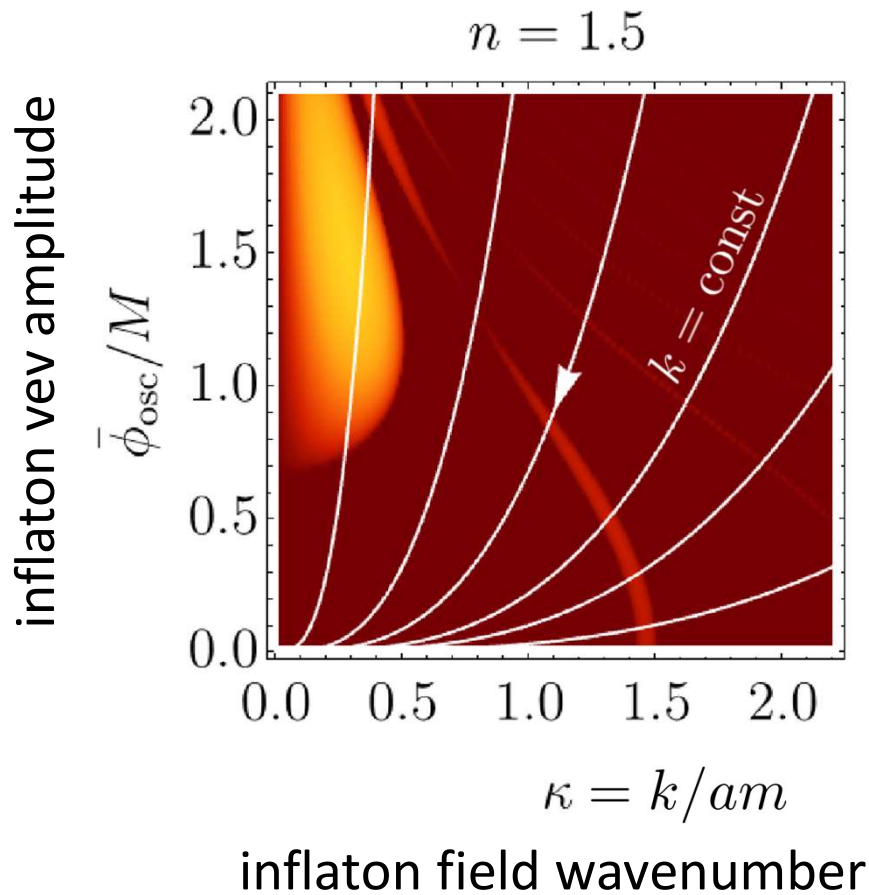


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Towards radiation domination

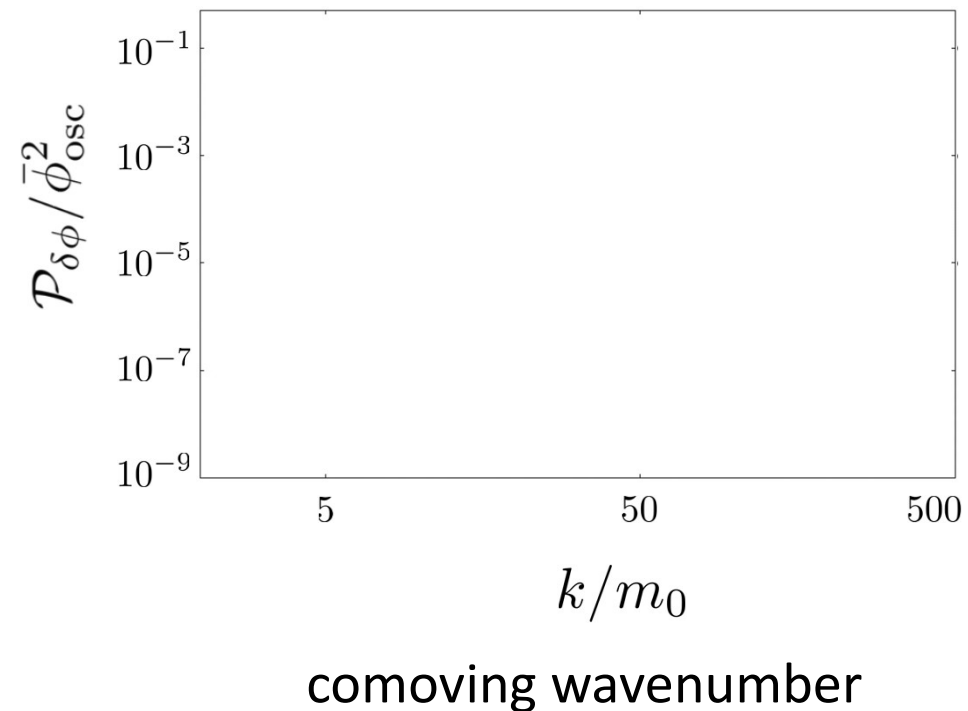
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Non-perturbative decay (parametric self-resonance)



Power spectrum:

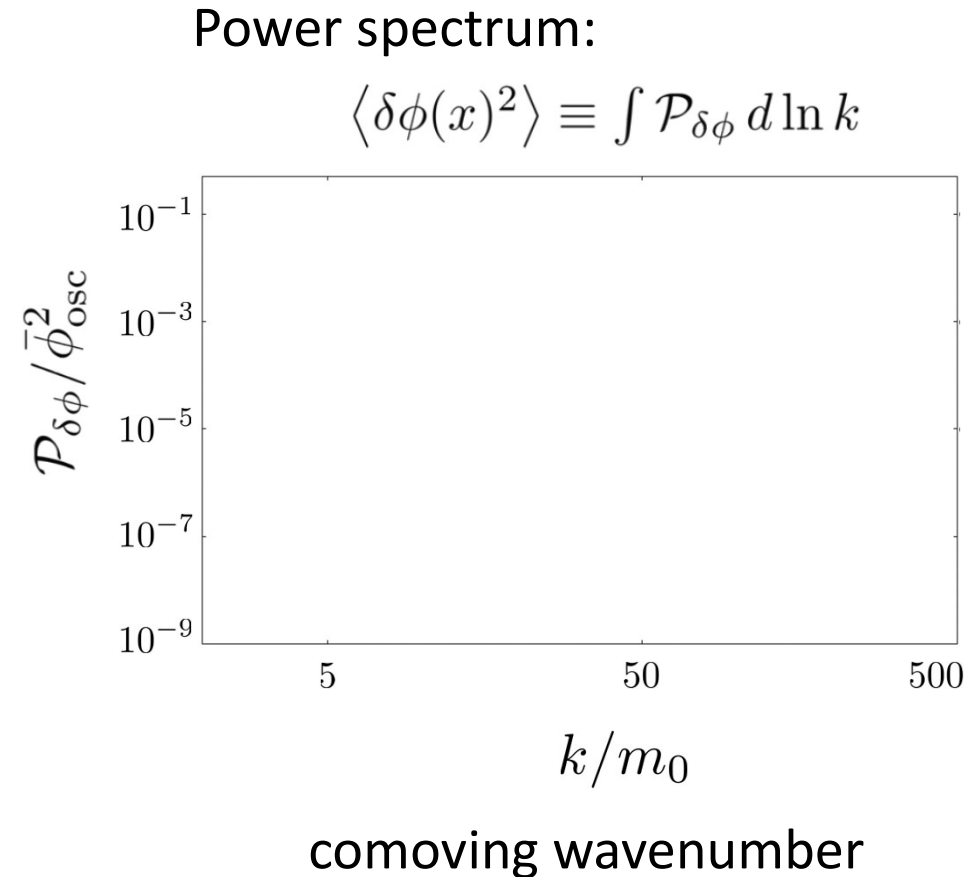
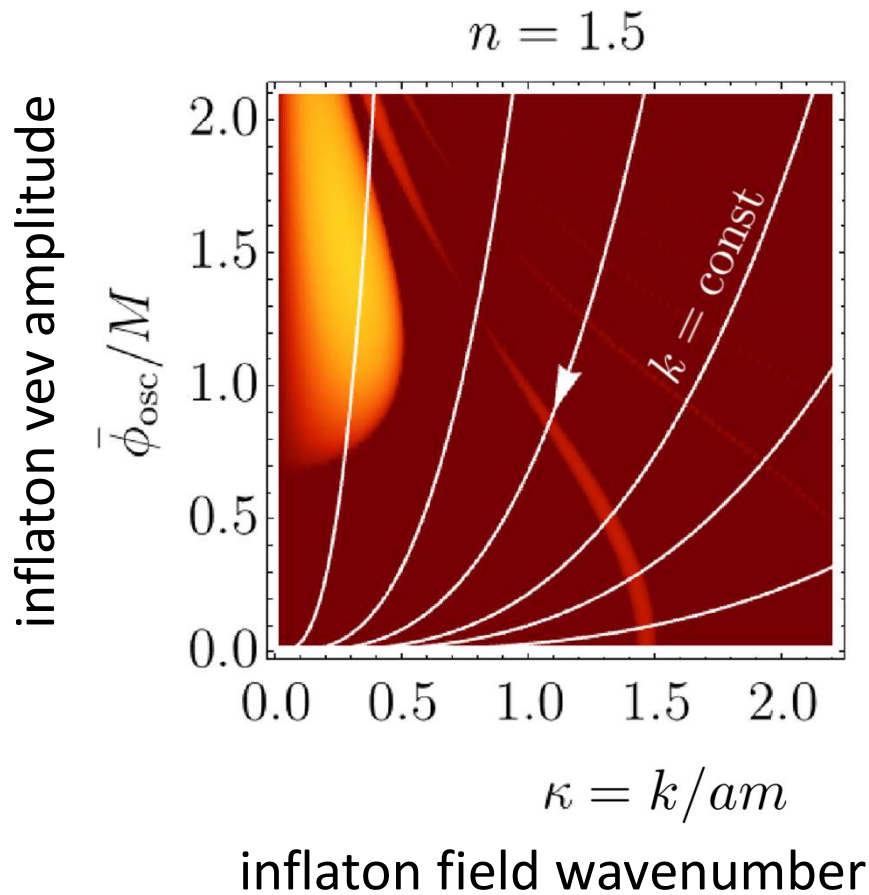
$$\langle \delta\phi(x)^2 \rangle \equiv \int \mathcal{P}_{\delta\phi} d \ln k$$



Towards radiation domination

$$n > 1$$

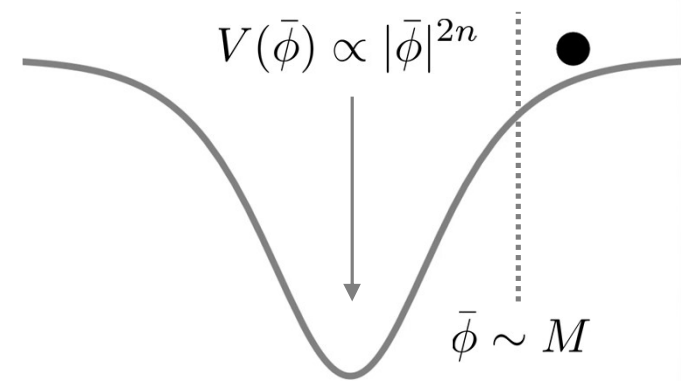
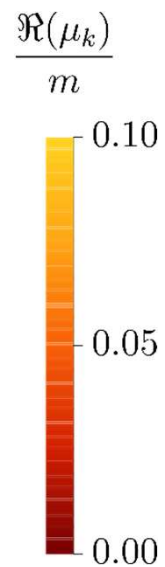
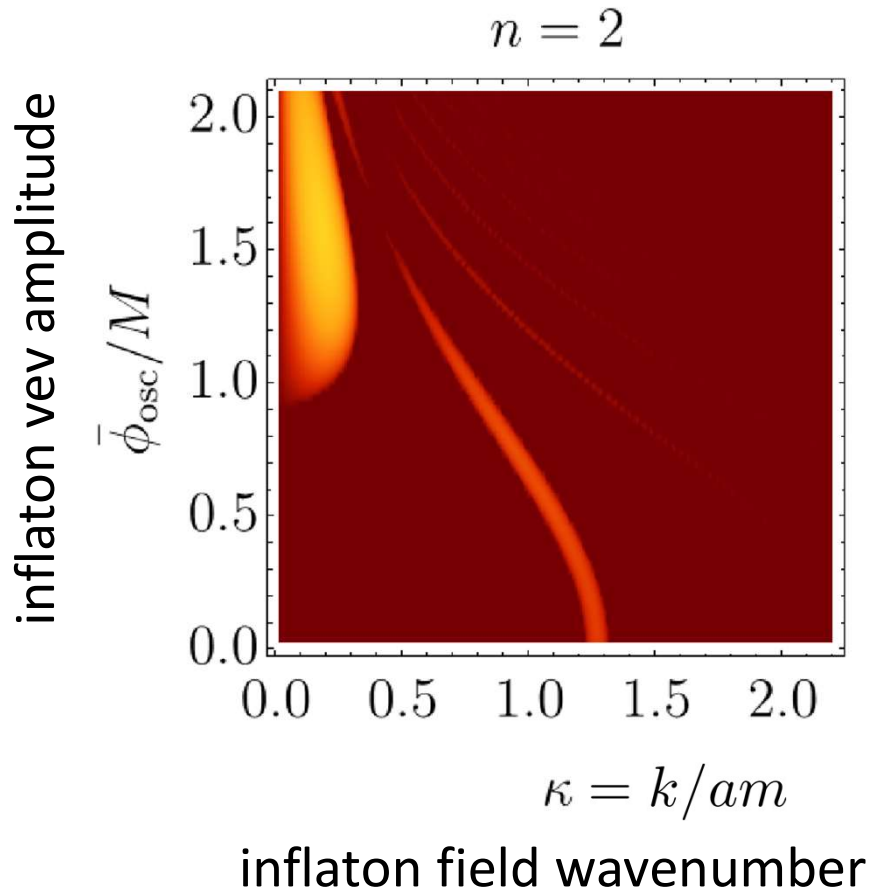
Non-perturbative decay (parametric self-resonance)



Towards radiation domination

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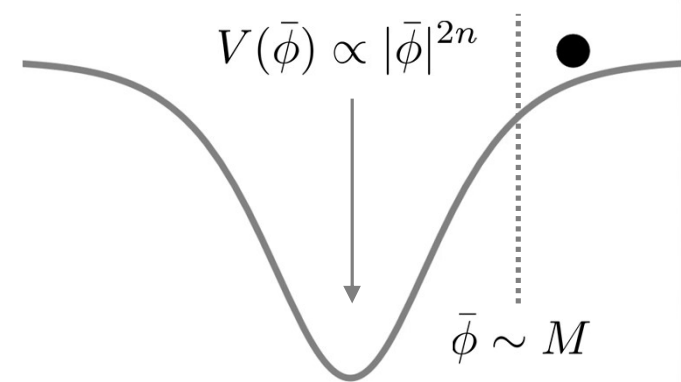
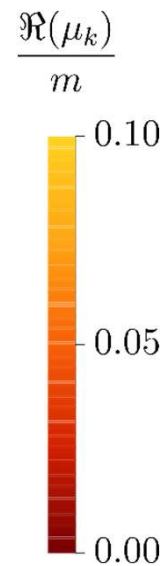
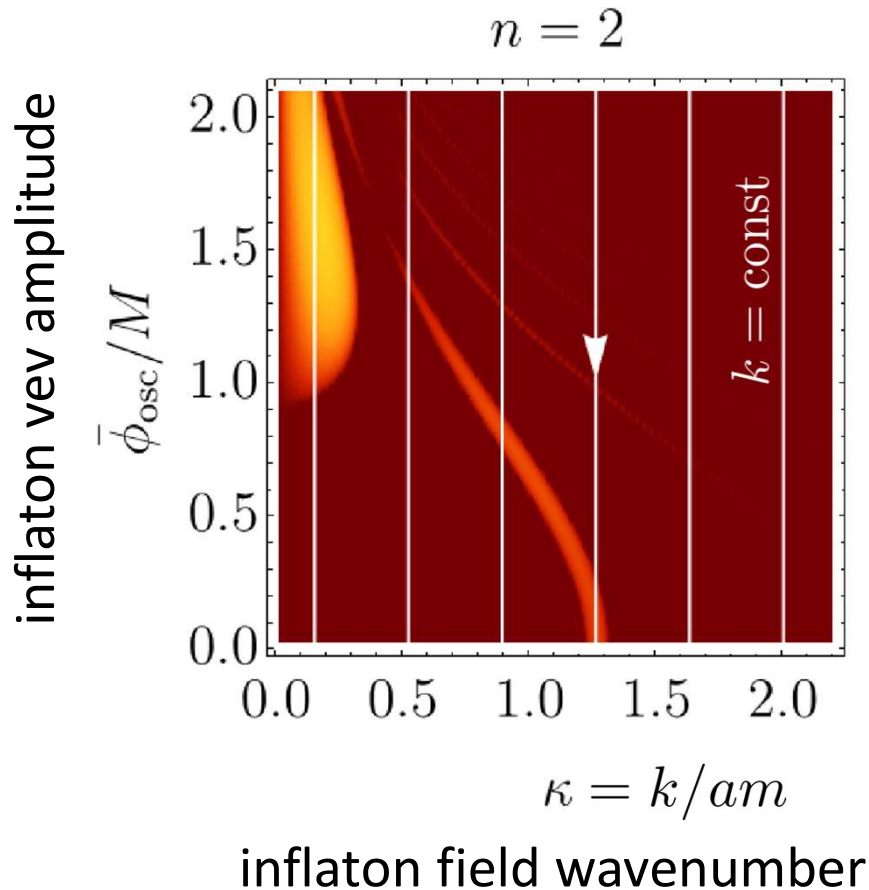
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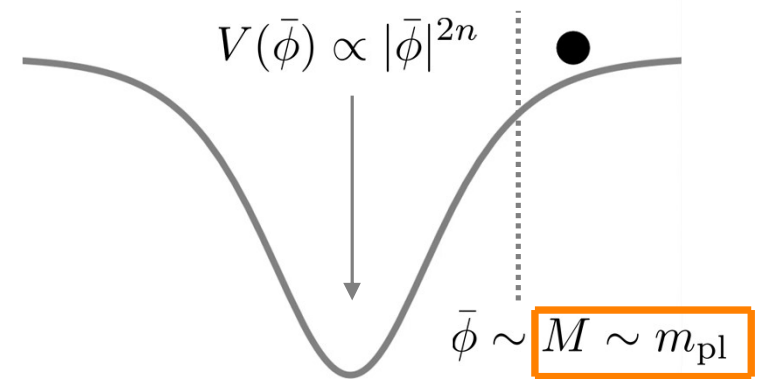
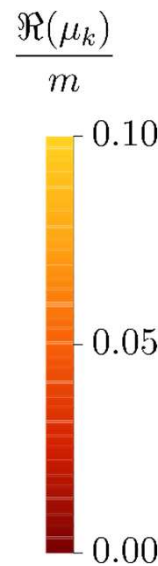
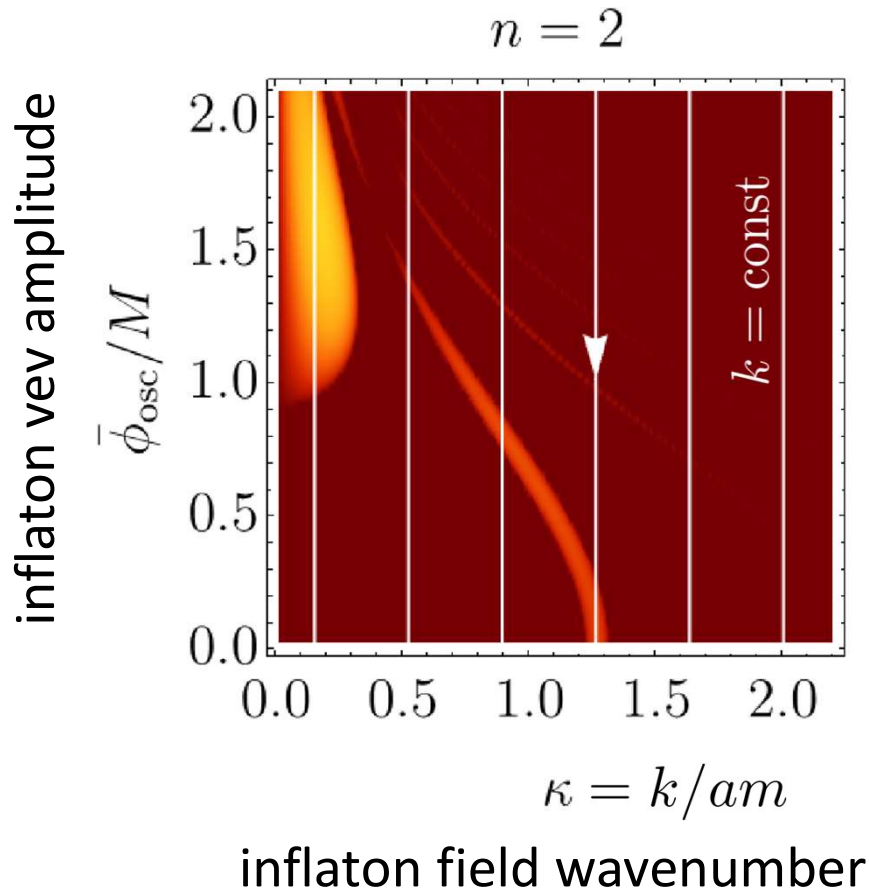
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Non-perturbative decay (parametric self-resonance)



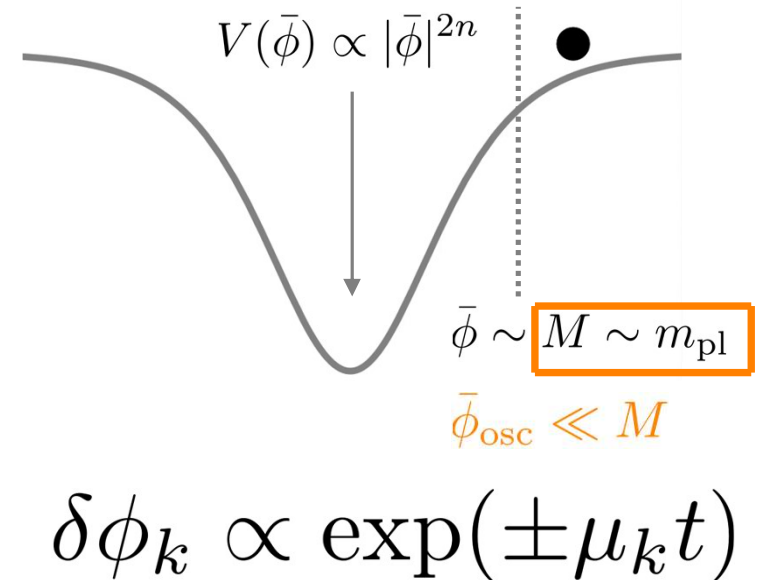
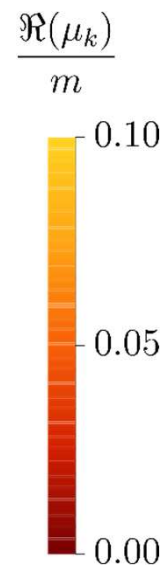
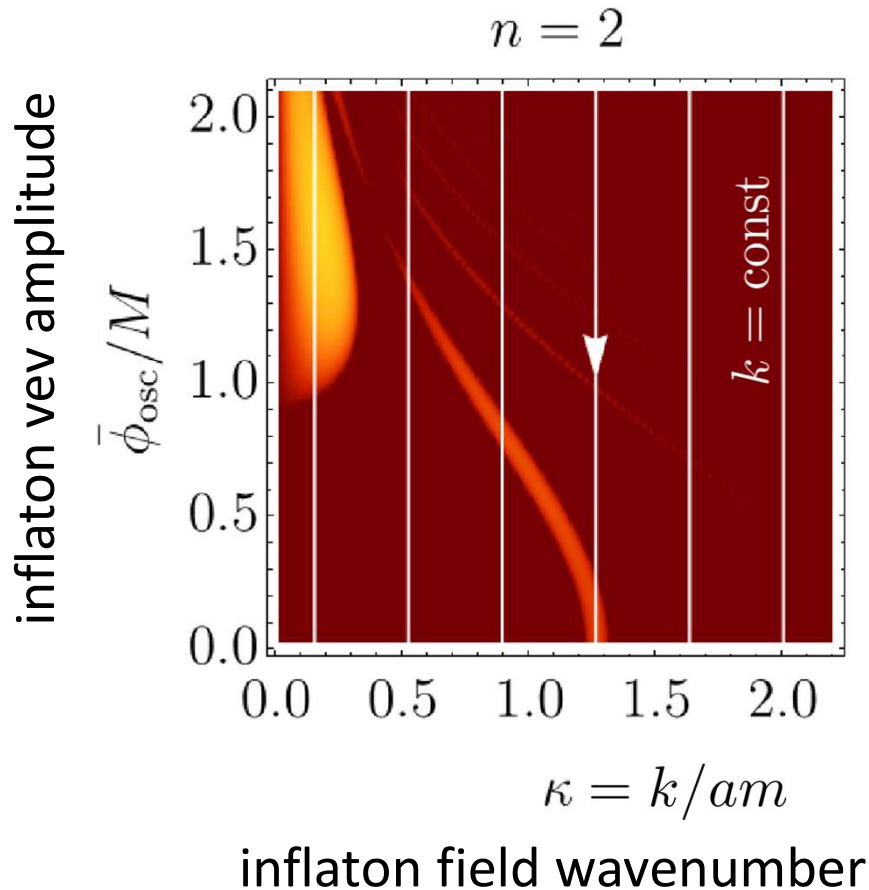
$$\delta\phi_k \propto \exp(\pm\mu_k t)$$

$$m^2 \equiv V'(\bar{\phi}_{\text{osc}})/\bar{\phi}_{\text{osc}}$$

Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)

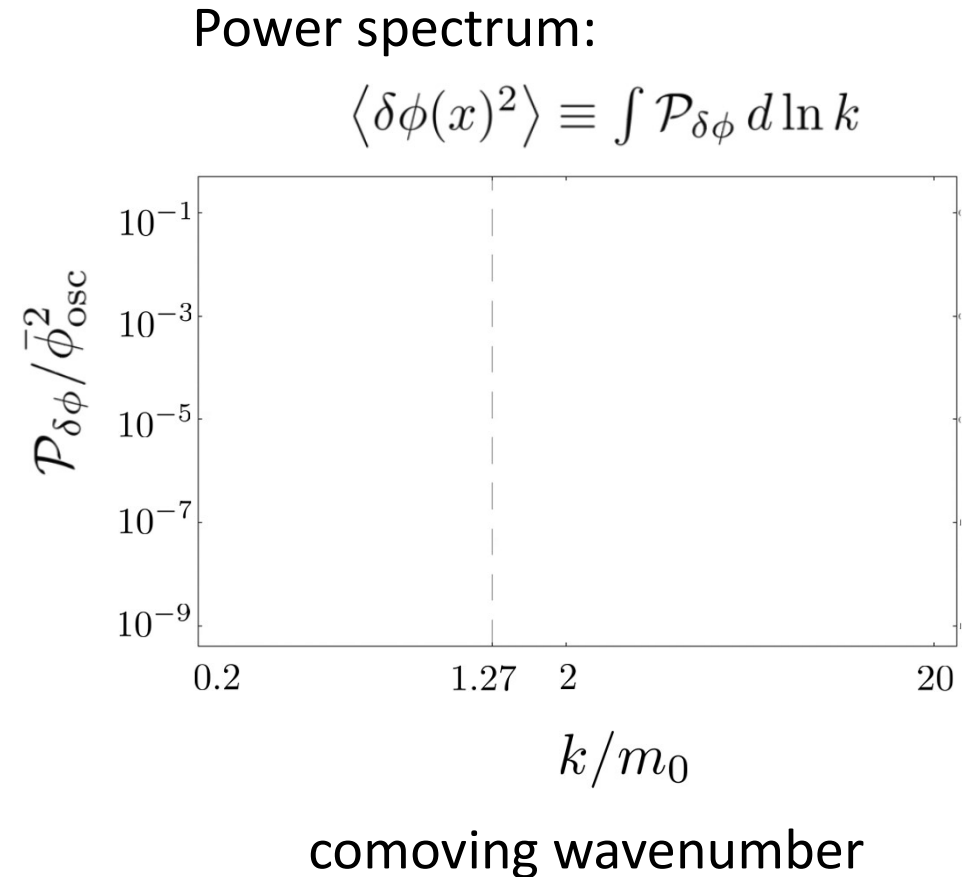
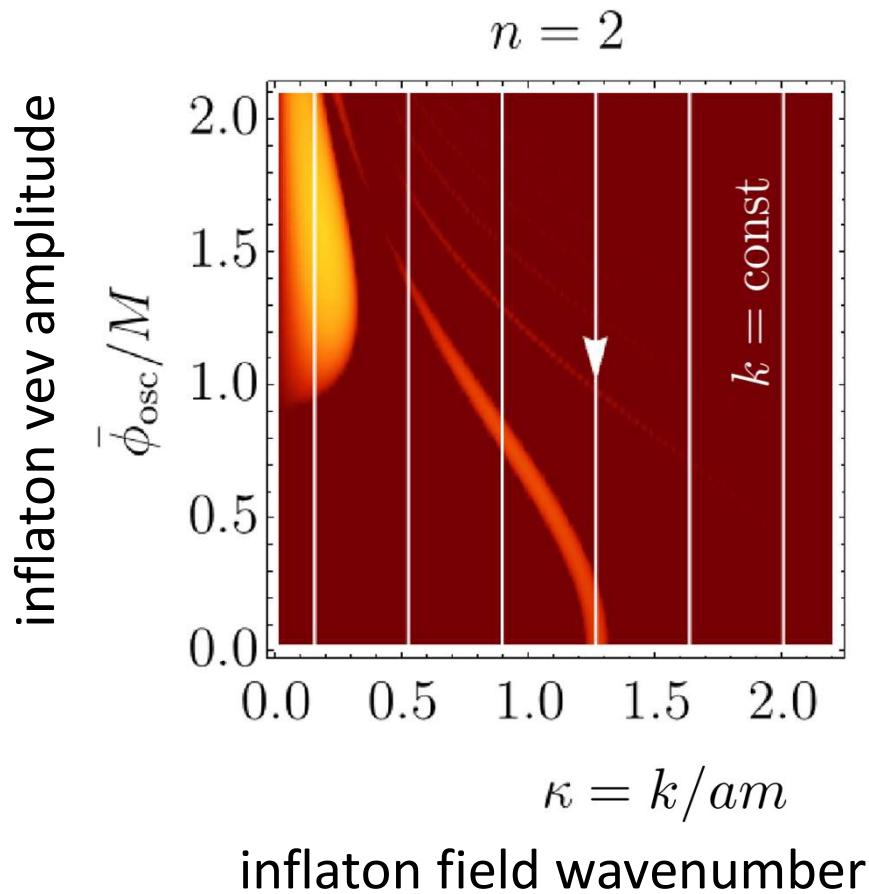


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Towards radiation domination

$$n > 1$$

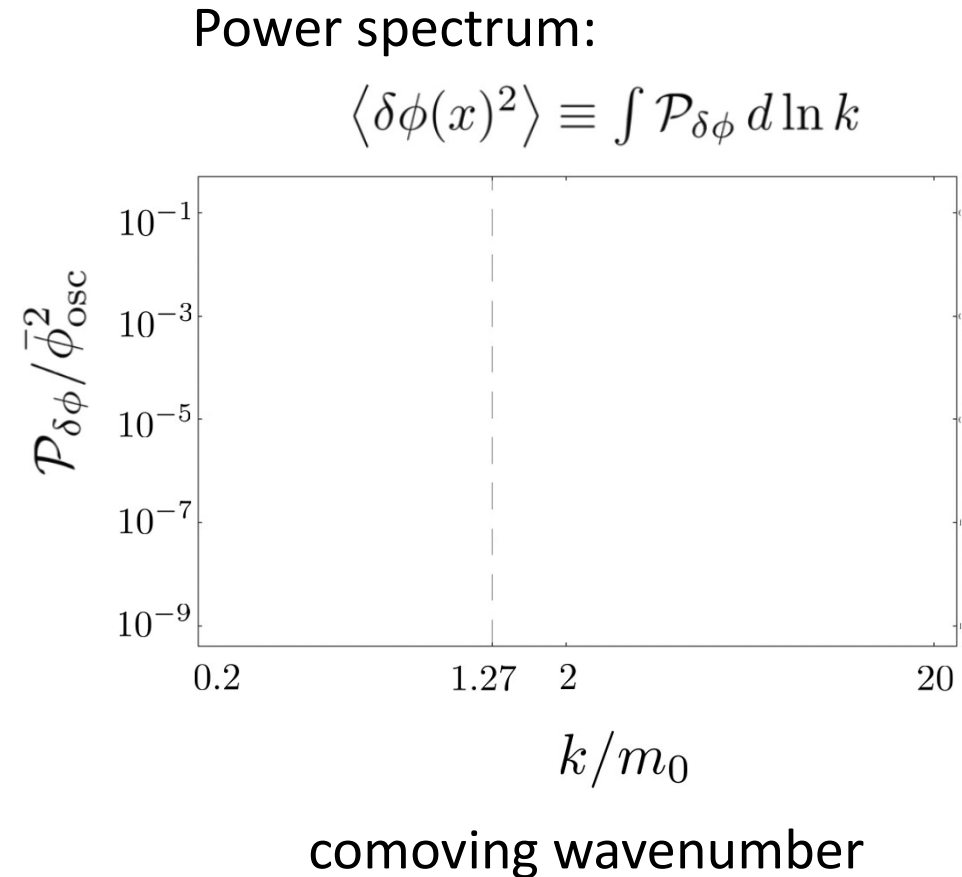
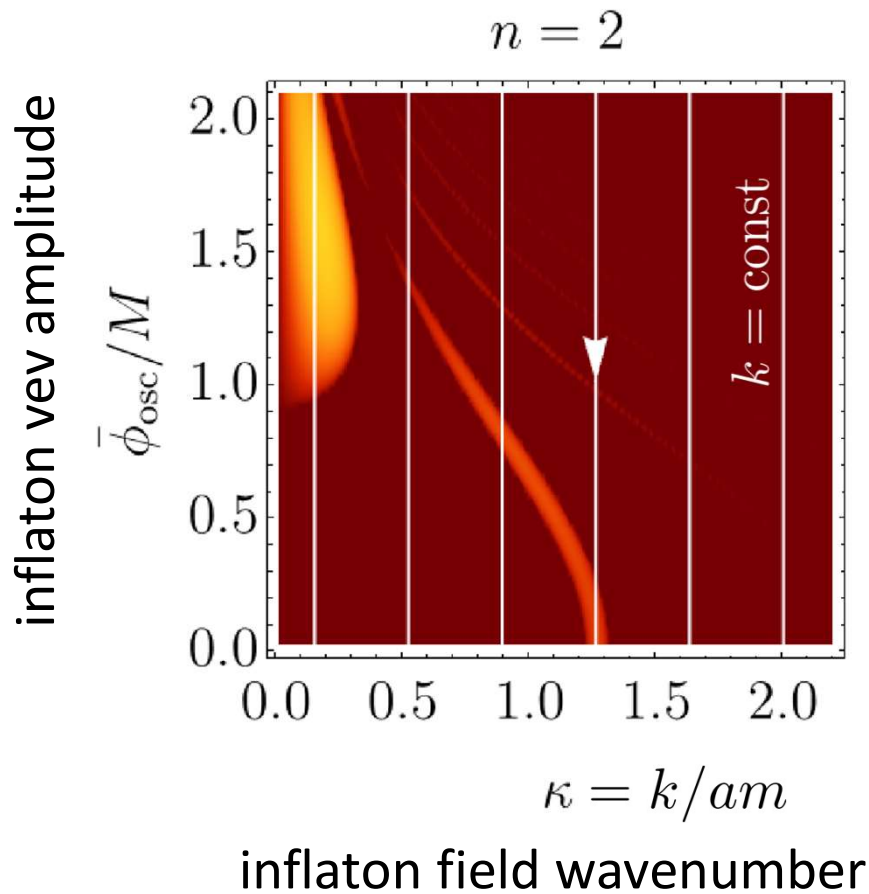
Non-perturbative decay (parametric self-resonance)



Towards radiation domination

$$n > 1$$

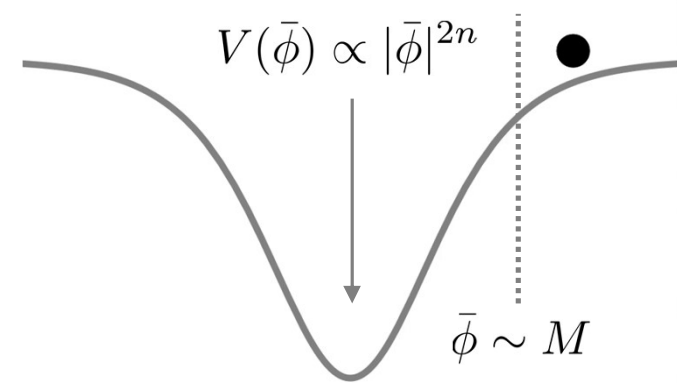
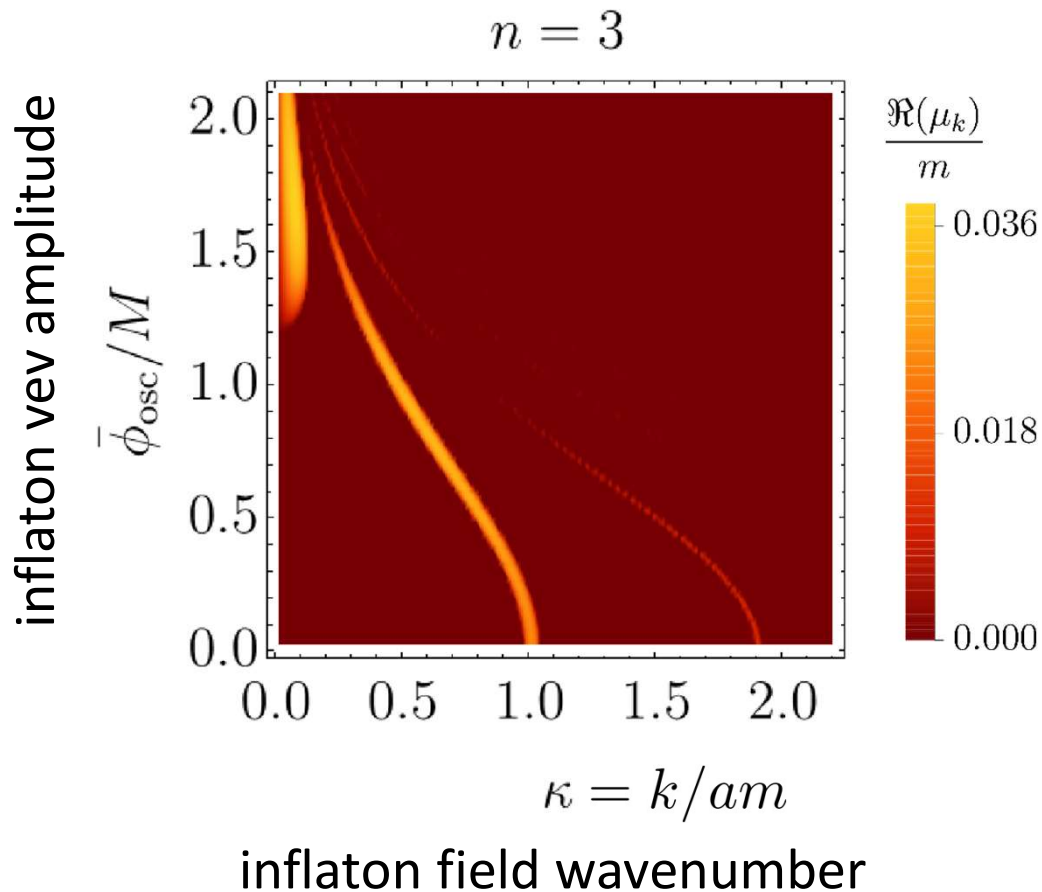
Non-perturbative decay (parametric self-resonance)



Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)



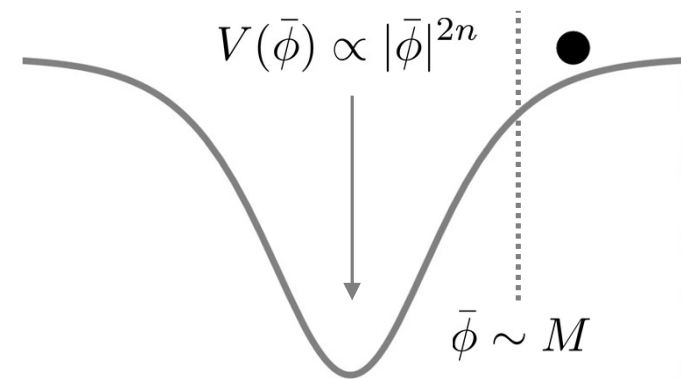
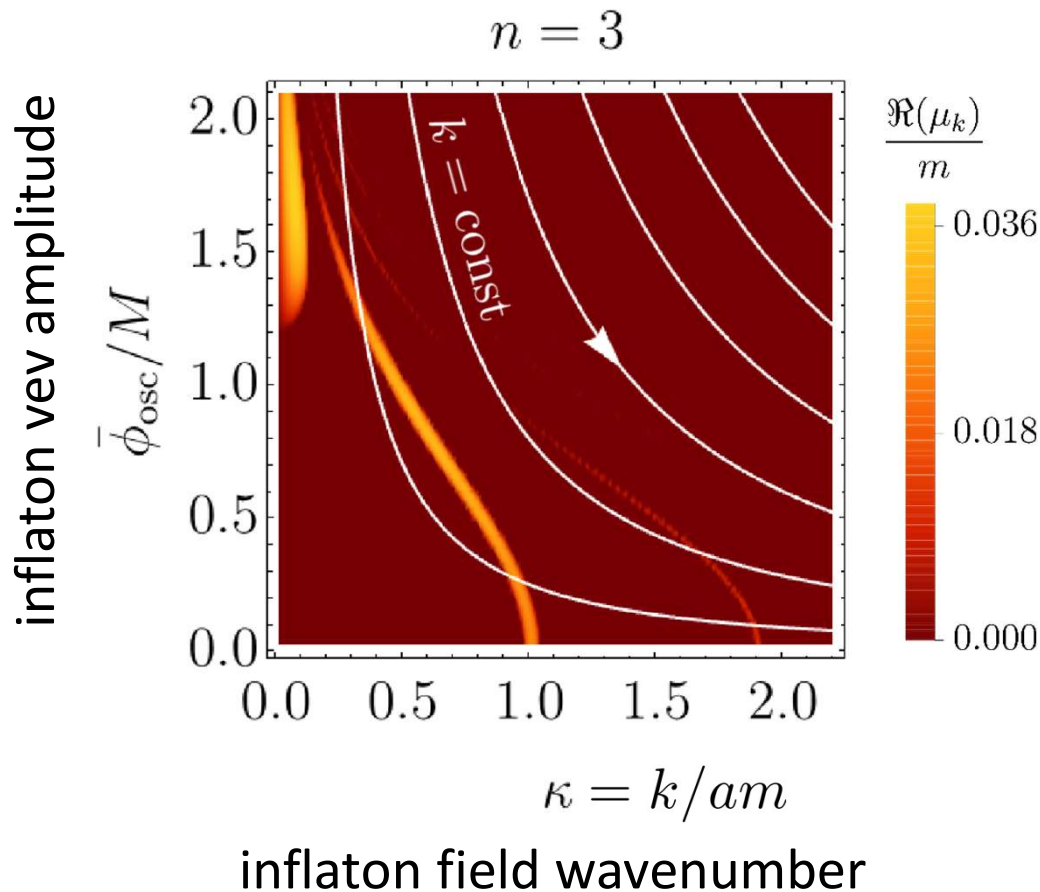
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Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)



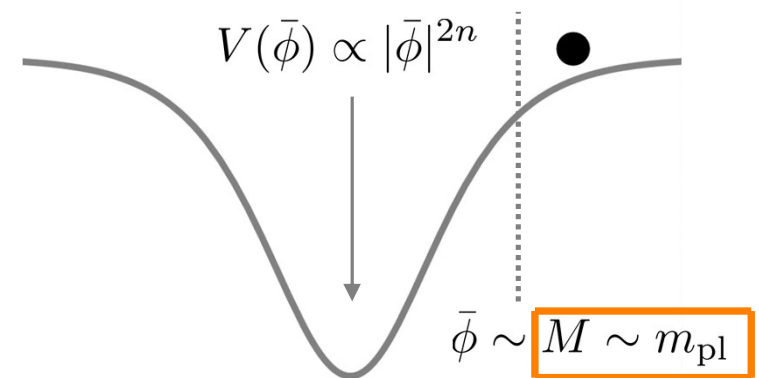
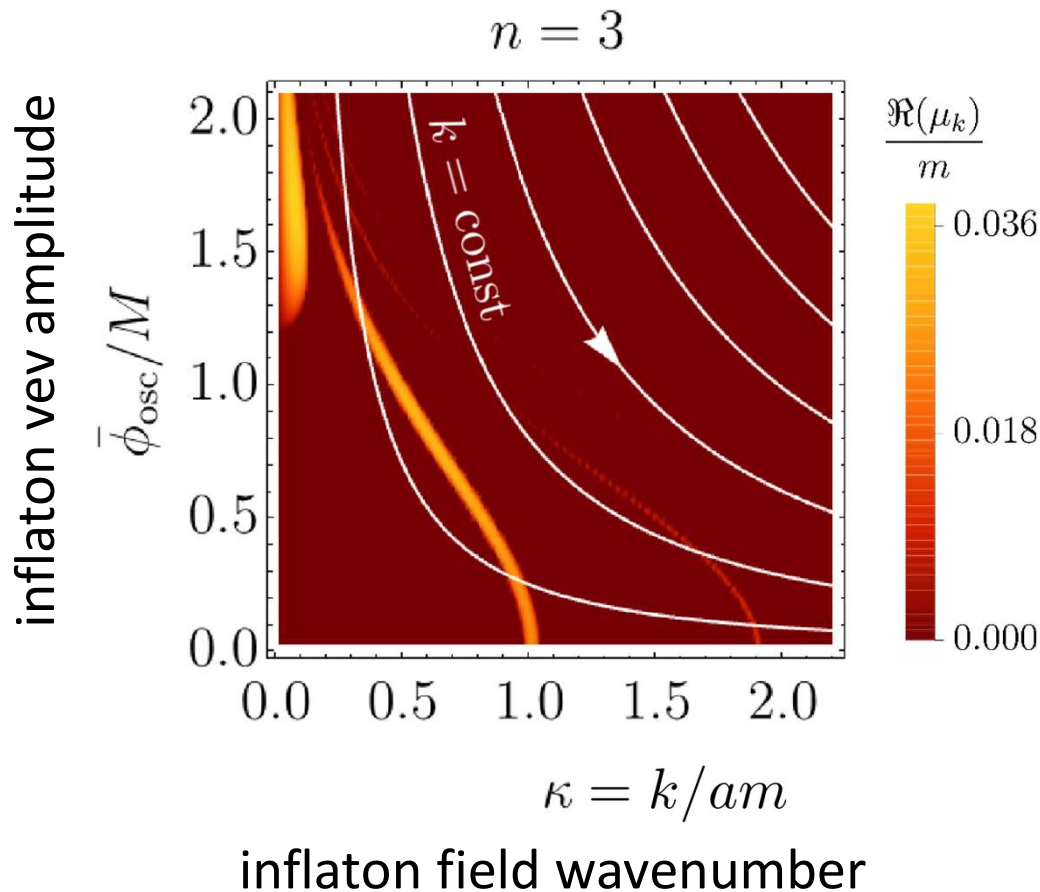
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Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)



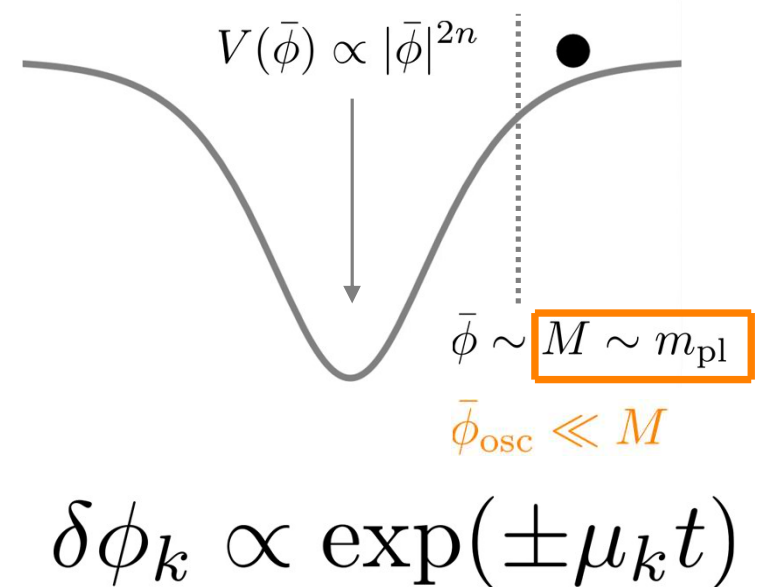
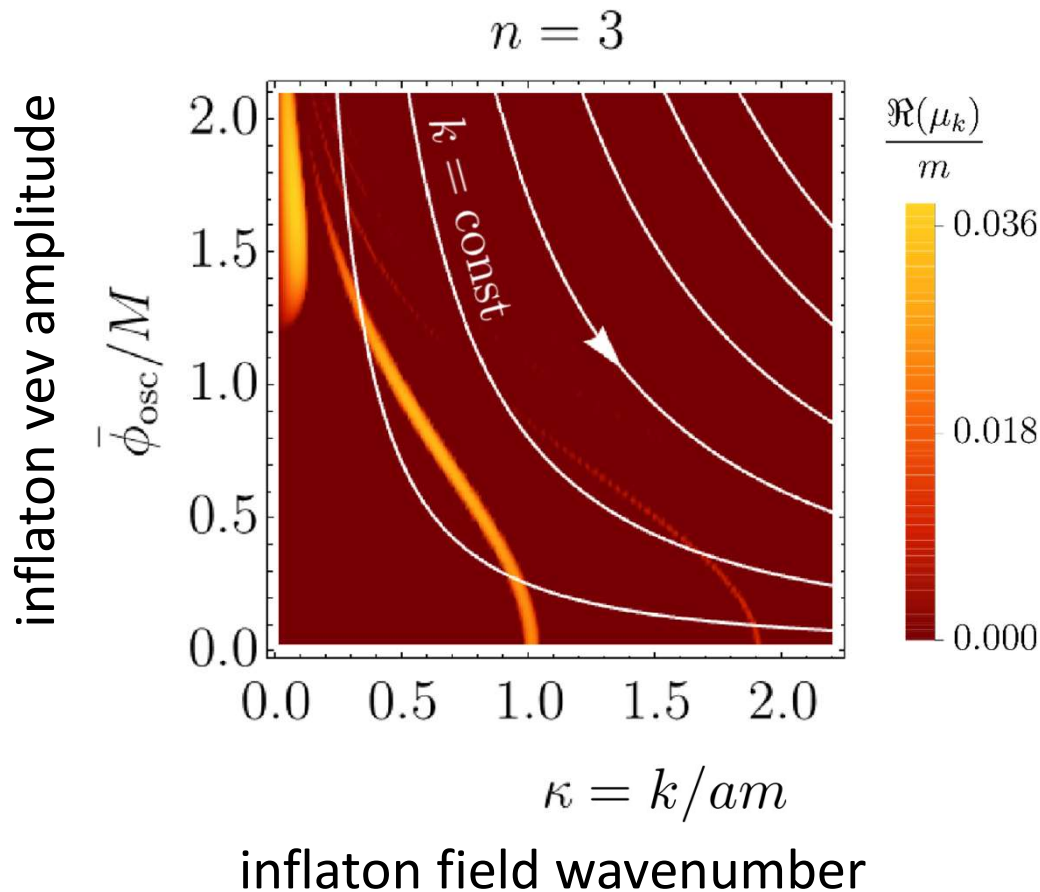
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Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)

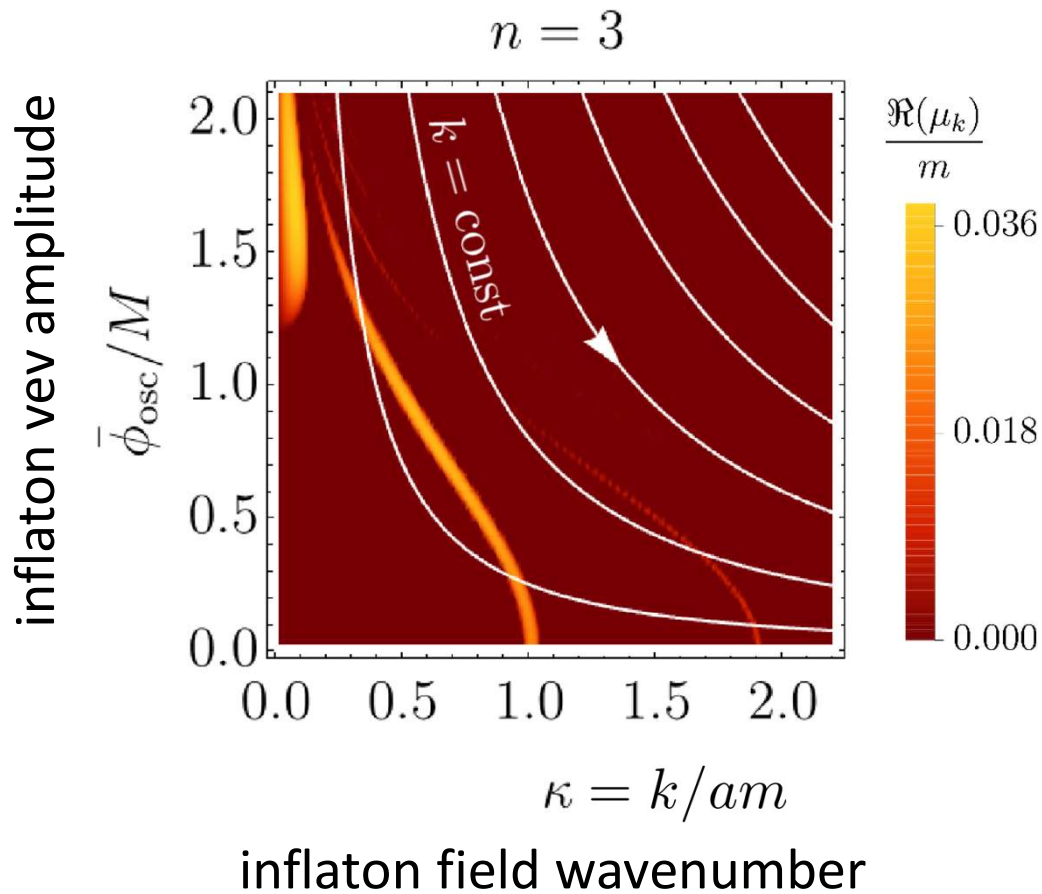


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Towards radiation domination

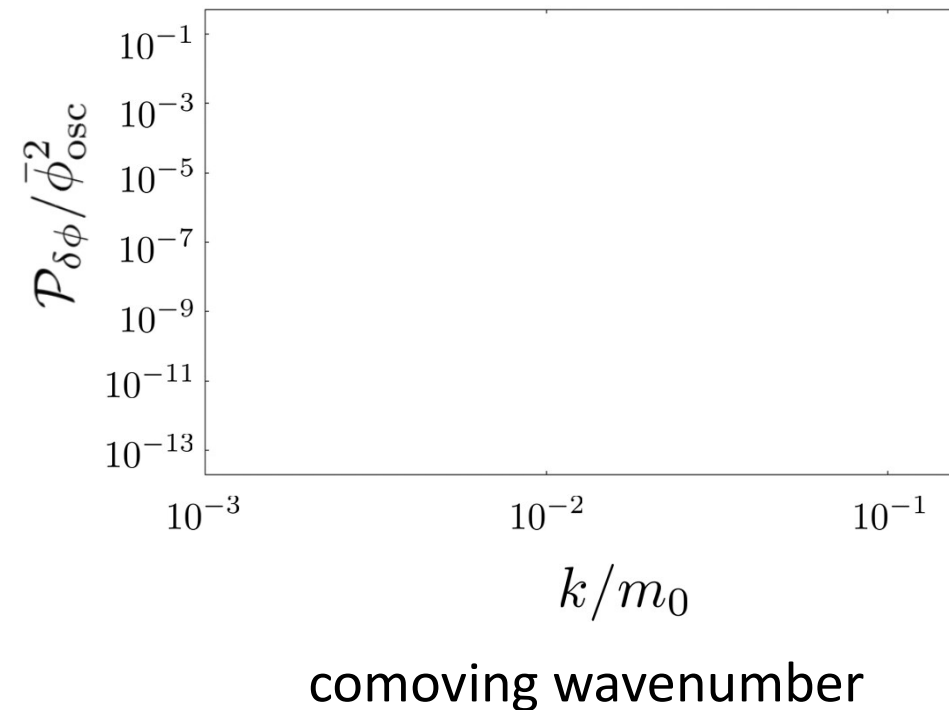
$$n > 1$$

Non-perturbative decay (parametric self-resonance)



Power spectrum:

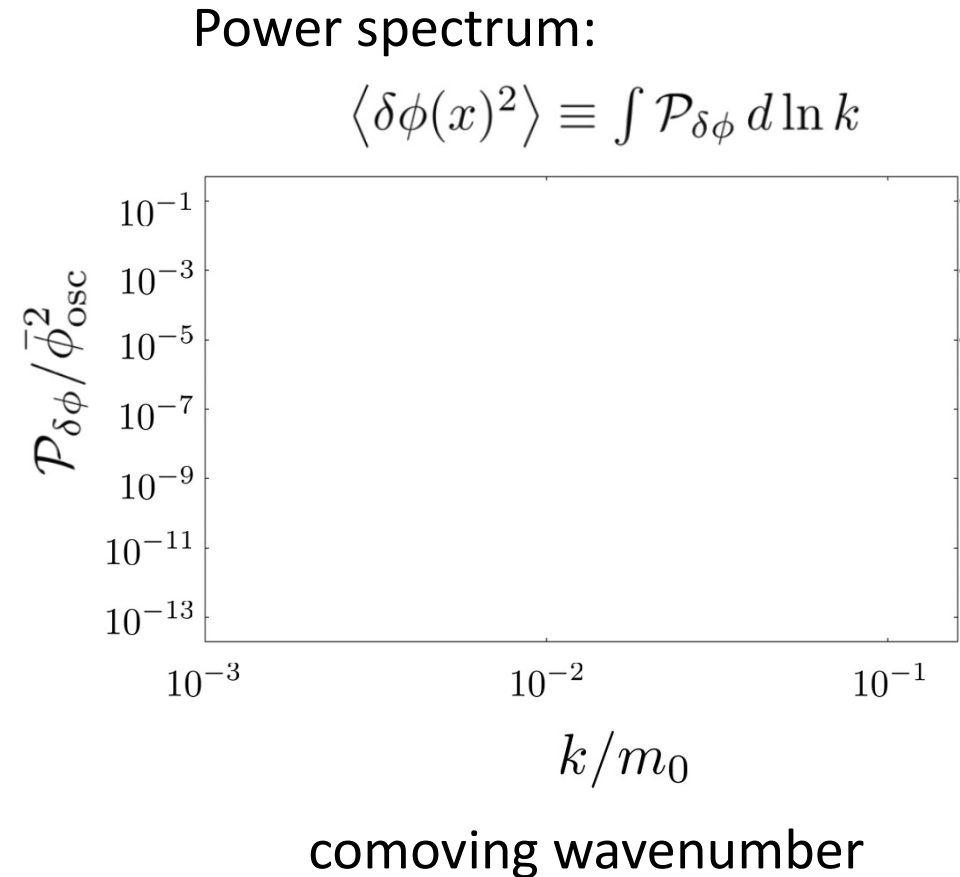
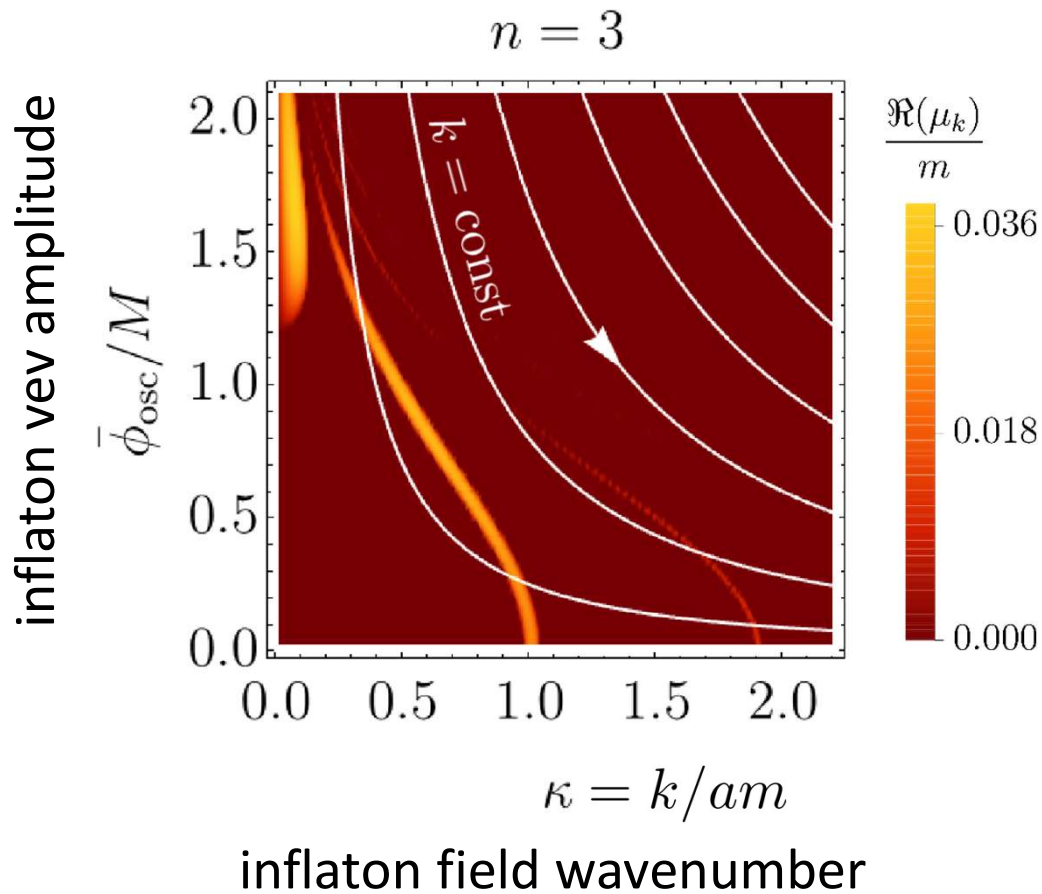
$$\langle \delta\phi(x)^2 \rangle \equiv \int \mathcal{P}_{\delta\phi} d \ln k$$



Towards radiation domination

$$n > 1$$

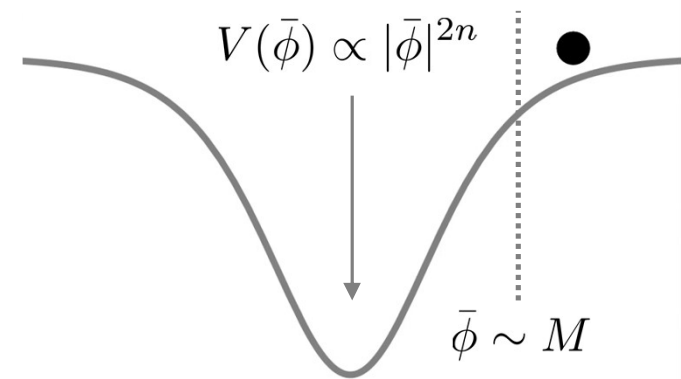
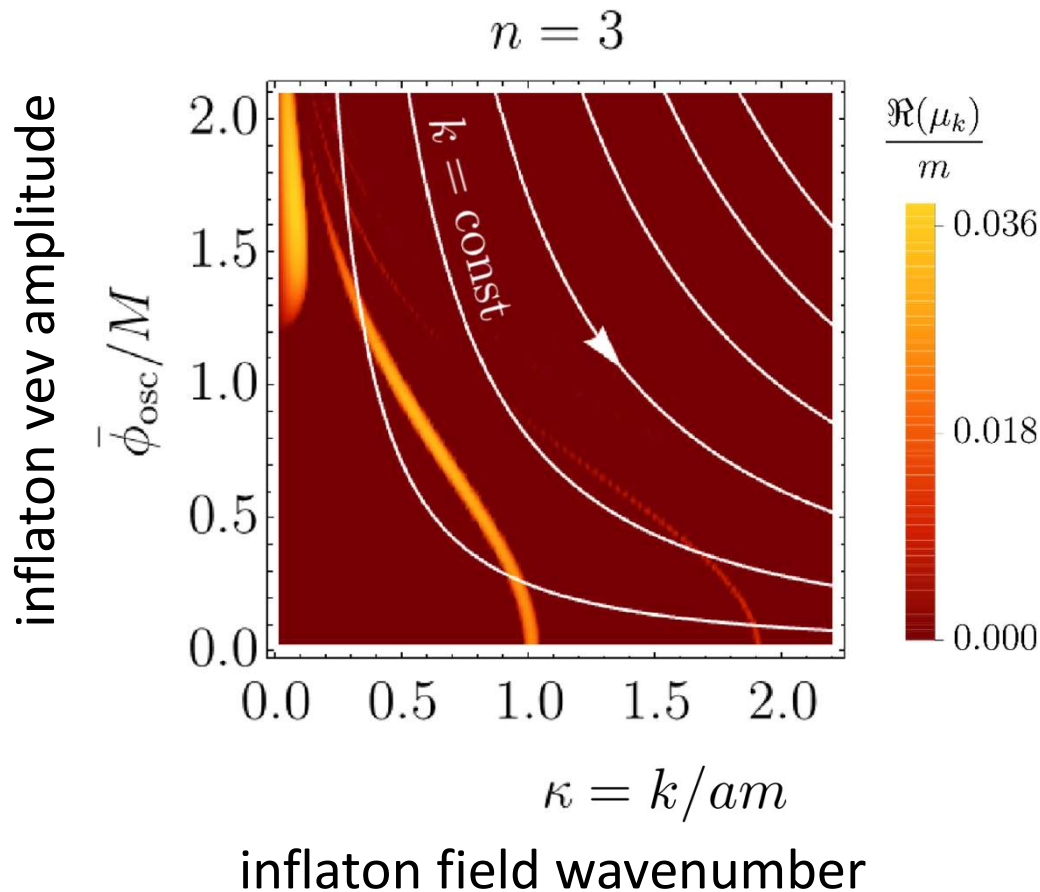
Non-perturbative decay (parametric self-resonance)



Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)



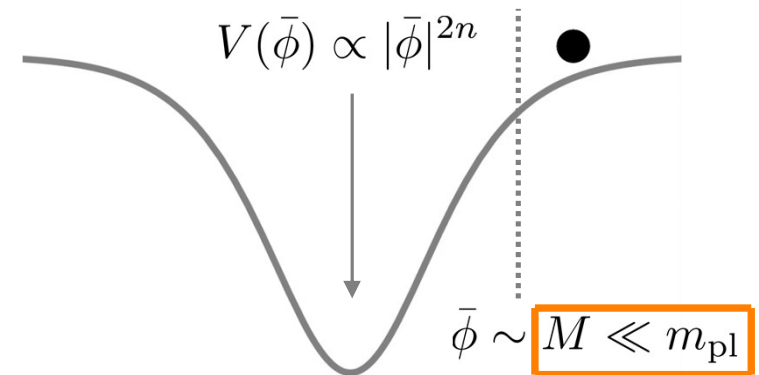
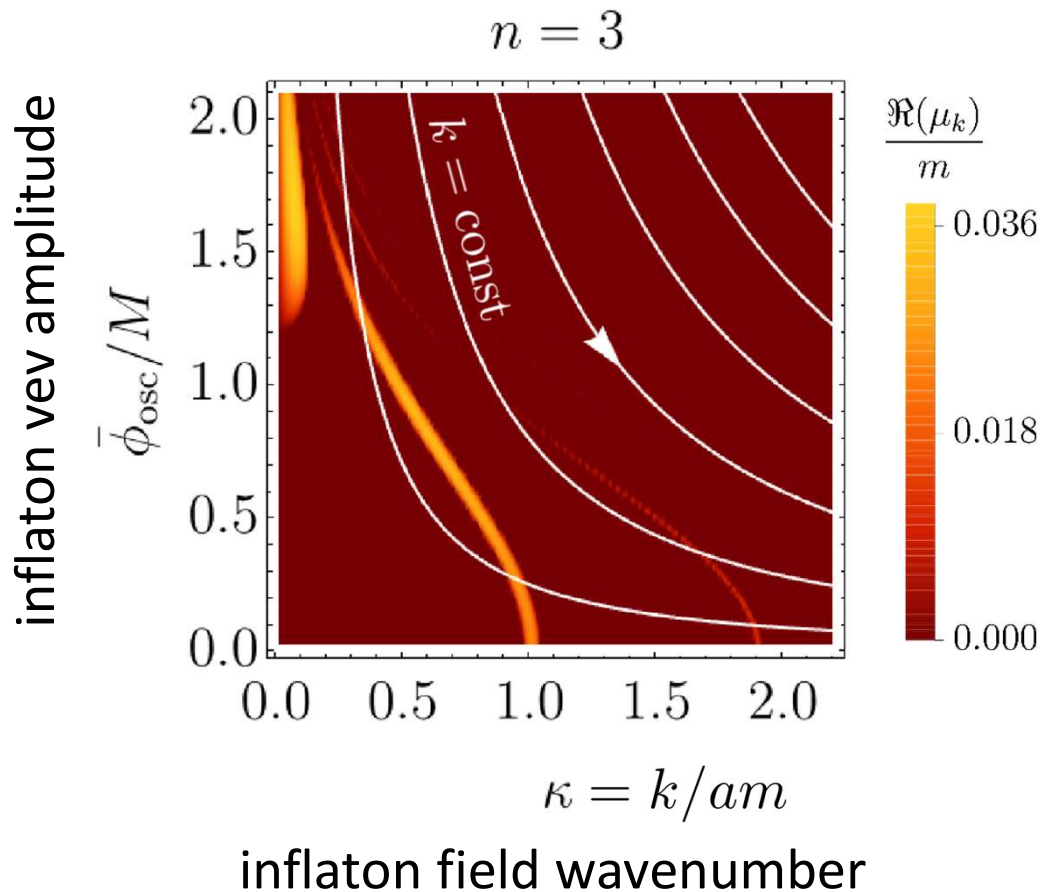
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Towards radiation domination

$$n > 1$$

Non-perturbative decay (parametric self-resonance)



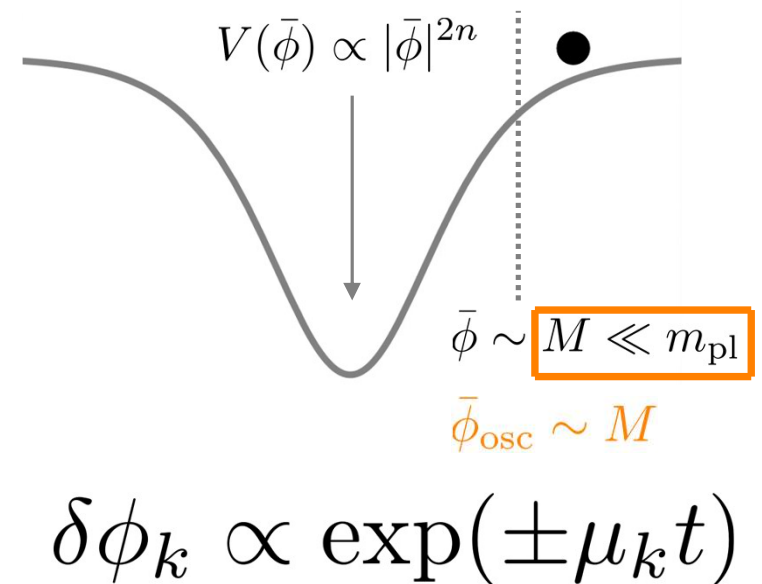
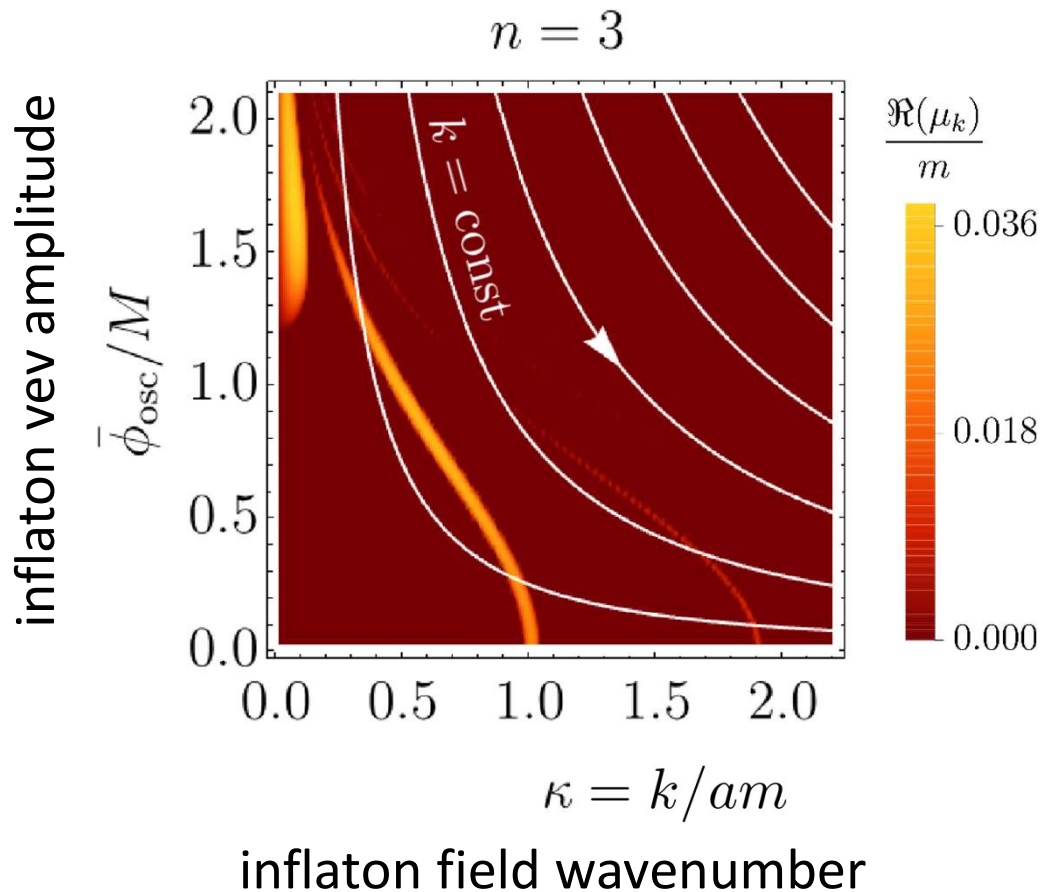
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Towards radiation domination

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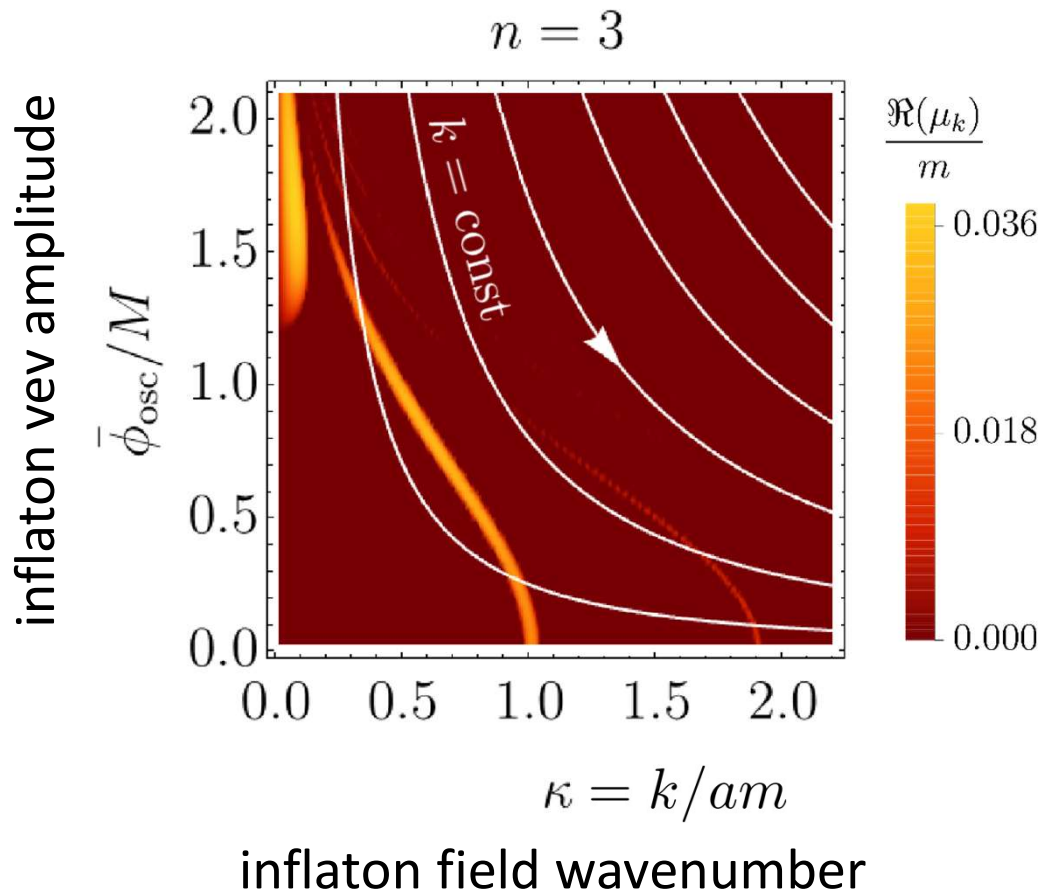


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Towards radiation domination

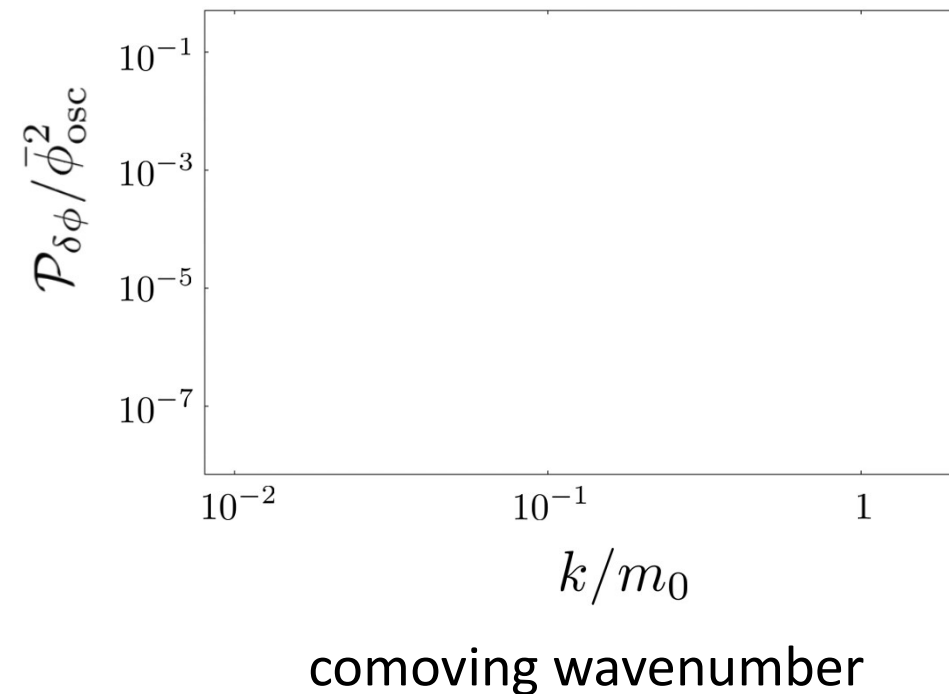
$$n > 1$$

Non-perturbative decay (parametric self-resonance)



Power spectrum:

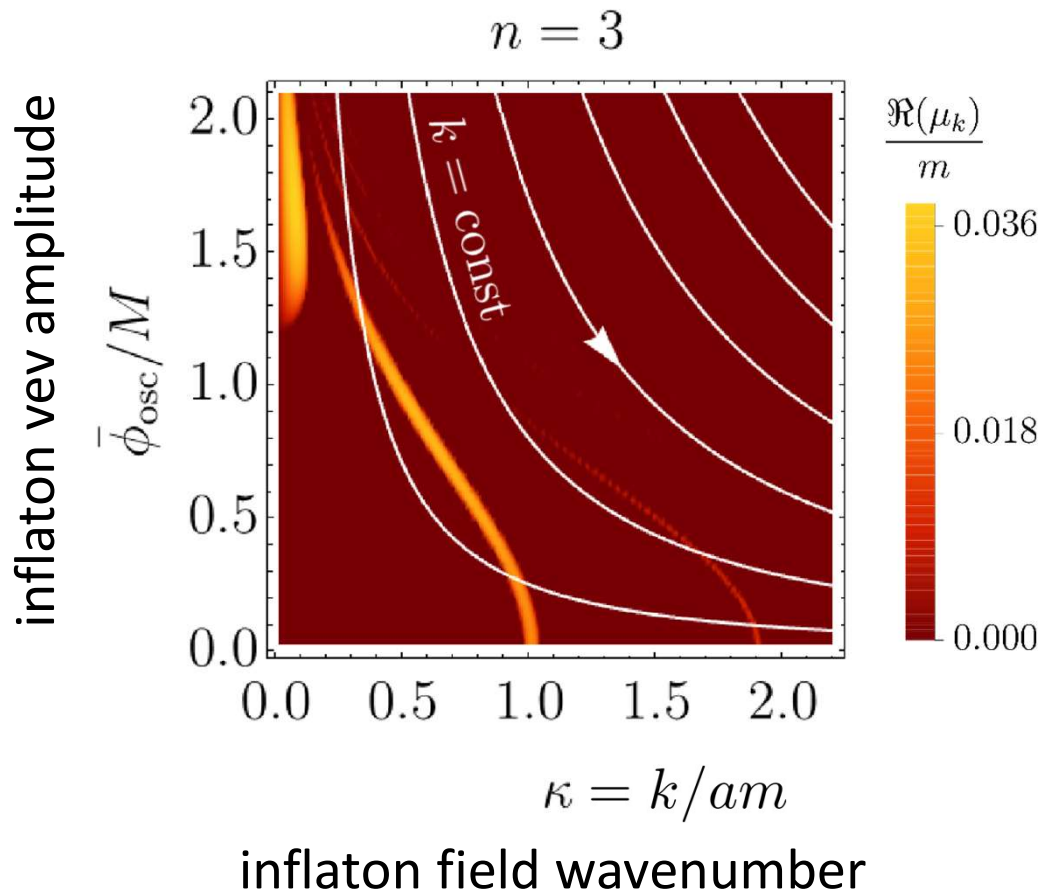
$$\langle \delta\phi(x)^2 \rangle \equiv \int \mathcal{P}_{\delta\phi} d \ln k$$



Towards radiation domination

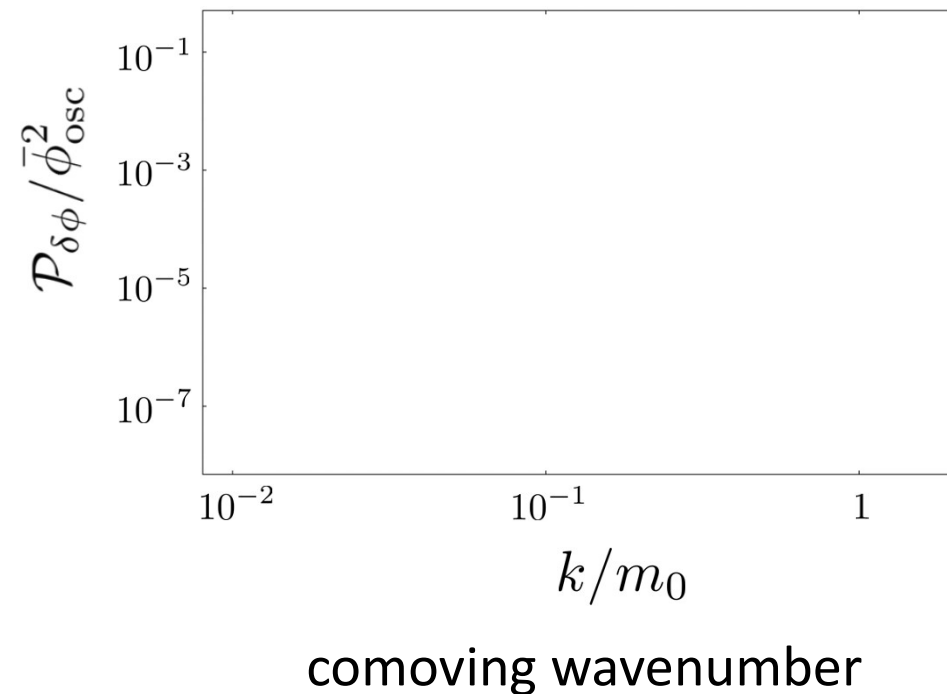
$$n > 1$$

Non-perturbative decay (parametric self-resonance)



Power spectrum:

$$\langle \delta\phi(x)^2 \rangle \equiv \int \mathcal{P}_{\delta\phi} d \ln k$$

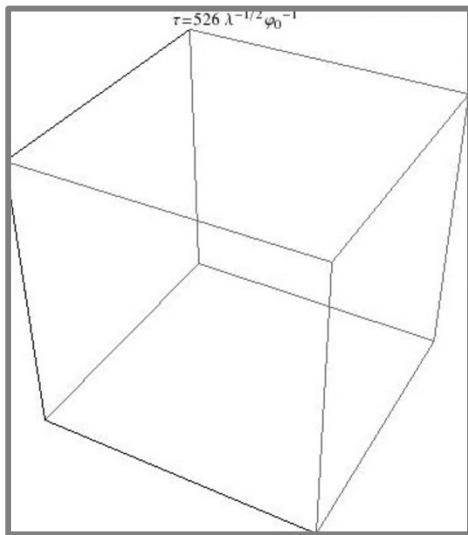


Towards radiation domination

$$n > 1$$

$$M \ll m_{\text{pl}}$$

$$M \sim m_{\text{pl}}$$



(transients)

- $\bar{\phi}$ fragments quickly

- slow production of $\delta\phi(t, \mathbf{x})$
- $\bar{\phi}$ fragments gradually

$$\Delta N_{\text{fr}} \approx \frac{n+1}{3} \ln \left(10^3 \frac{M}{m_{\text{pl}}} \right)$$

at sufficiently late times:

$$\phi \text{ virialized + turbulent} \rightarrow w = \frac{1}{3}$$

Expansion history effects

Expansion history effects

Spectral index: n_s

Tensor-to-scalar ratio: r

Expansion history effects

Spectral index: $n_s = n_s(M, n, N_*)$

Tensor-to-scalar ratio: $r = r(M, n, N_*)$

Expansion history effects

Spectral index: $n_s = n_s(M, n, \textcircled{N_*})$

Tensor-to-scalar ratio: $r = r(M, n, \textcircled{N_*})??$

Expansion history effects

Spectral index: $n_s = n_s(M, n, N_*)$

Tensor-to-scalar ratio: $r = r(M, n, N_*)??$

$$50 \leq N_* \leq 60$$

Expansion history effects

Spectral index: $n_s = n_s(M, n, \textcircled{N_*})$

Tensor-to-scalar ratio: $r = r(M, n, \textcircled{N_*})$ ✓

$$\text{50} \leq \text{N}_* \leq \text{60}$$

Expansion history effects

Spectral index: $n_s = n_s(M, n, N_*)$

Tensor-to-scalar ratio: $r = r(M, n, N_*)$ ✓

$$N_* = 66.89 - \frac{1}{12} \ln g_{\text{th}} + \frac{1}{4} \ln \frac{V_*^4}{m_{\text{pl}}^4 \rho_{\text{end}}} - \ln \frac{k_*}{a_0 H_0} + \frac{3\bar{w}_{\text{int}} - 1}{4} \Delta N_{\text{rad}}$$

$$\Delta N_{\text{rad}} \equiv \int_{a_{\text{end}}}^{a_{\text{rad}}} d \ln a$$

Expansion history effects

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reheating

$$\Delta N_{\text{rad}} \equiv \int_{a_{\text{end}}}^{a_{\text{rad}}} d \ln a$$

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$$\Delta N_{\text{fr}} = \begin{cases} 1 & \text{if } M \ll m_{\text{pl}} \\ \frac{n+1}{3} \ln \left(10 \frac{\kappa}{\Delta \kappa} \frac{M}{m_{\text{pl}}} \right) & \text{if } M \sim m_{\text{pl}} \end{cases}$$

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Expansion history effects

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$$\Delta N_{\text{rad}} \leq \Delta N_{\text{fr}} = \begin{cases} 1 & \text{if } M \ll m_{\text{pl}} \\ \frac{n+1}{3} \ln \left(10 \frac{\kappa}{\Delta \kappa} \frac{M}{m_{\text{pl}}} \right) & \text{if } M \sim m_{\text{pl}} \end{cases}$$

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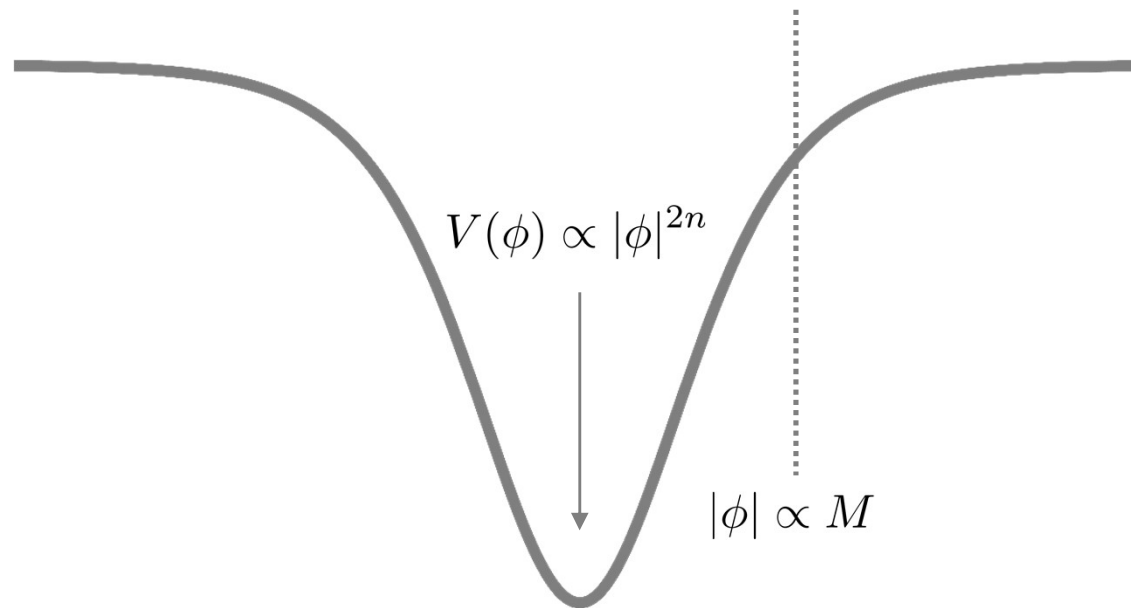
$$w_{\text{int}}(\Delta N) = \begin{cases} \frac{n-1}{n+1} & \text{if } 0 < \Delta N < \Delta N_{\text{rad}} \\ \frac{1}{3} & \text{if } \Delta N > \Delta N_{\text{rad}} \end{cases}$$

$$\Delta N_{\text{rad}} \equiv \int_{a_{\text{end}}}^{a_{\text{rad}}} d \ln a$$

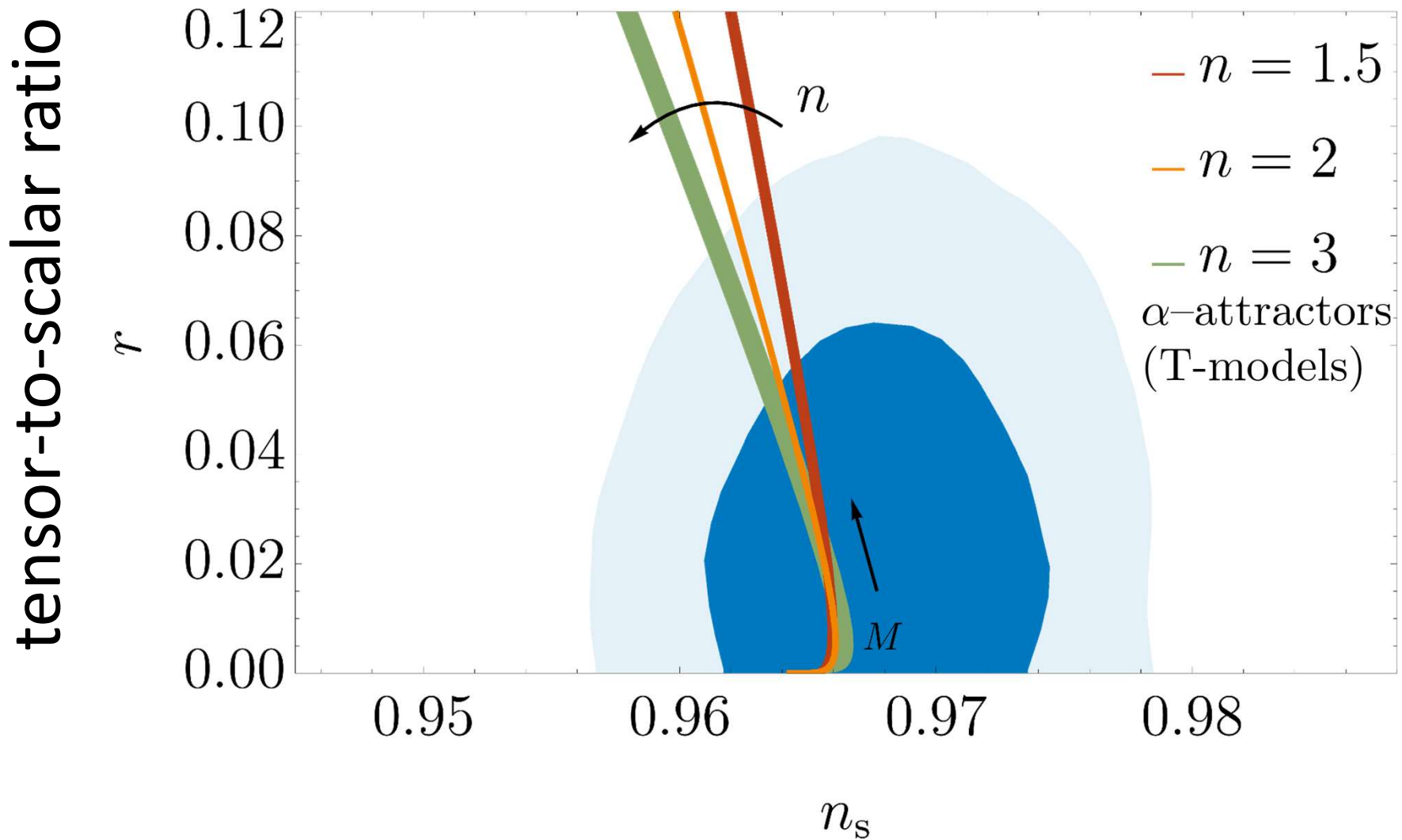
Expansion history effects

$$V(\phi) \propto \tanh^{2n} \left(\frac{|\phi|}{M} \right)$$

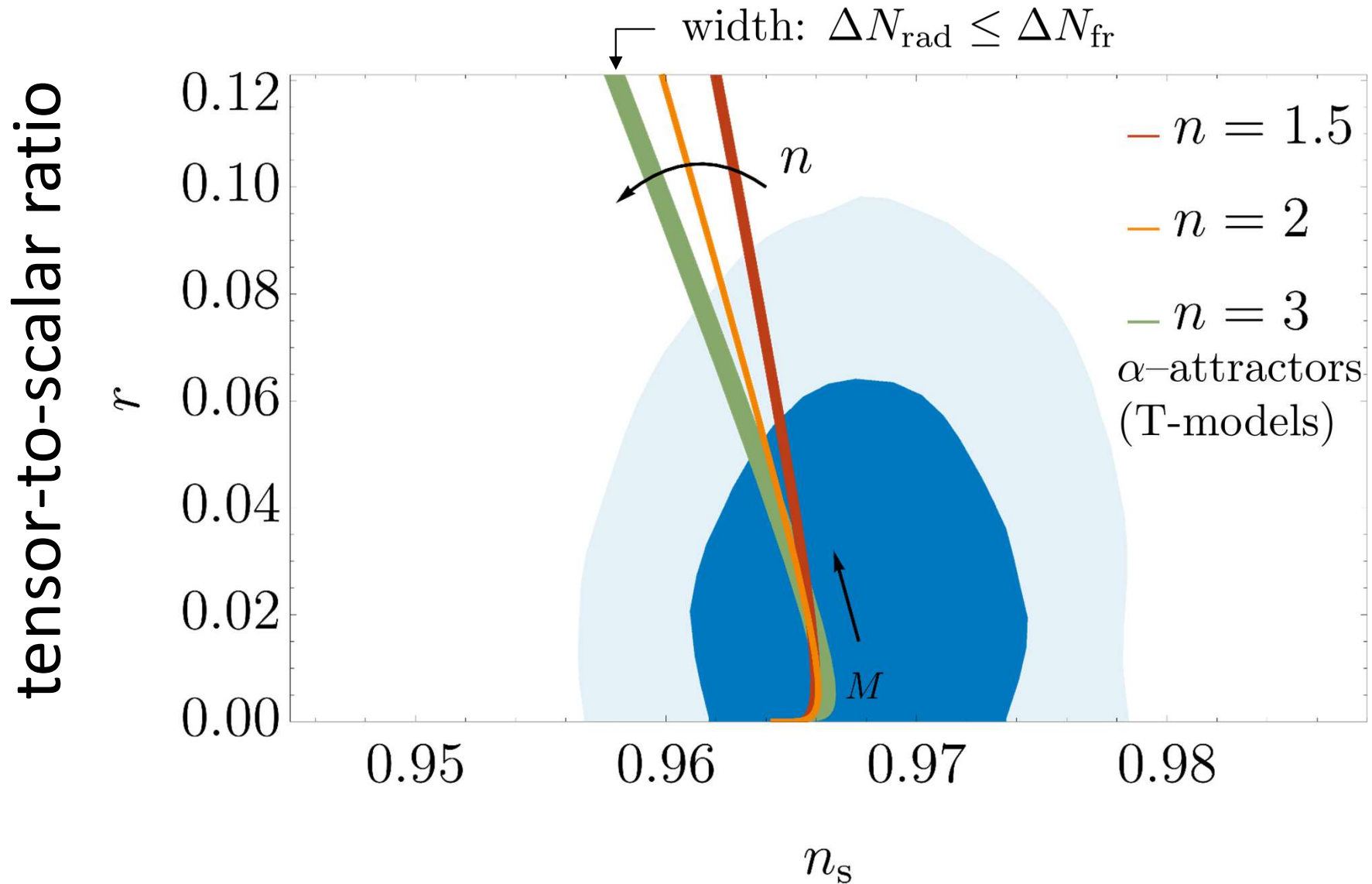
α -attractors
(T-models)



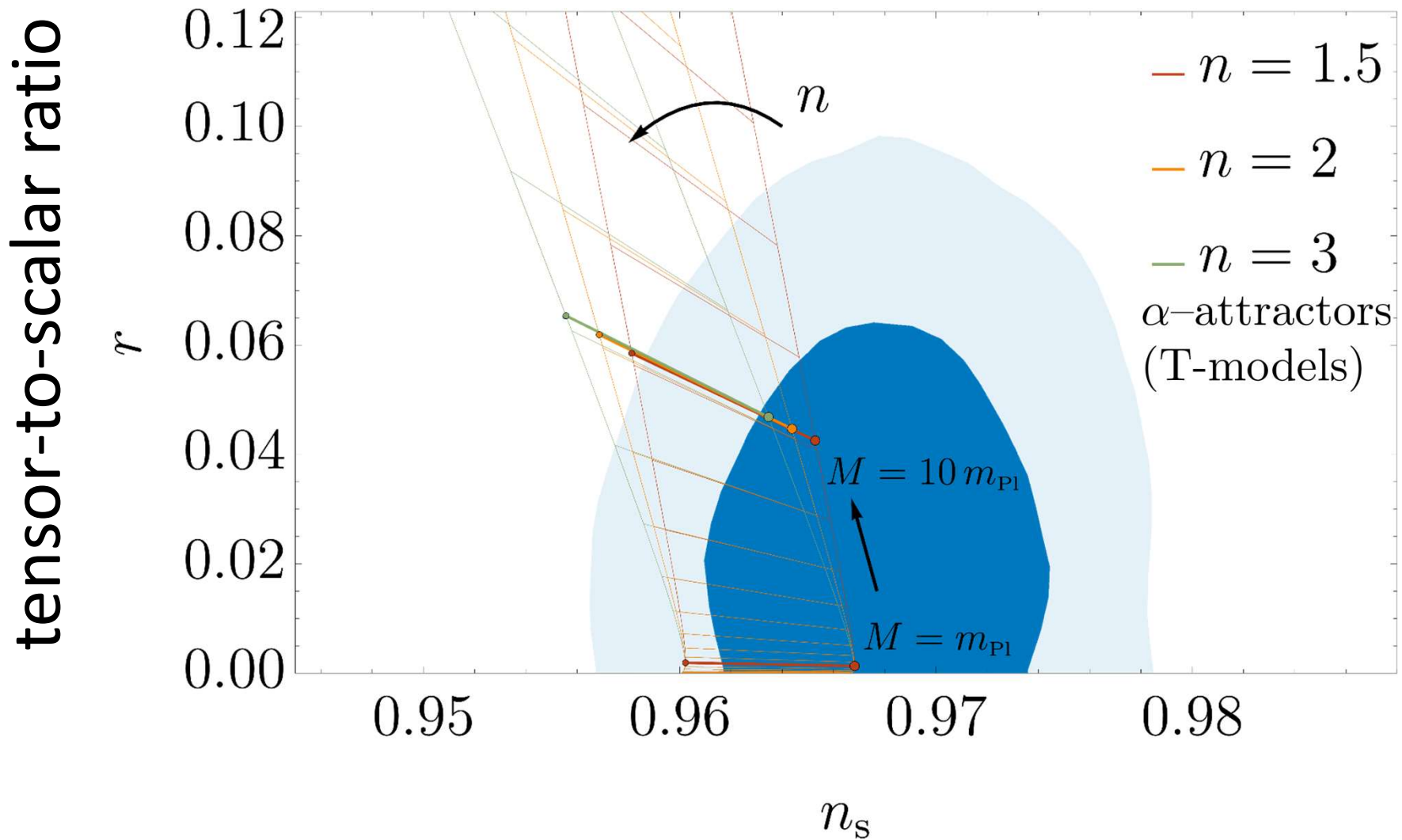
Expansion history effects



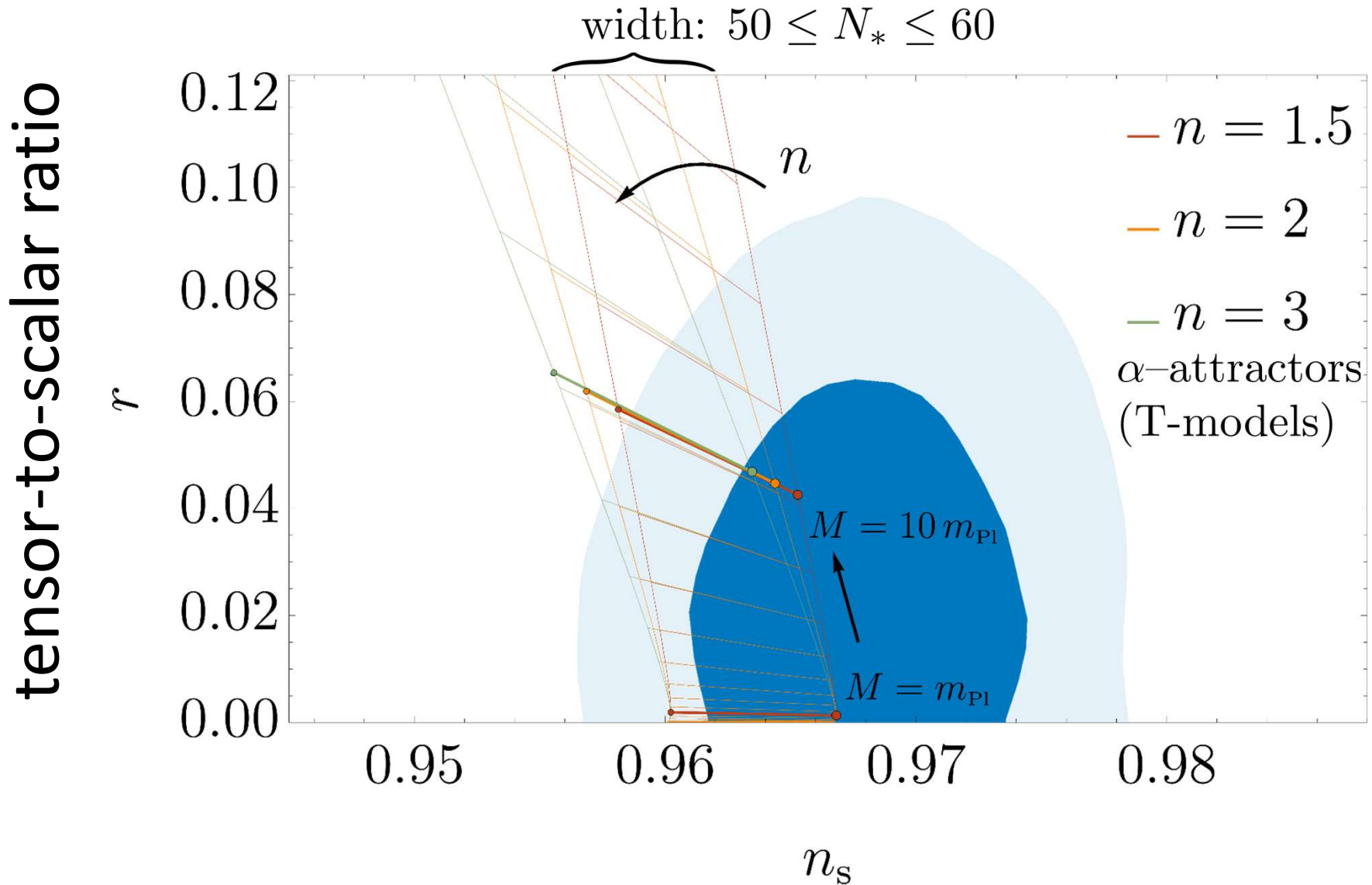
Expansion history effects



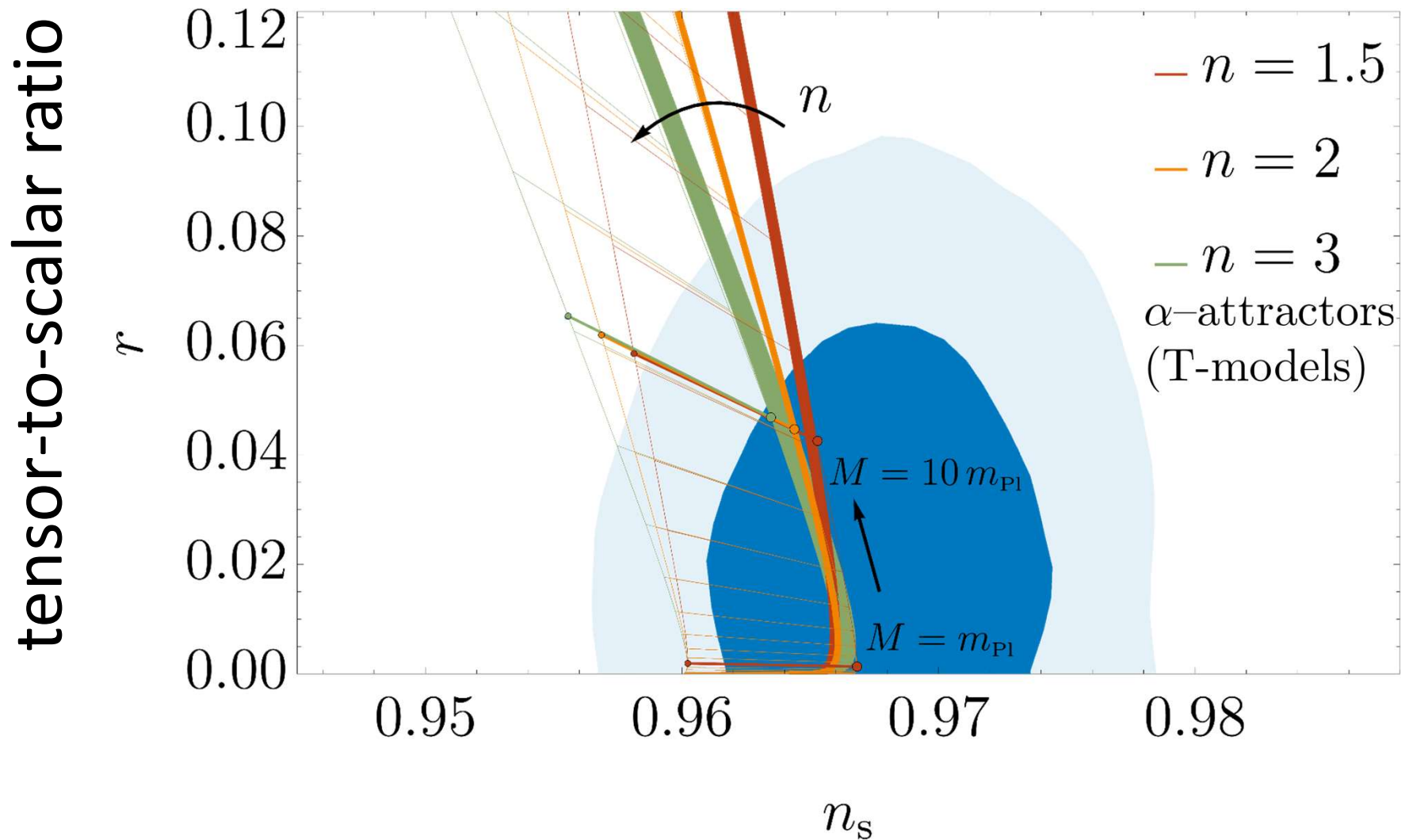
Expansion history effects



Expansion history effects



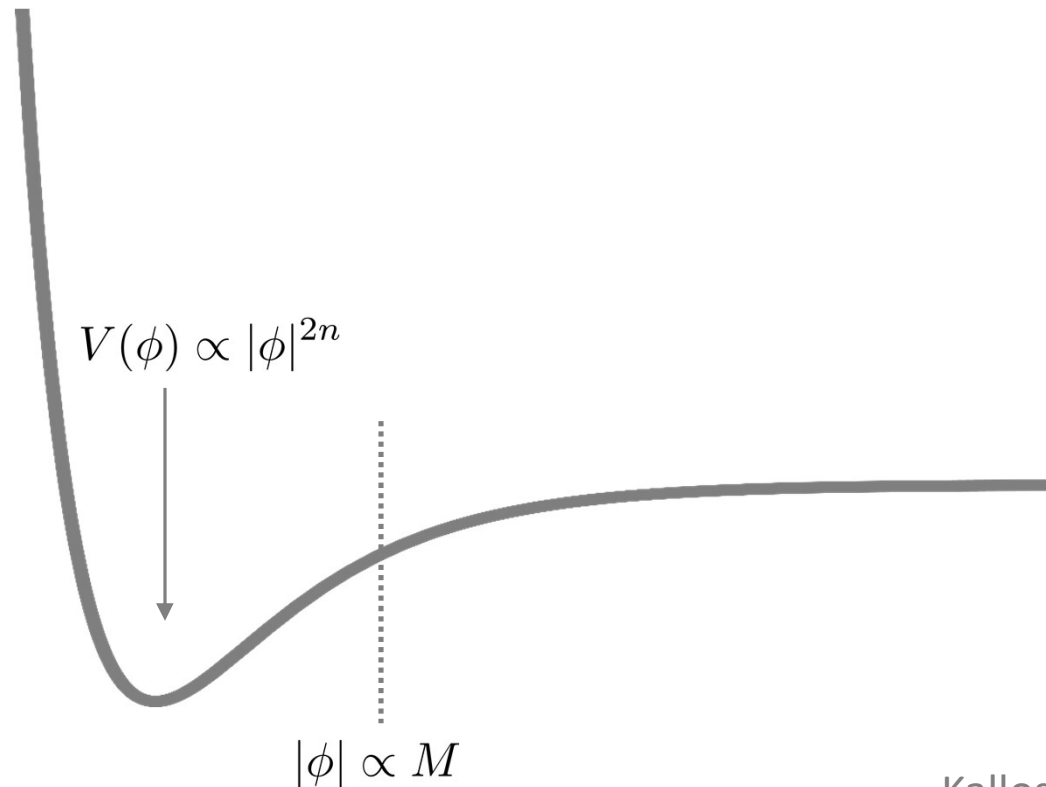
Expansion history effects



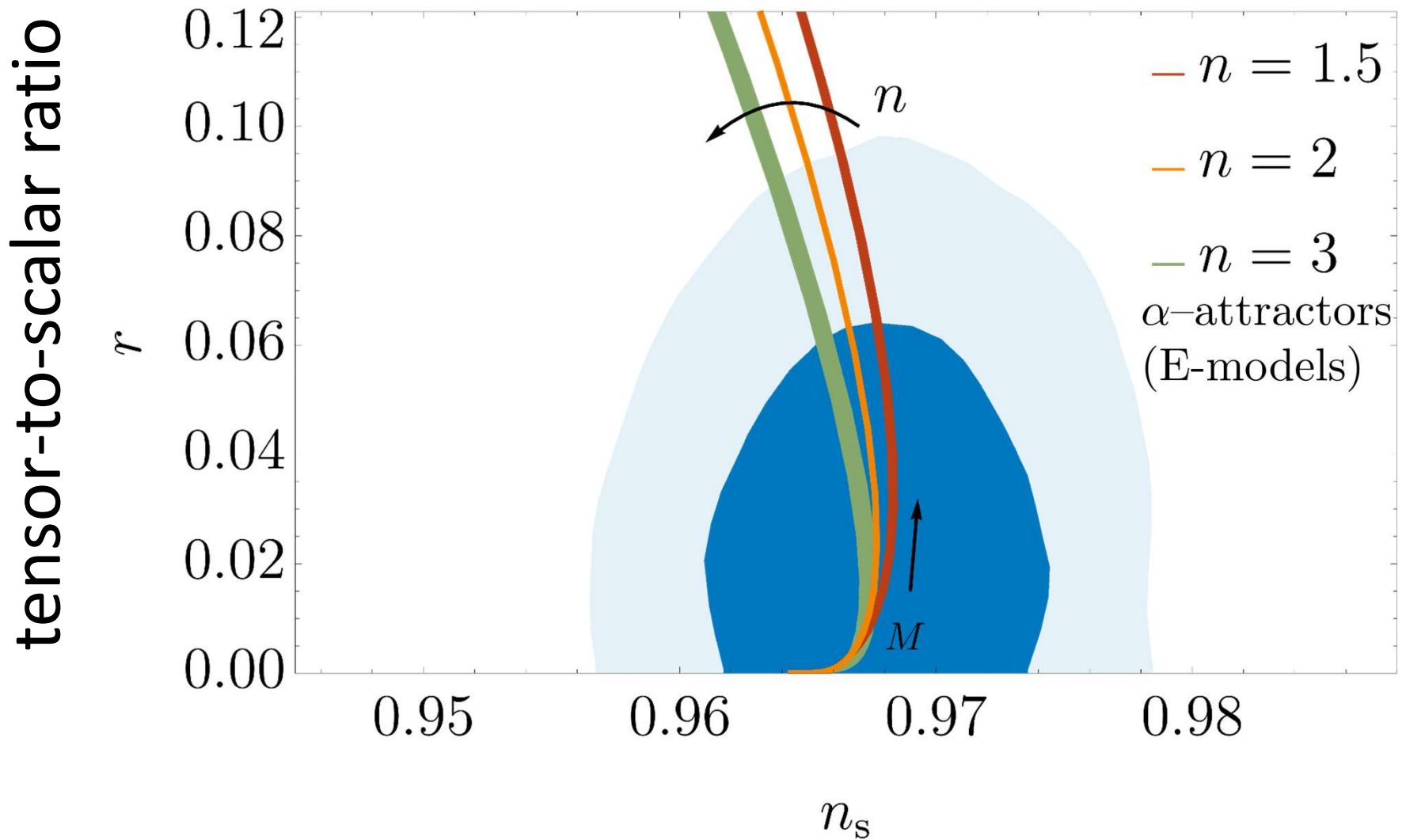
Expansion history effects

$$V(\phi) \propto \left| 1 - e^{-\phi/M} \right|^{2n}$$

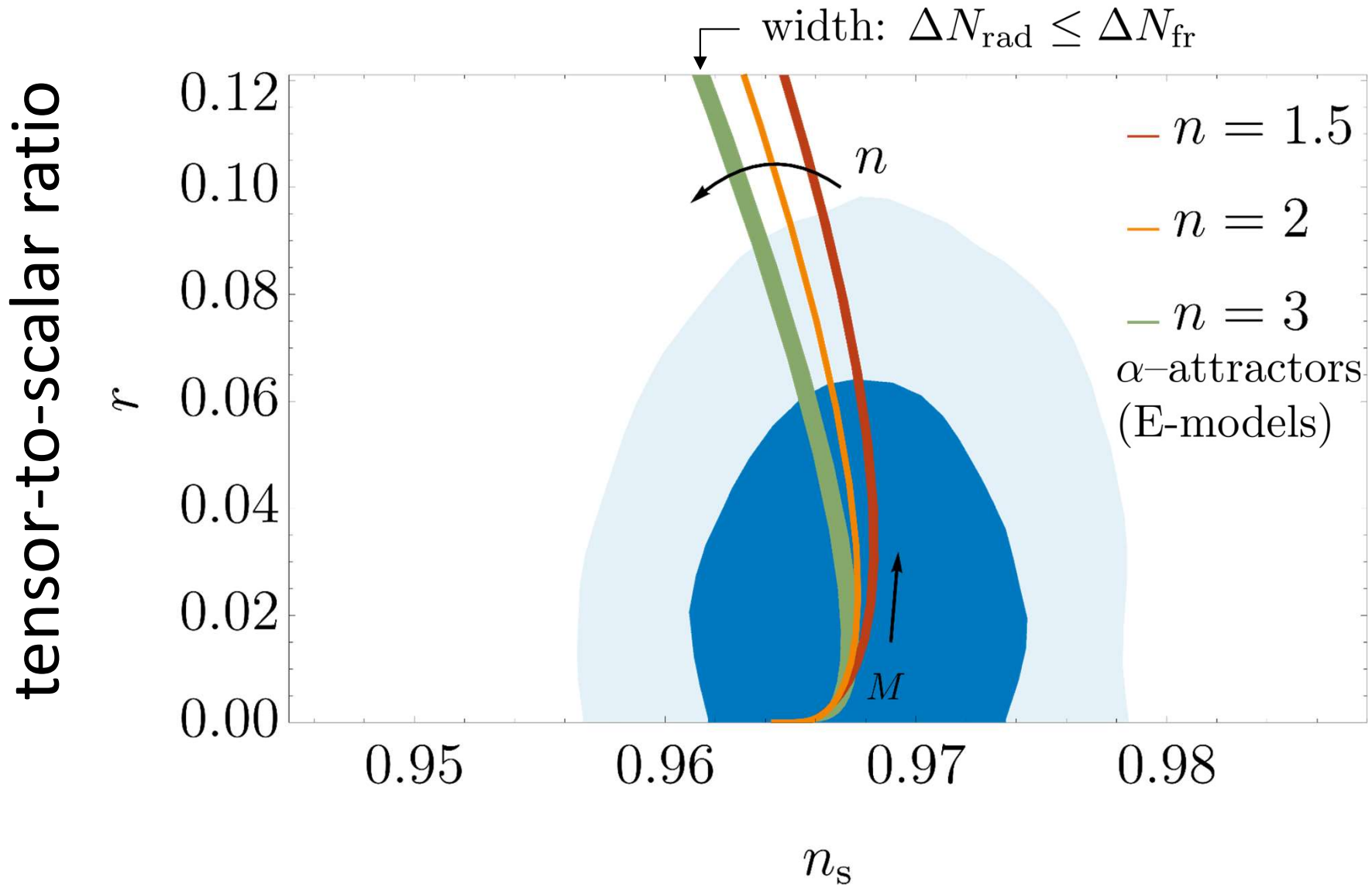
α -attractors
(E-models)



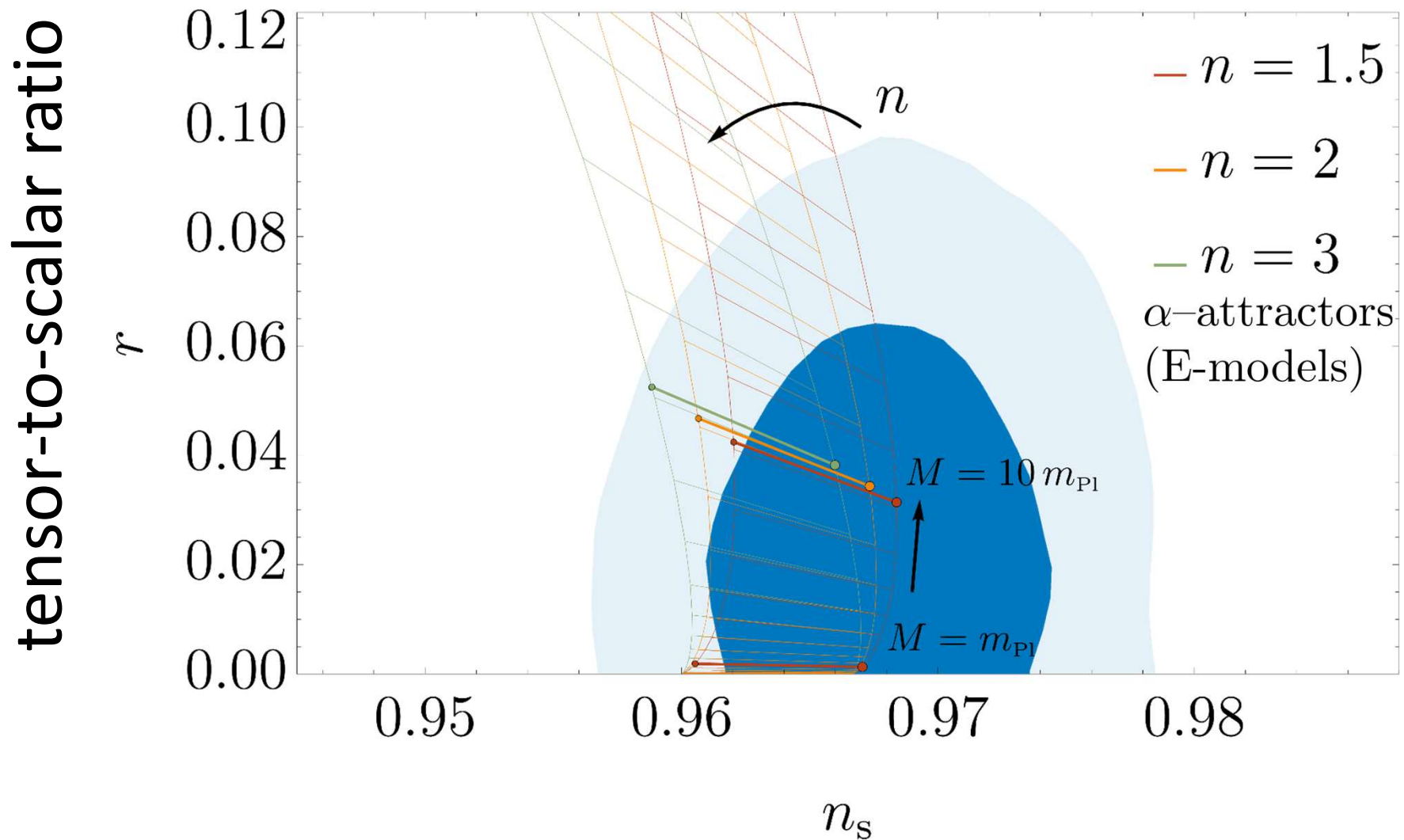
Expansion history effects



Expansion history effects

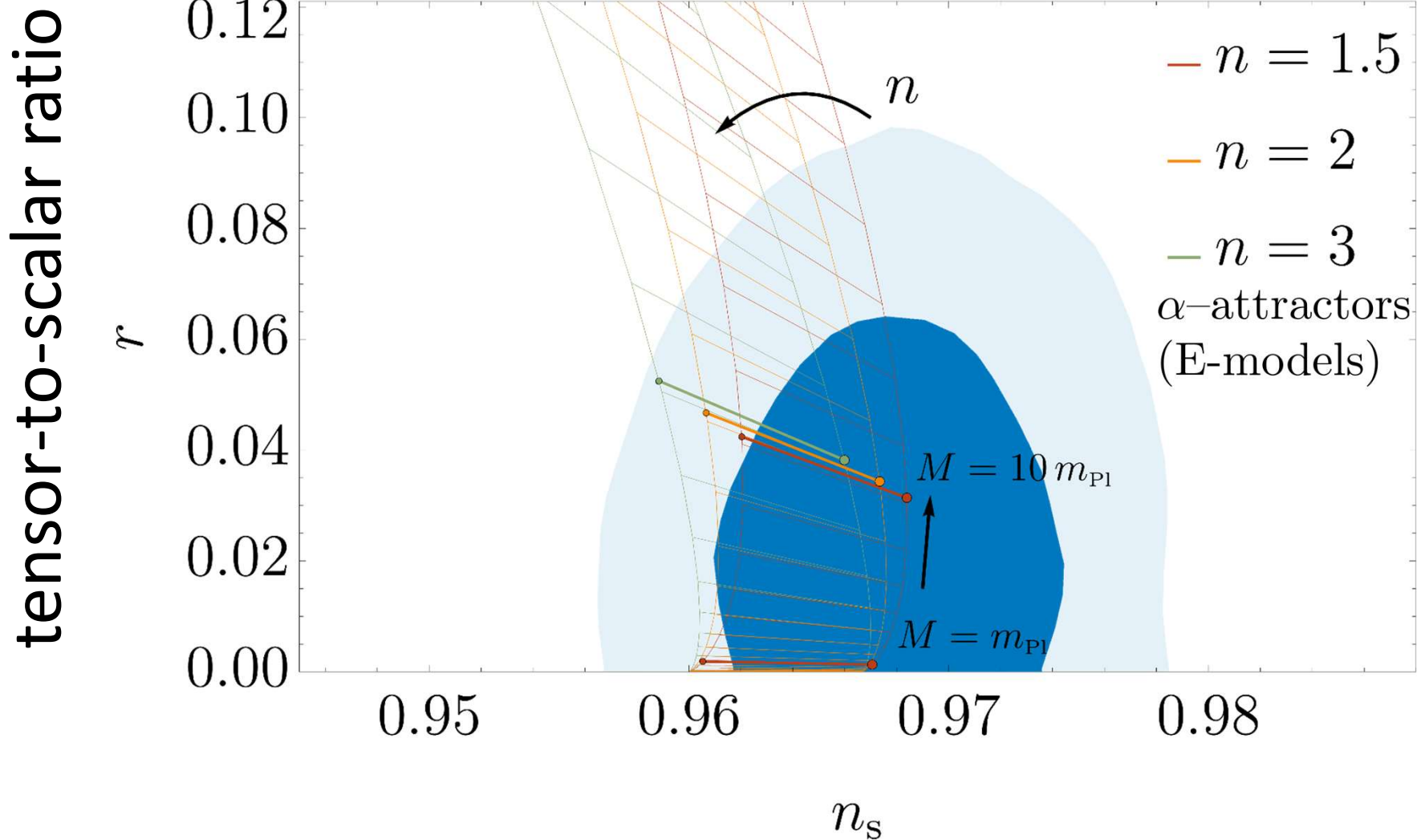


Expansion history effects

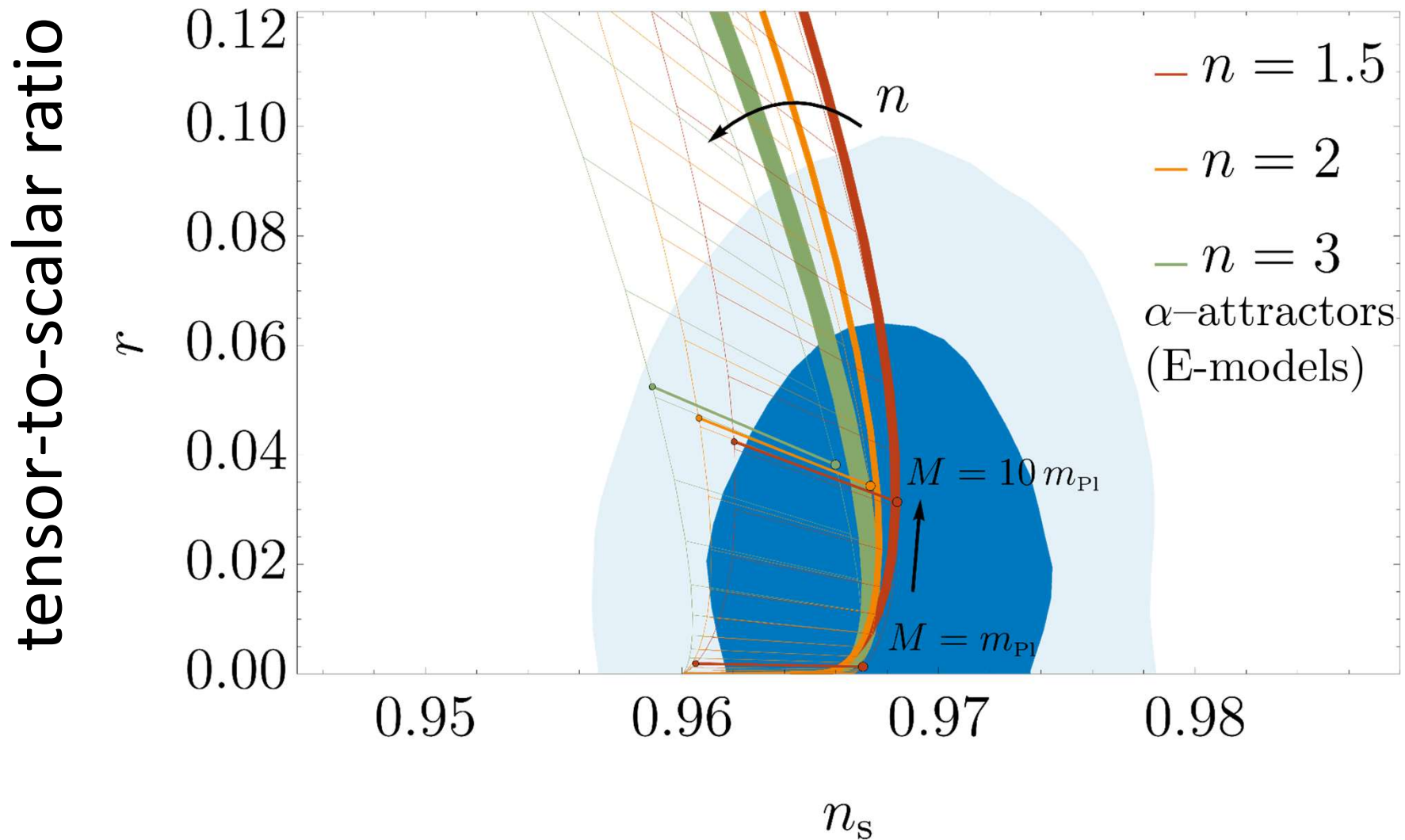


Expansion history effects

width: $50 \leq N_* \leq 60$

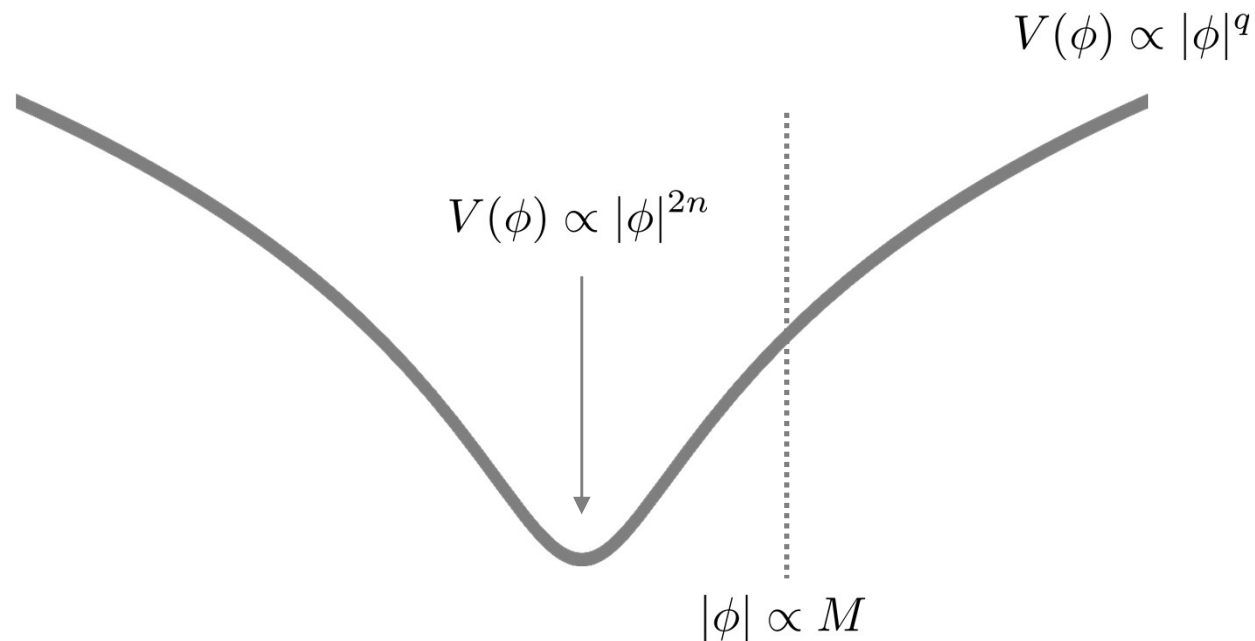


Expansion history effects



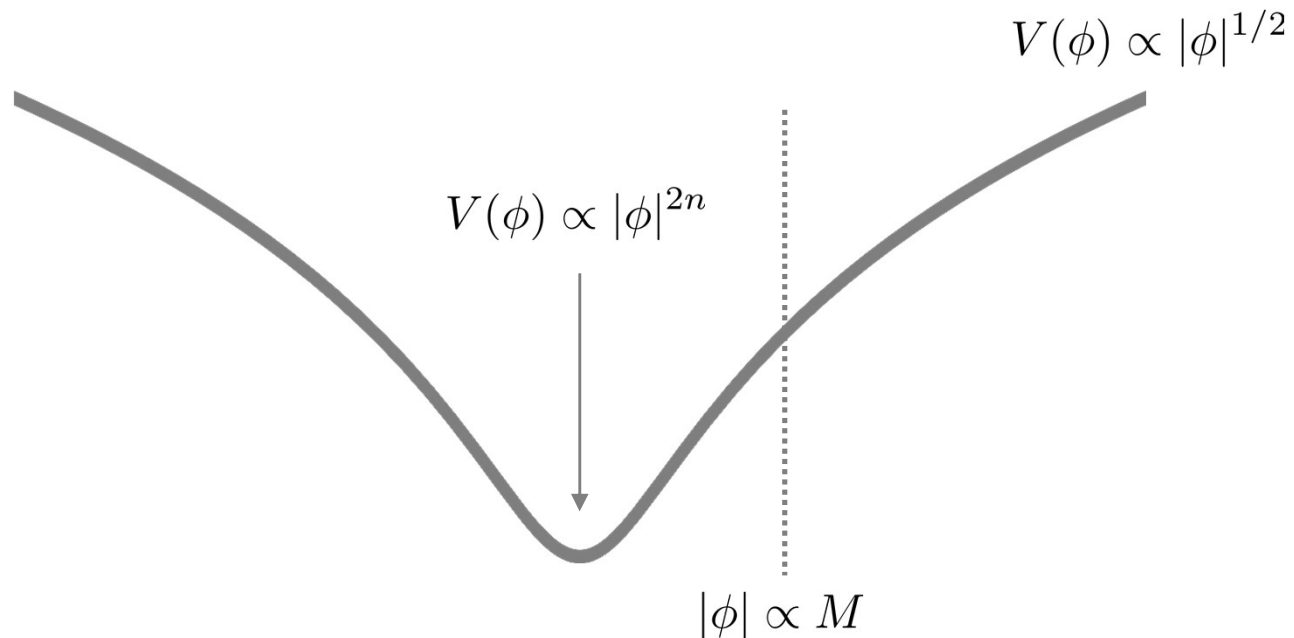
Expansion history effects

$$V(\phi) \propto \left[1 + \left| \frac{\phi}{M} \right|^{2n} \right]^{\frac{q}{2n}} - 1 \quad \text{Monodromy}$$

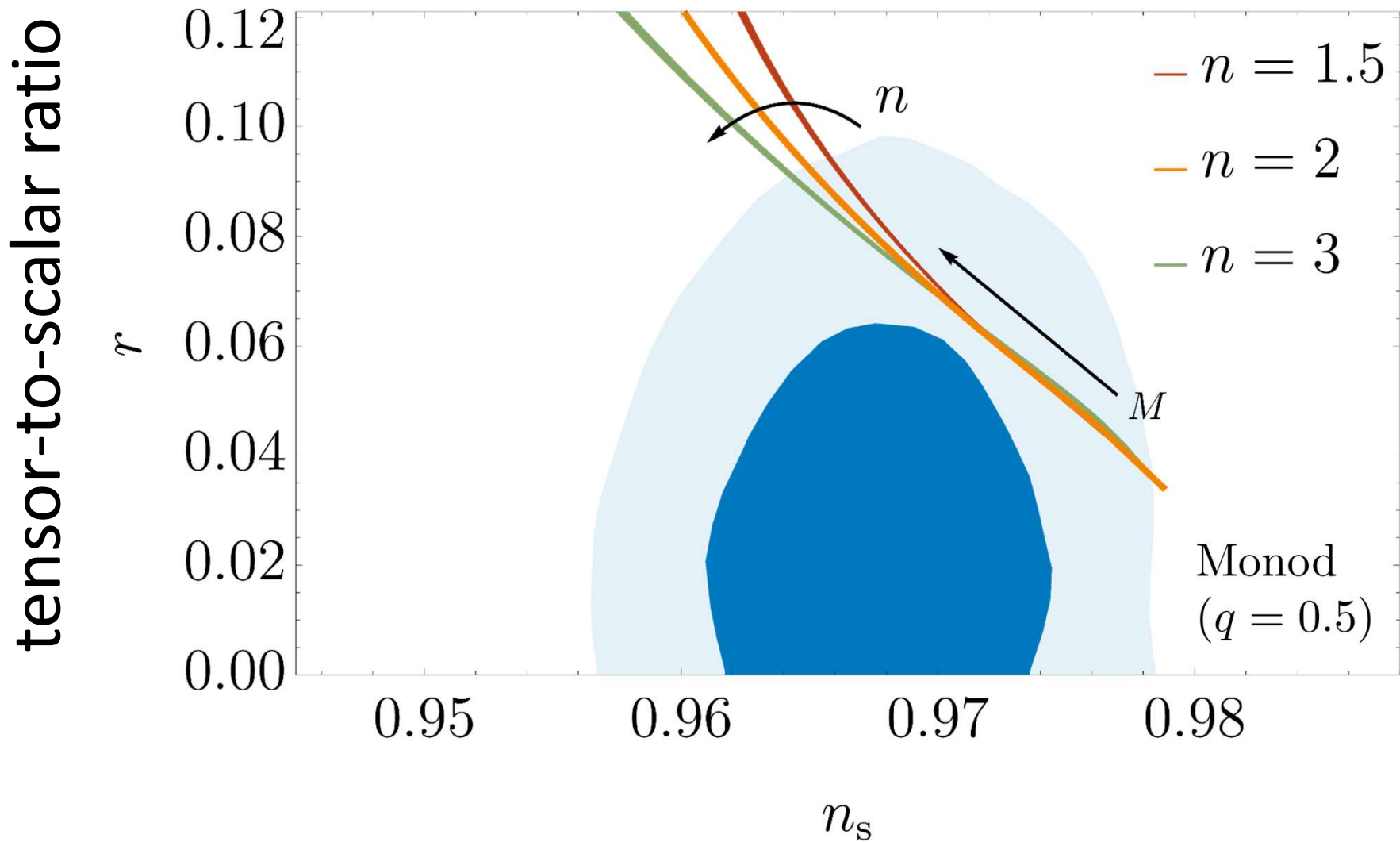


Expansion history effects

$$V(\phi) \propto \left[1 + \left| \frac{\phi}{M} \right|^{2n} \right]^{\frac{1}{4n}} - 1 \quad \text{Monodromy} \\ q = 1/2$$



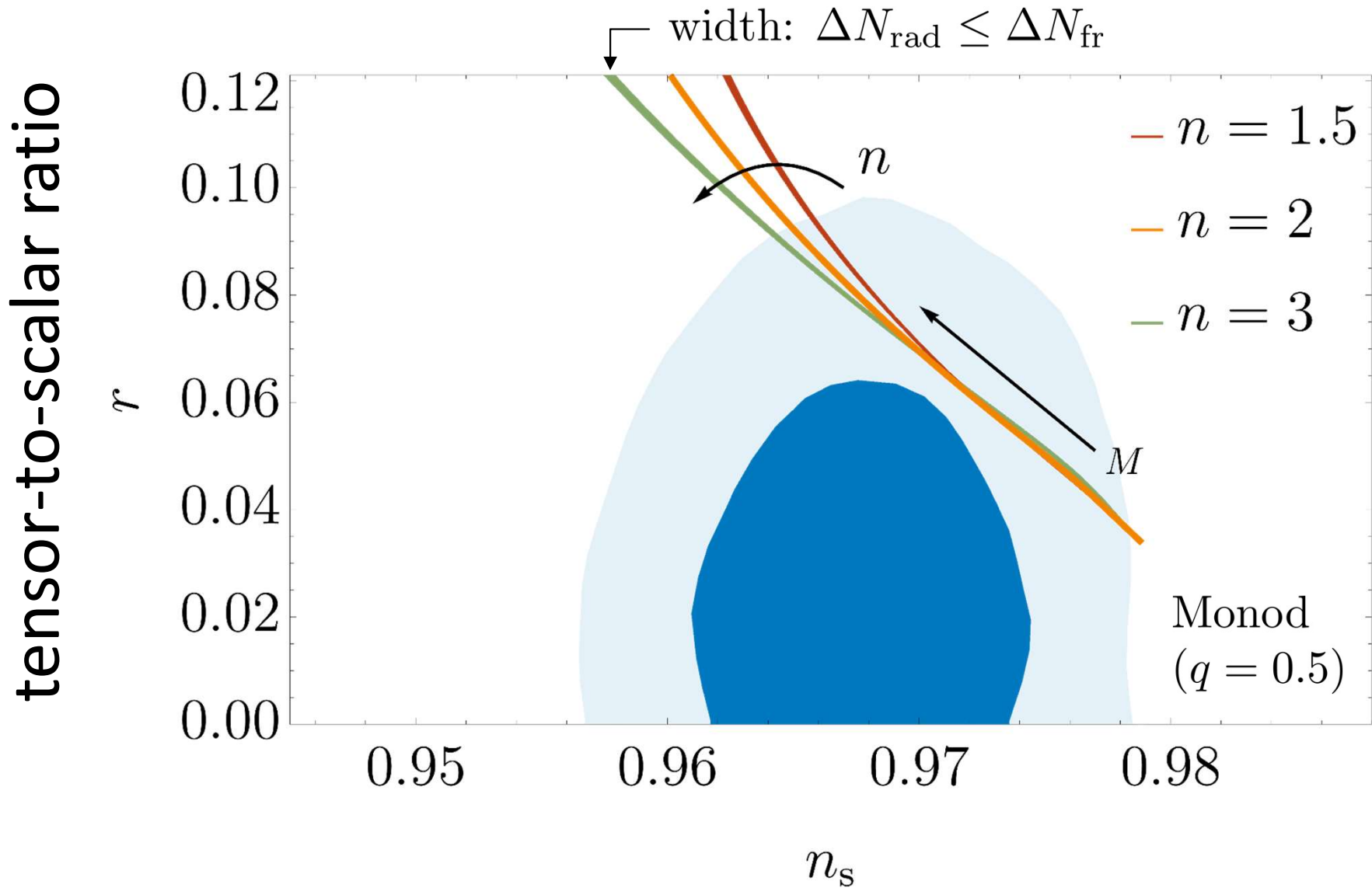
Expansion history effects



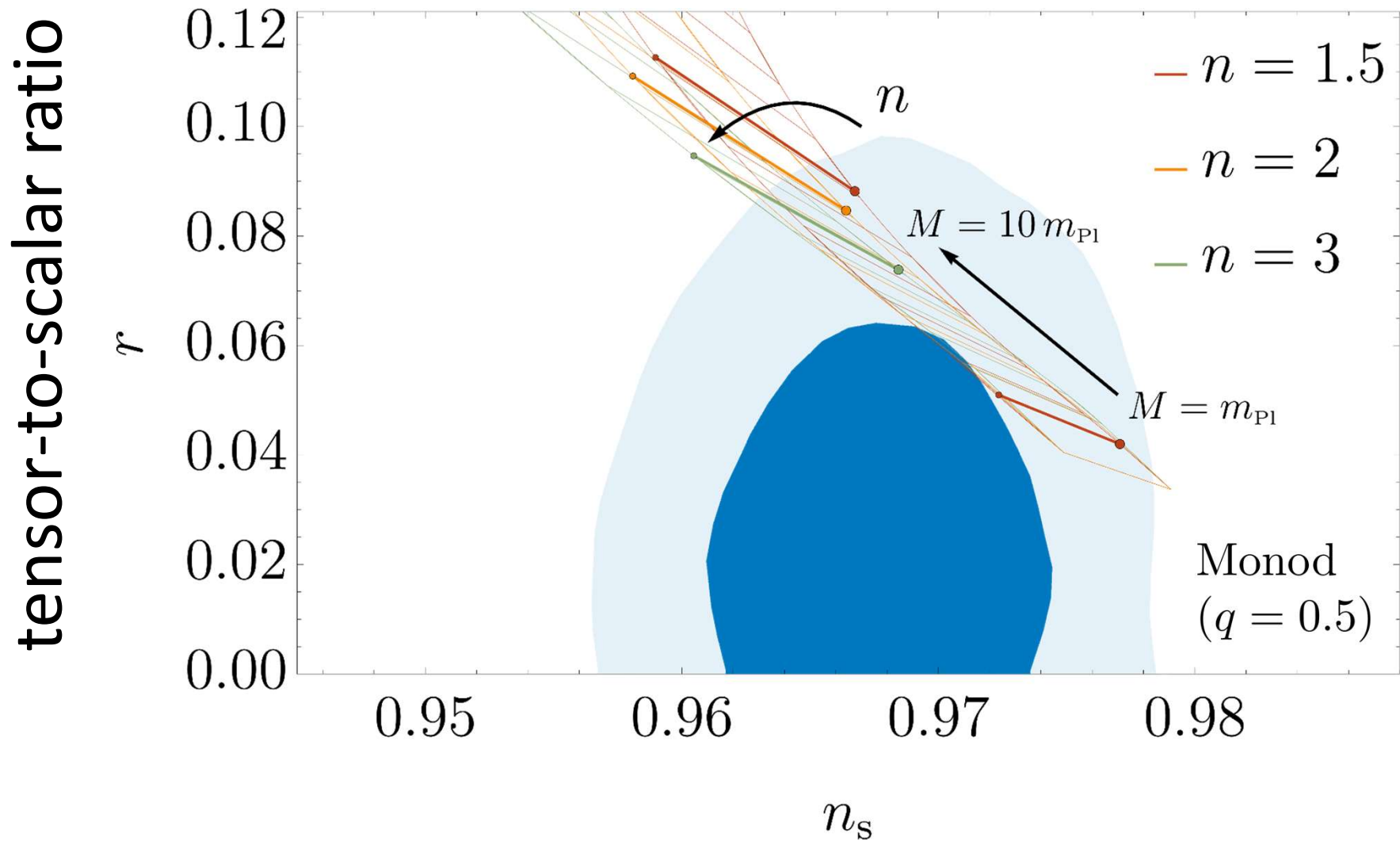
spectral index

KL and M. Amin (2018)
Planck Collaboration (2015)

Expansion history effects

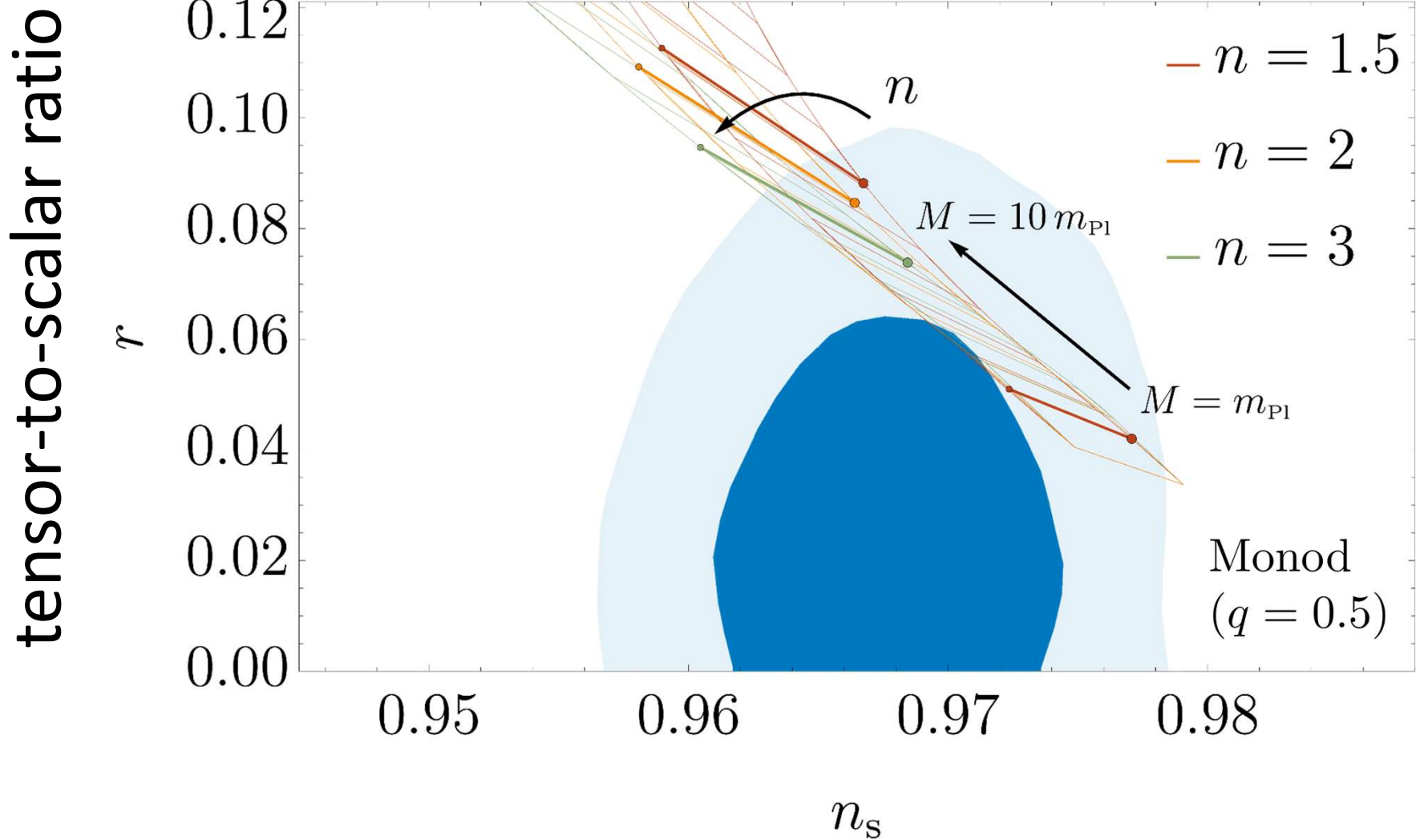


Expansion history effects

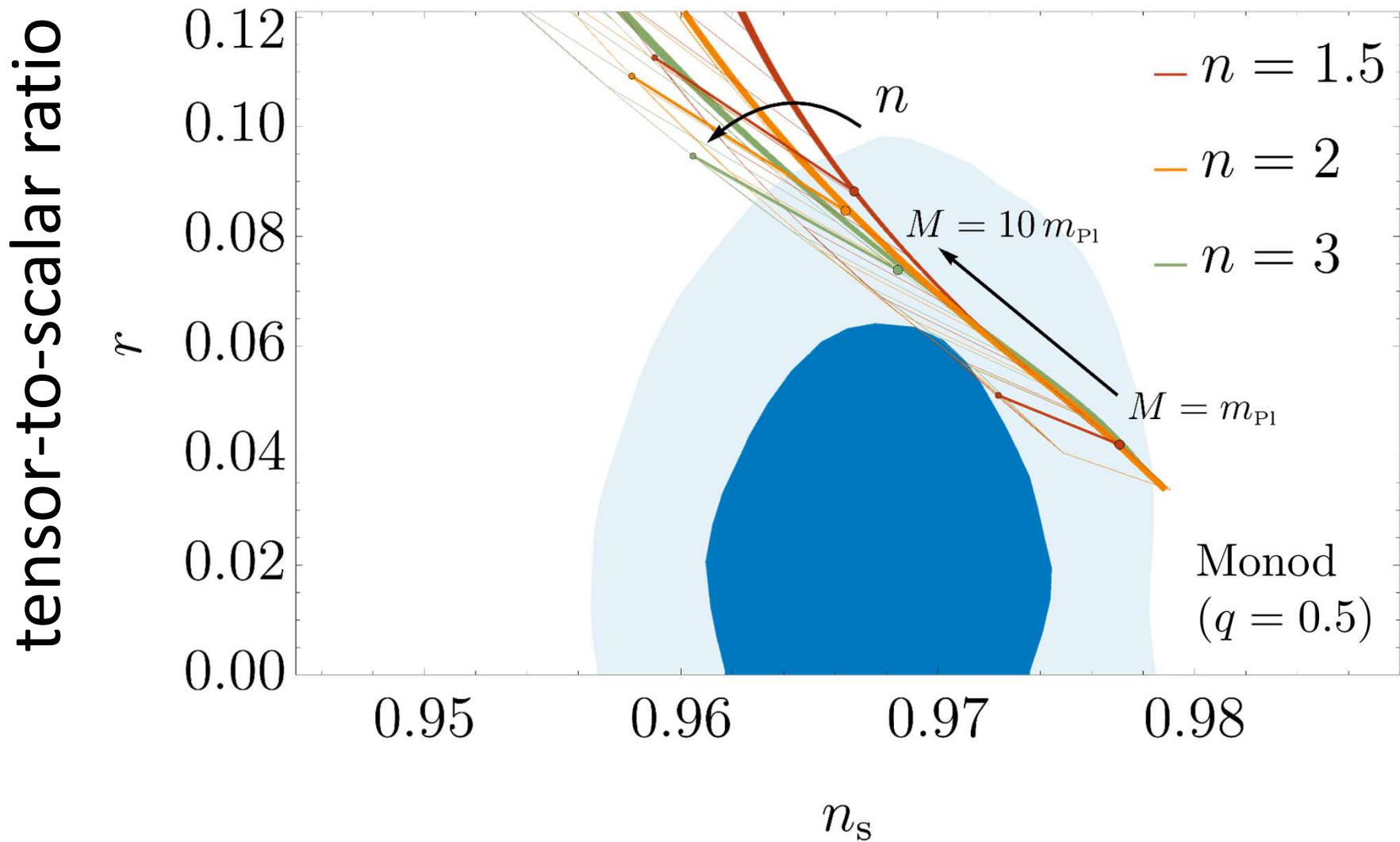


Expansion history effects

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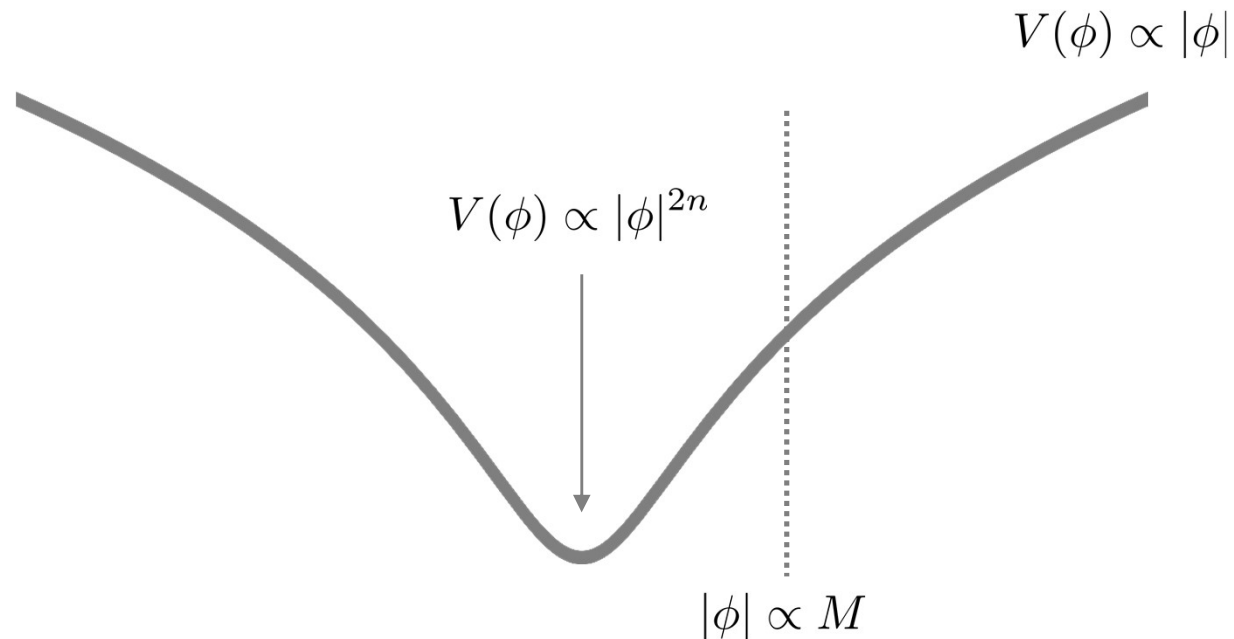


Expansion history effects

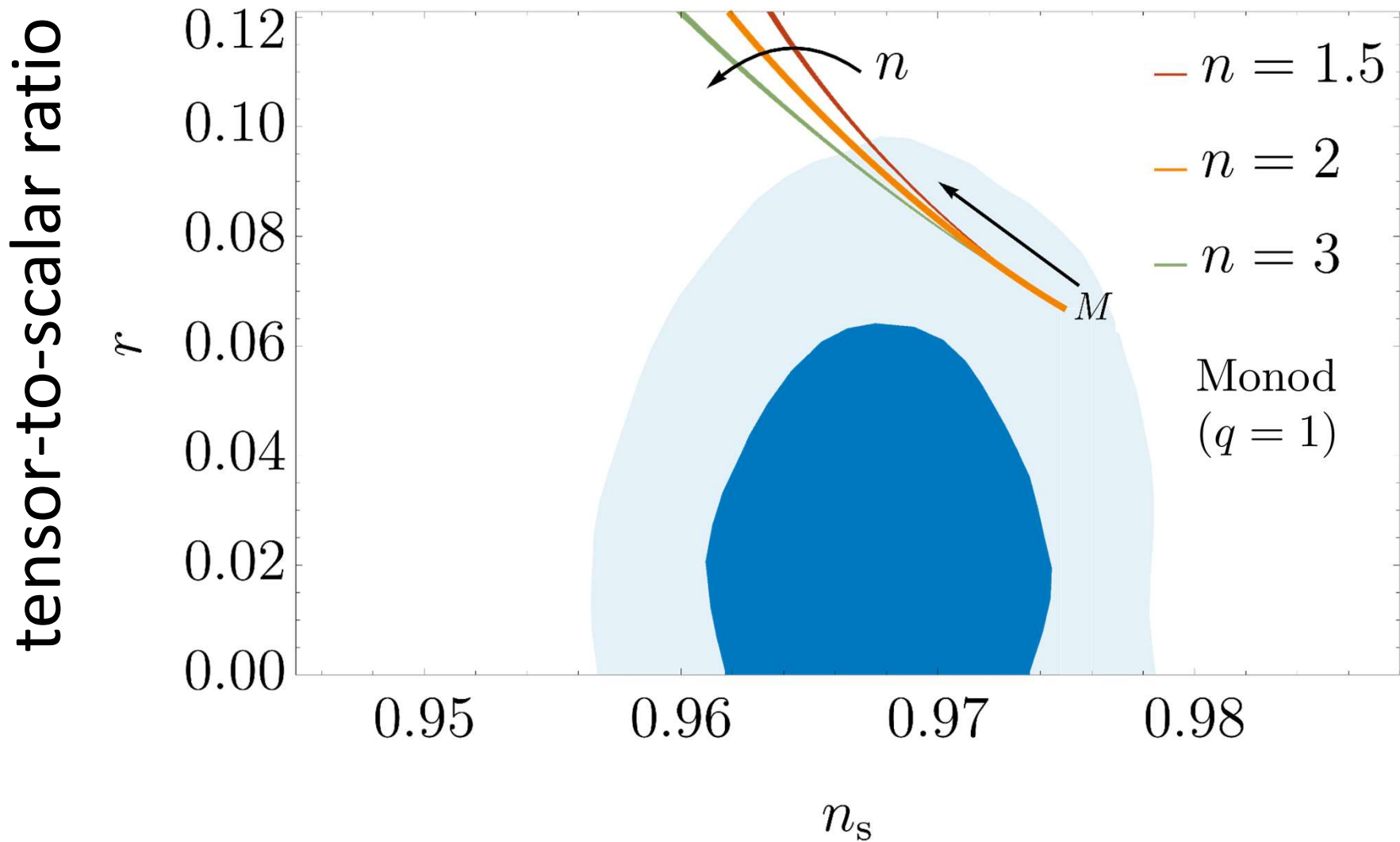


Expansion history effects

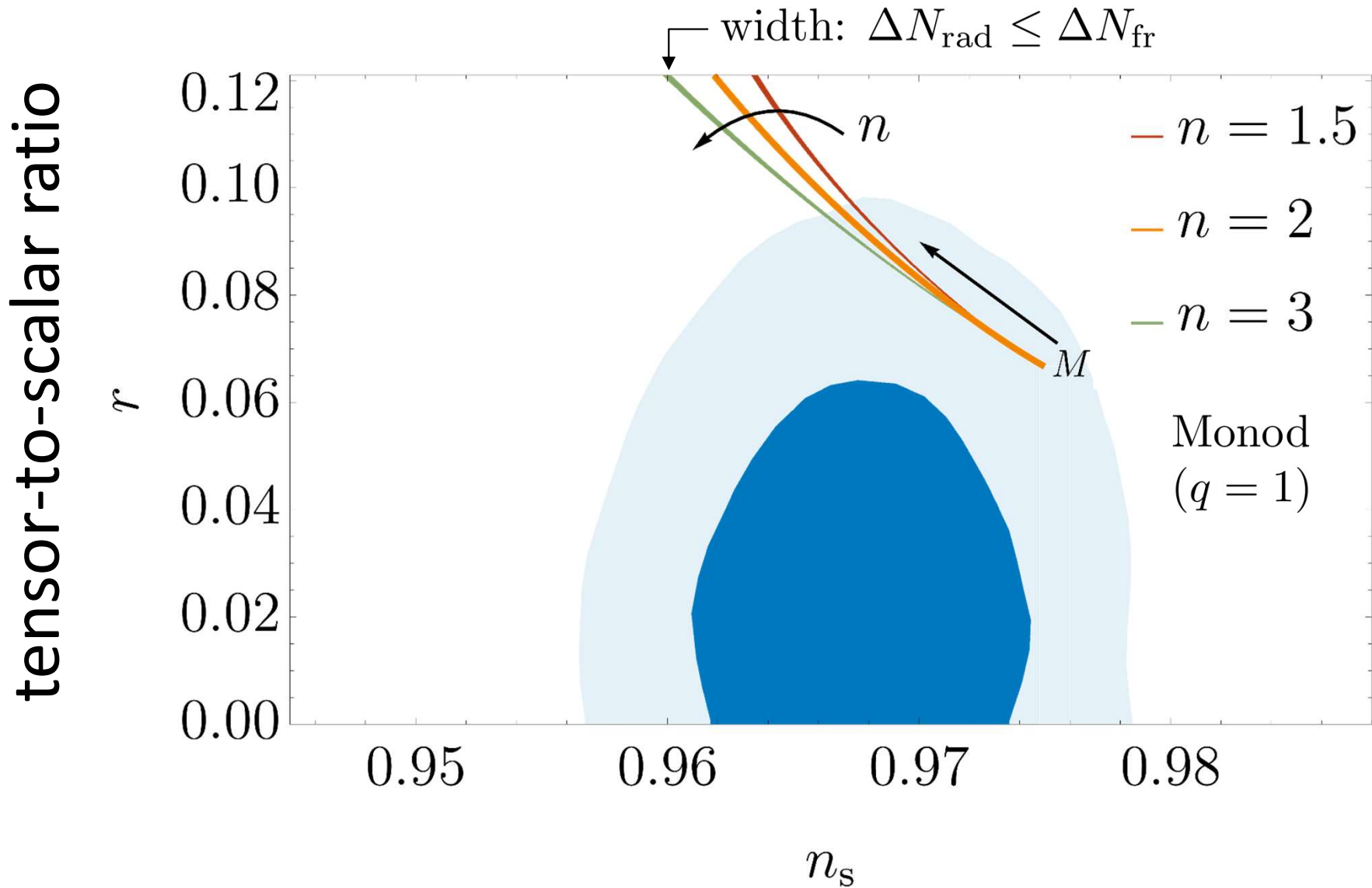
$$V(\phi) \propto \left[1 + \left| \frac{\phi}{M} \right|^{2n} \right]^{\frac{1}{2n}} - 1 \quad \text{Monodromy} \\ q = 1$$



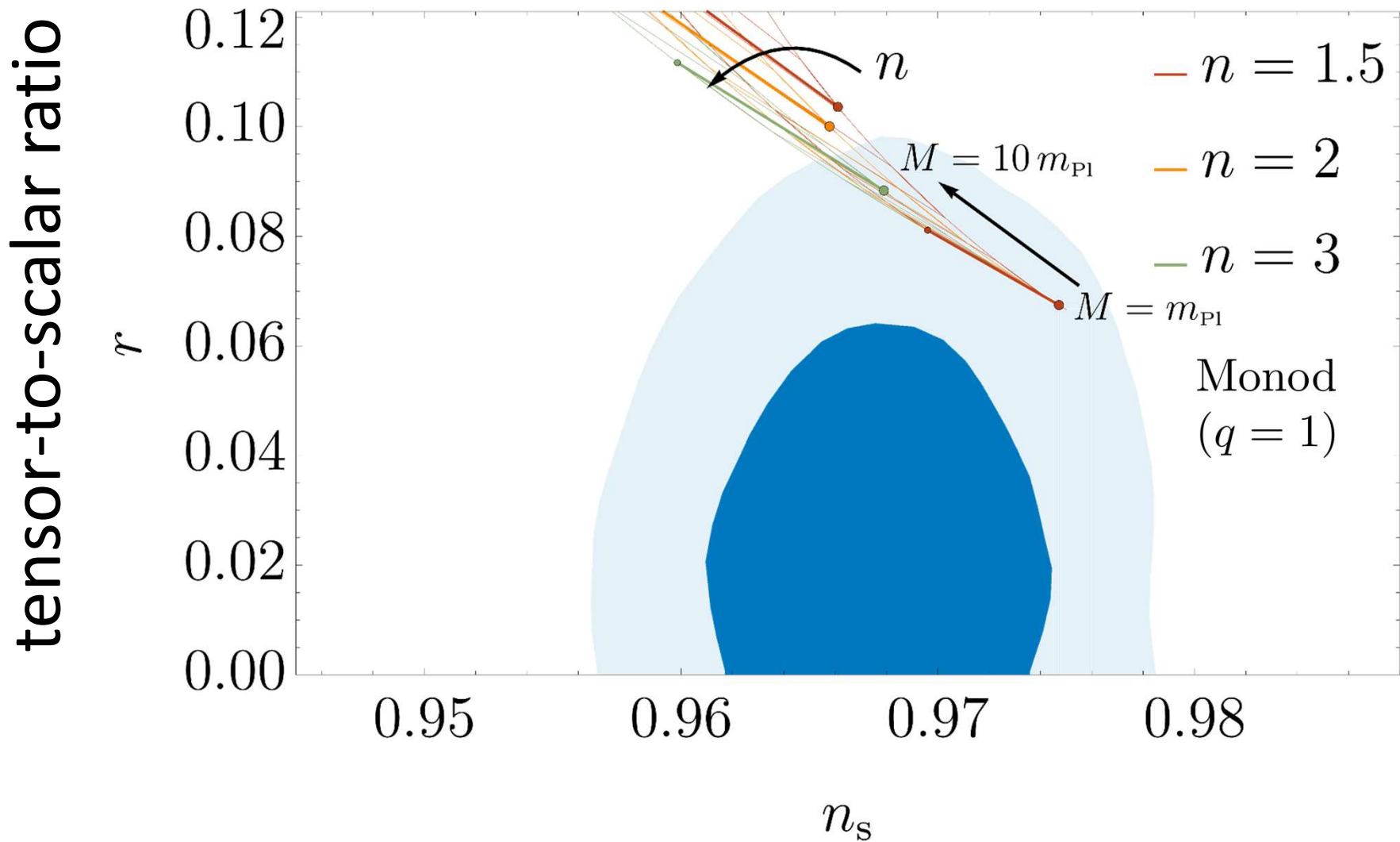
Expansion history effects



Expansion history effects

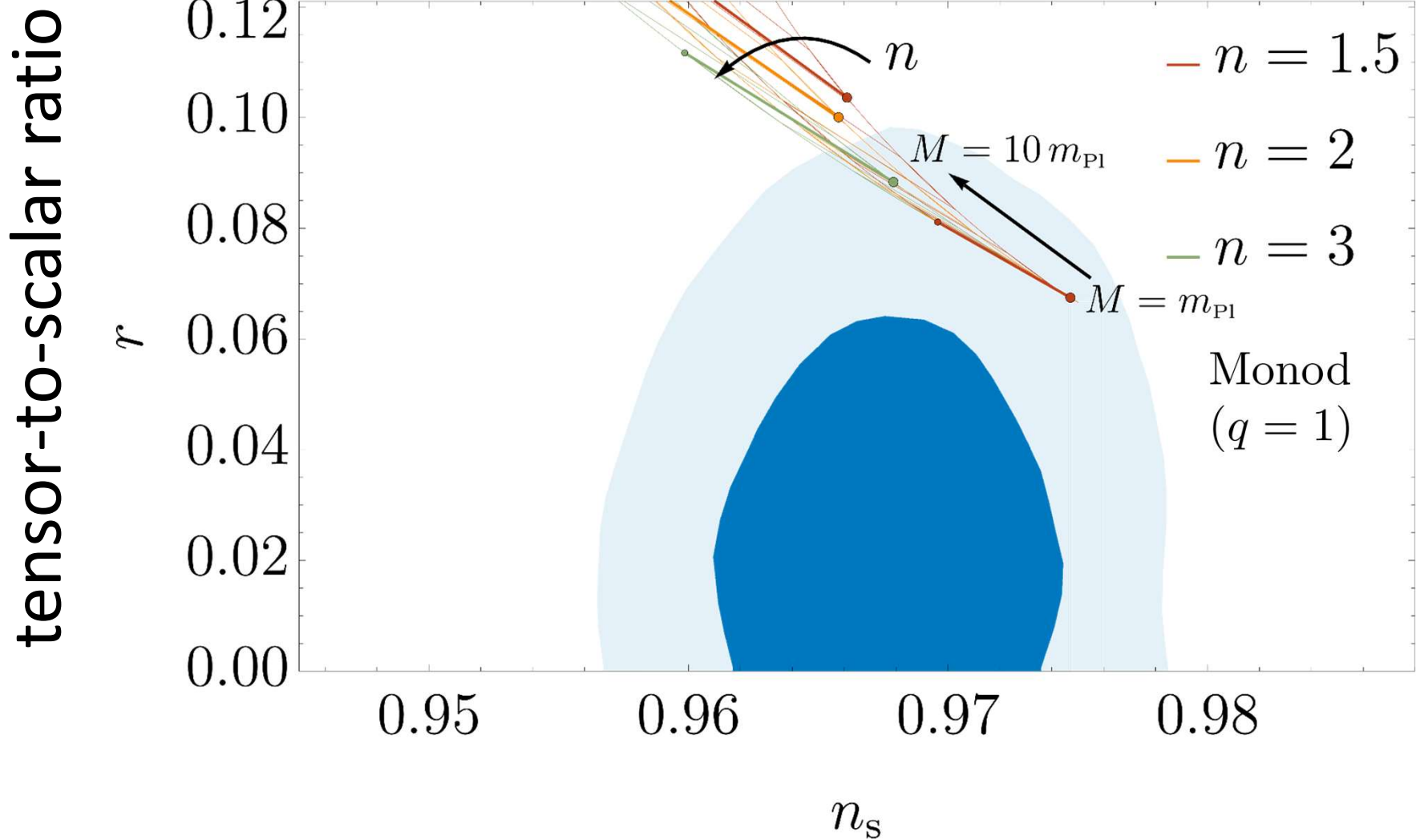


Expansion history effects

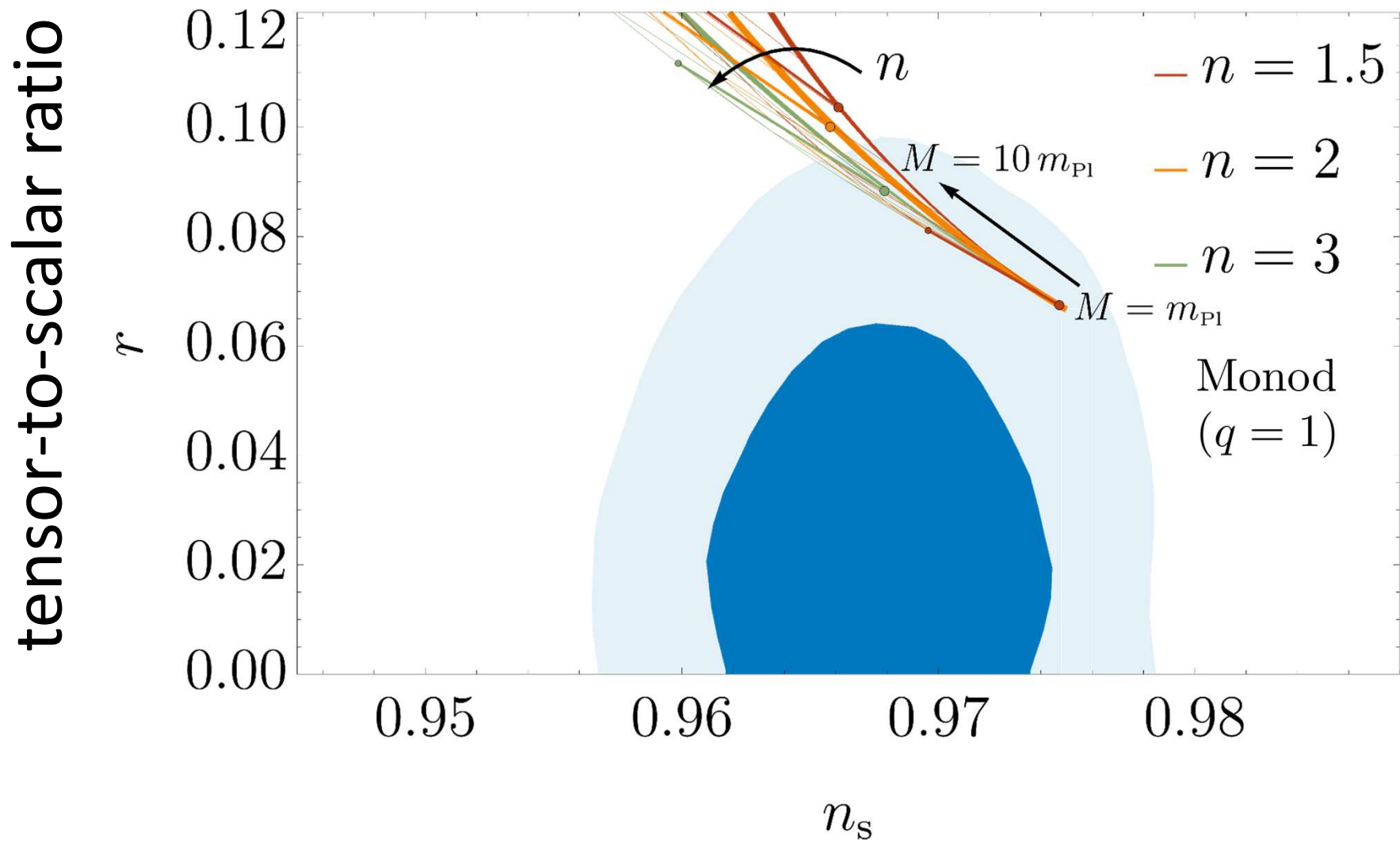


Expansion history effects

width: $50 \leq N_* \leq 60$



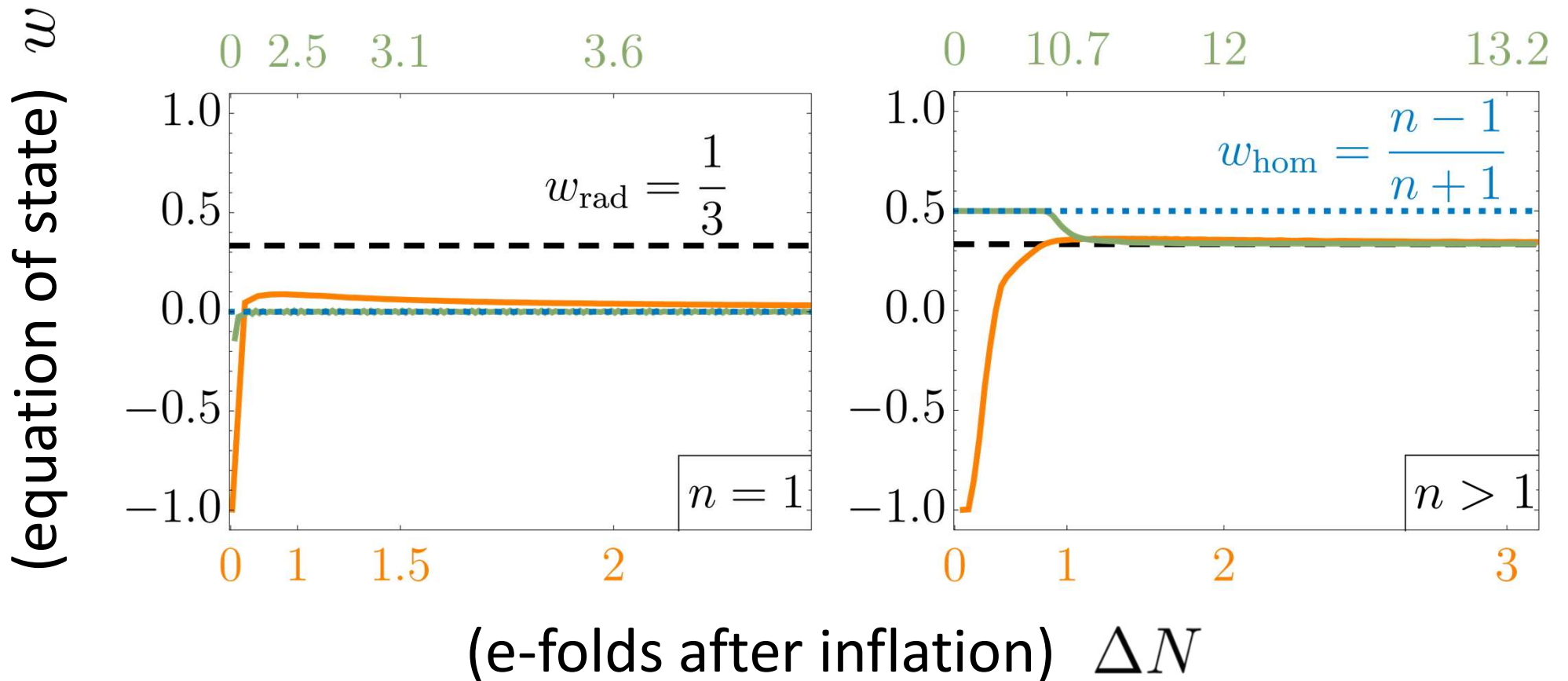
Expansion history effects



Equation of state

— $M \ll m_{\text{pl}}$ (efficient resonance)

— $M \sim m_{\text{pl}}$ (inefficient resonance)



$$\Delta N \equiv \int_{a_{\text{end}}}^a d \ln a$$

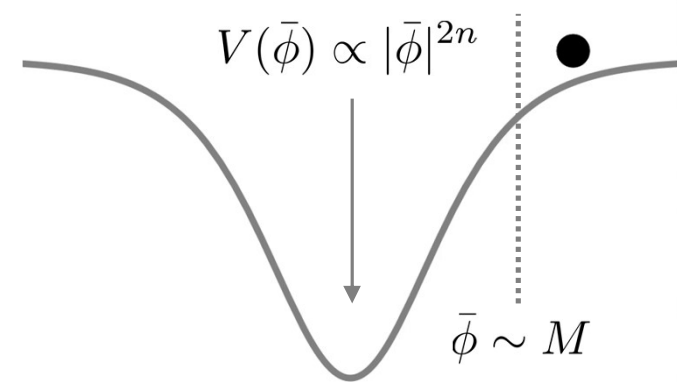
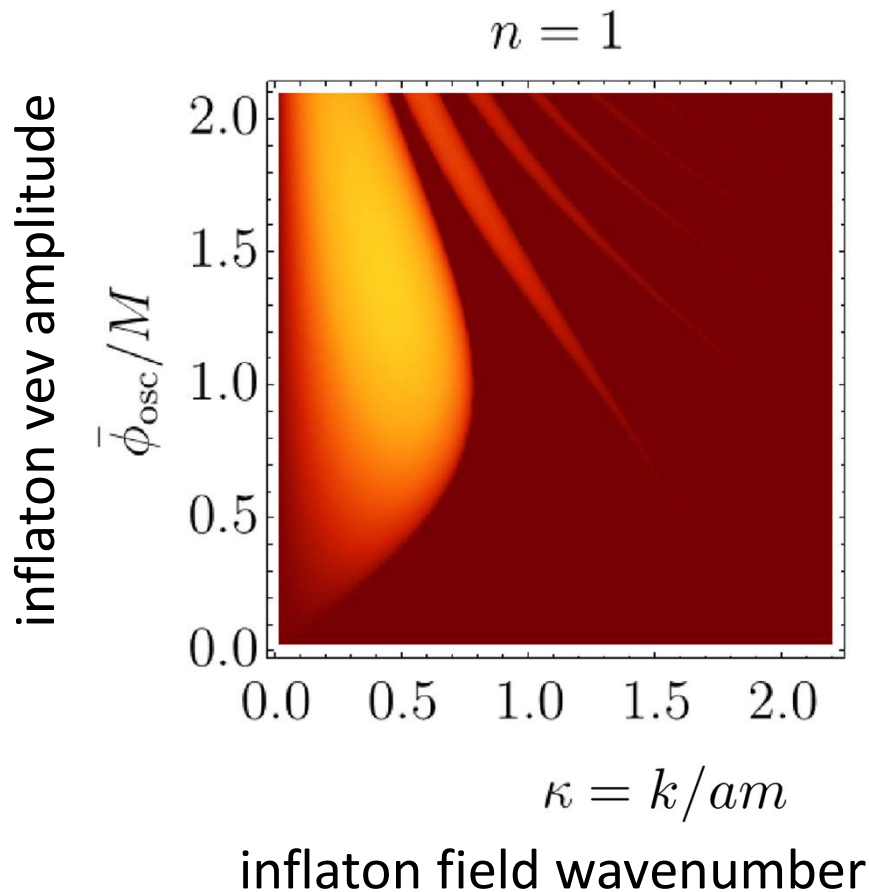
Matter domination?

$$n = 1$$

Matter domination?

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Non-perturbative decay (parametric self-resonance)



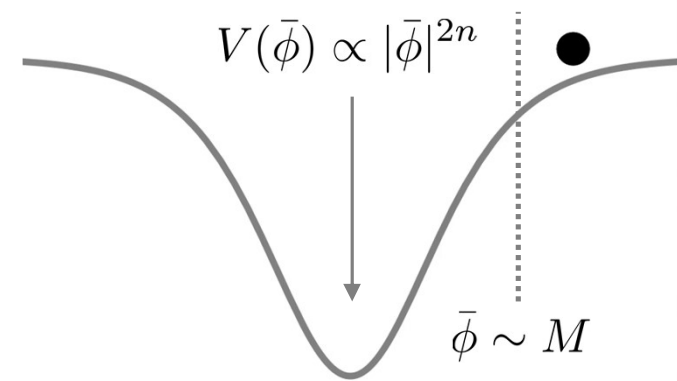
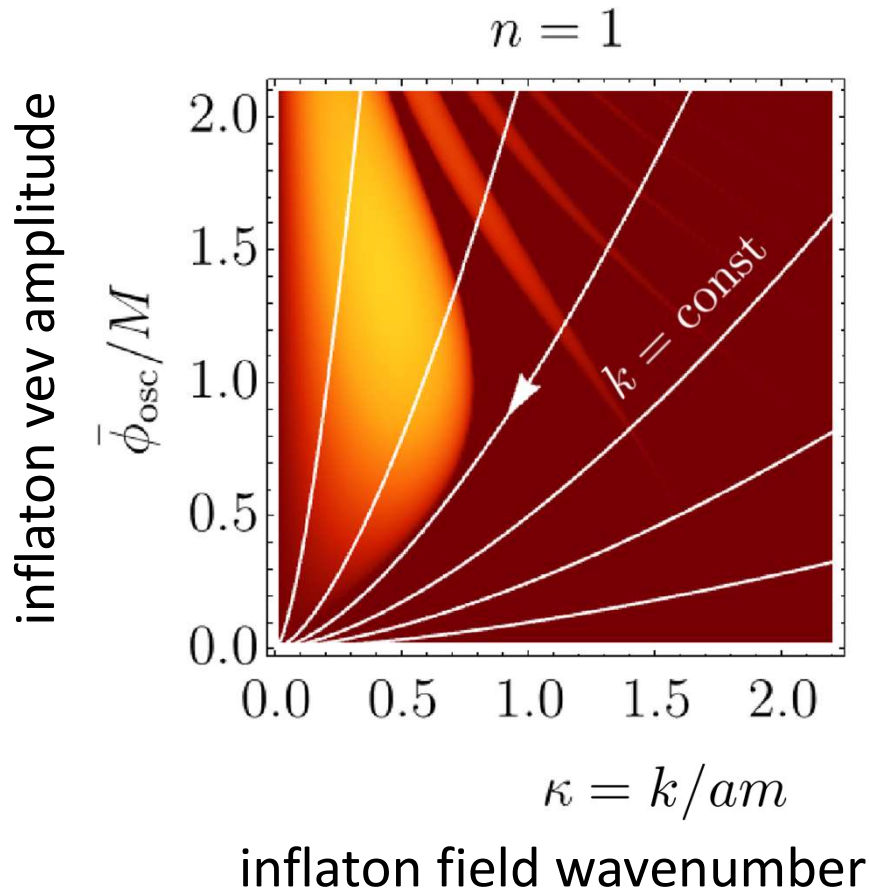
$$\delta\phi_k \propto \exp(\pm\mu_k t)$$

$$m^2 \equiv V'(\bar{\phi}_{\text{osc}})/\bar{\phi}_{\text{osc}}$$

Matter domination?

$$n = 1$$

Non-perturbative decay (parametric self-resonance)



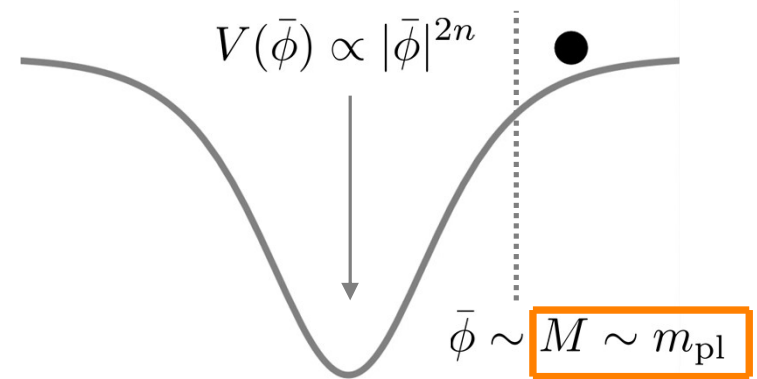
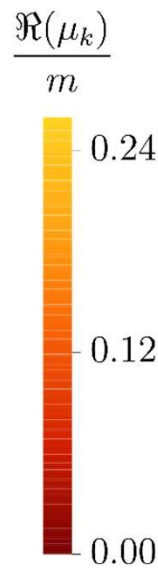
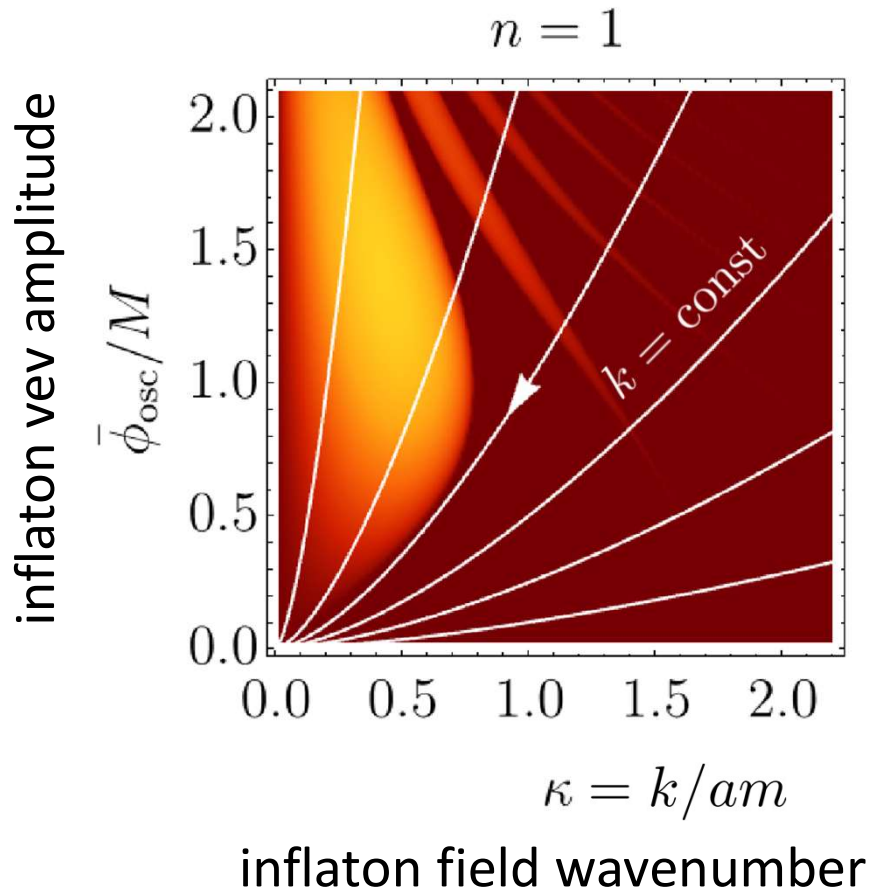
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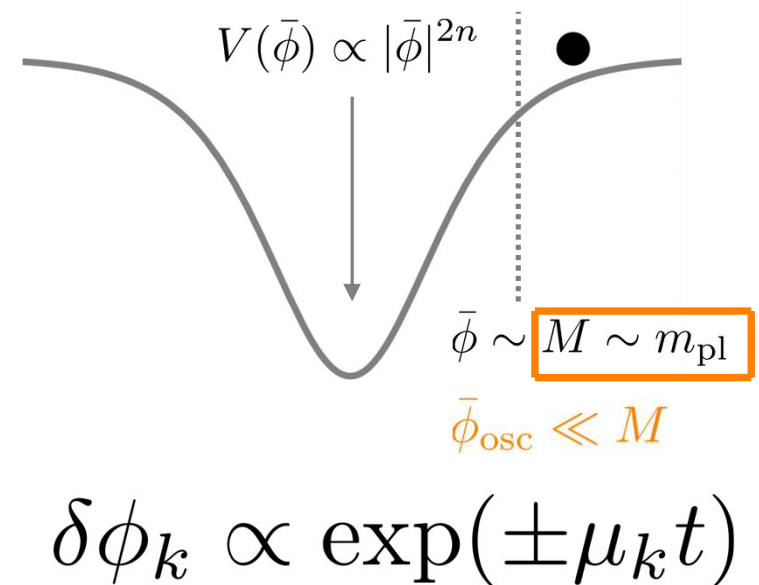
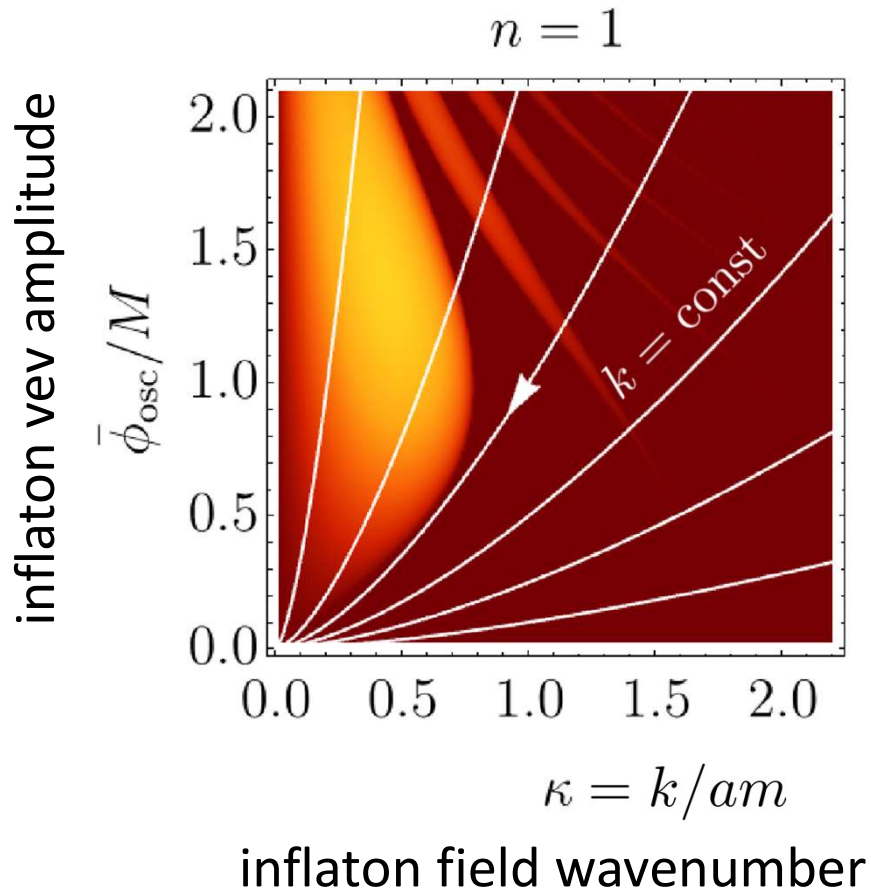
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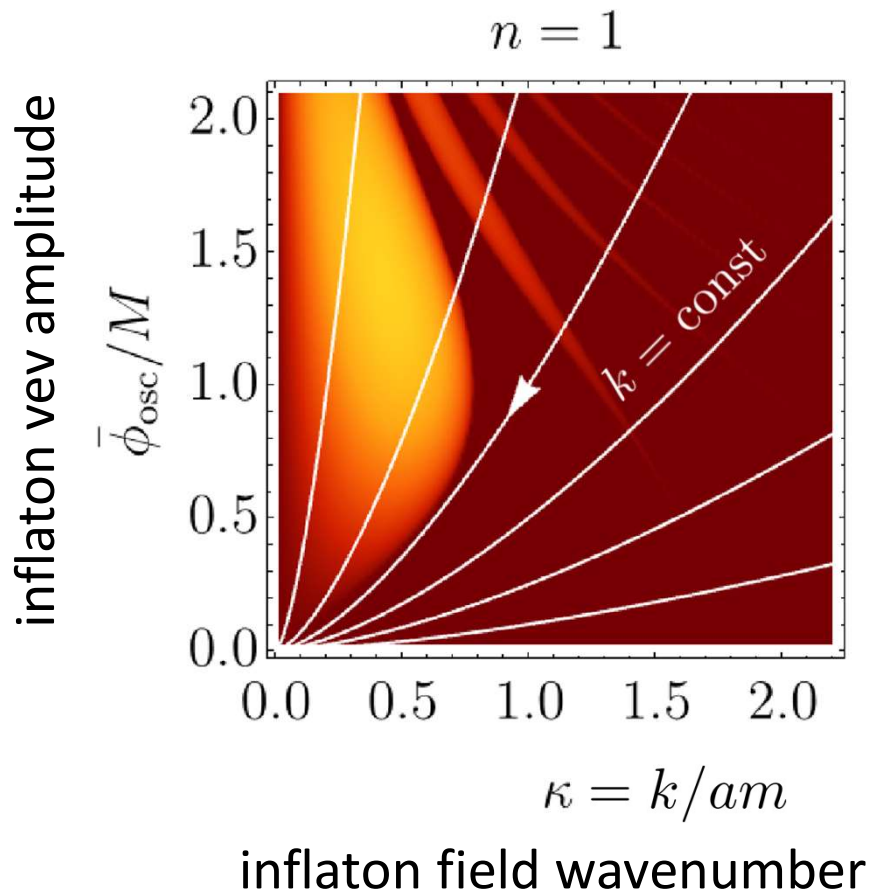


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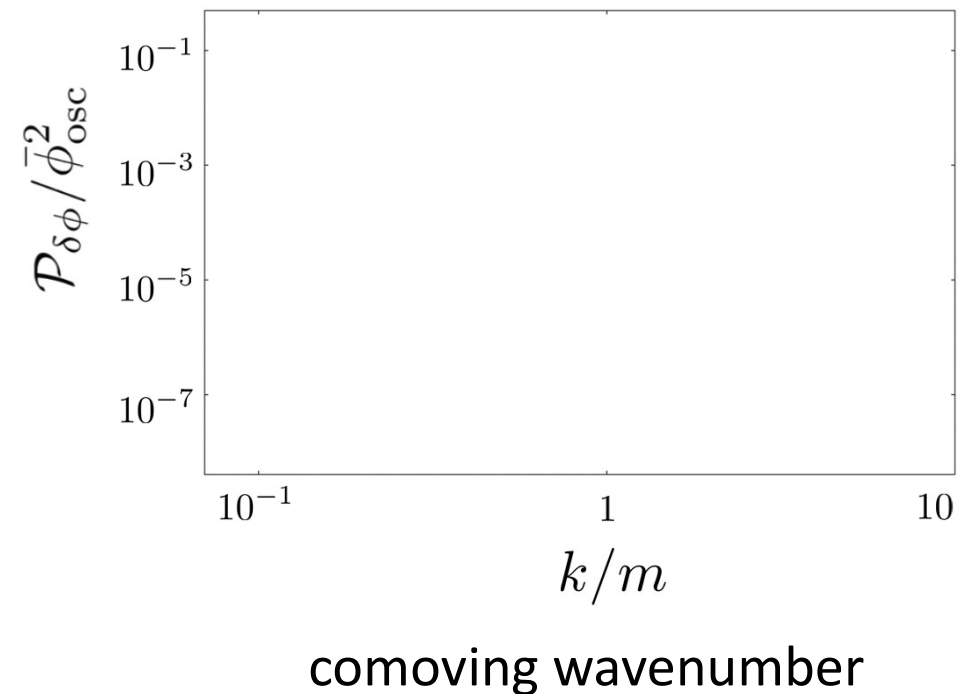
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Non-perturbative decay (parametric self-resonance)



Power spectrum:

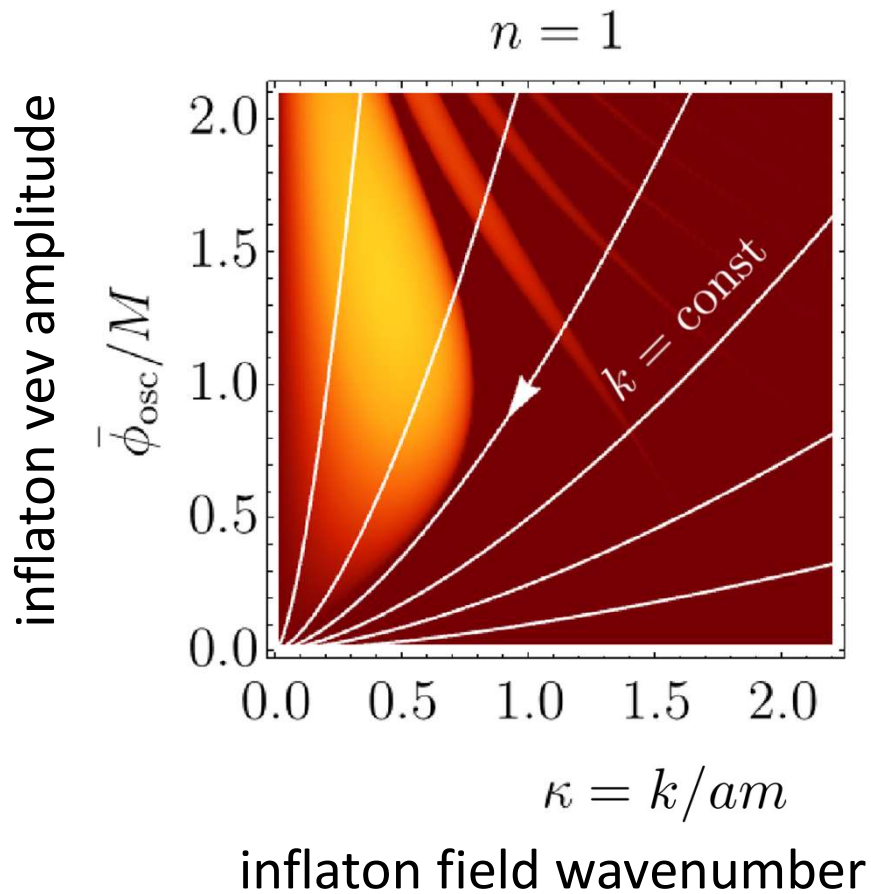
$$\langle \delta\phi(x)^2 \rangle \equiv \int \mathcal{P}_{\delta\phi} d\ln k$$



Matter domination?

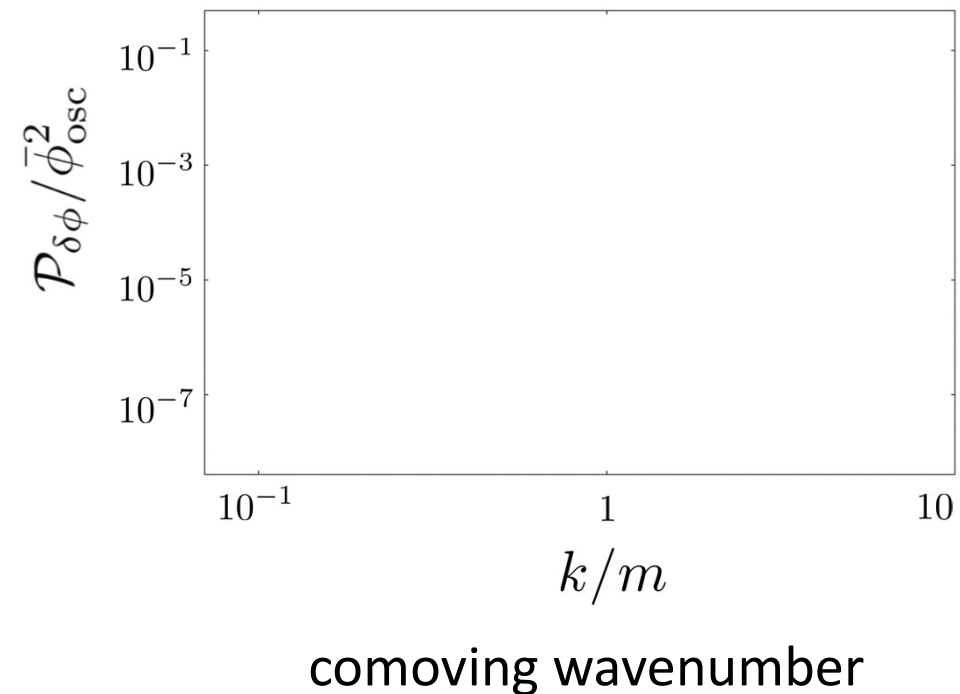
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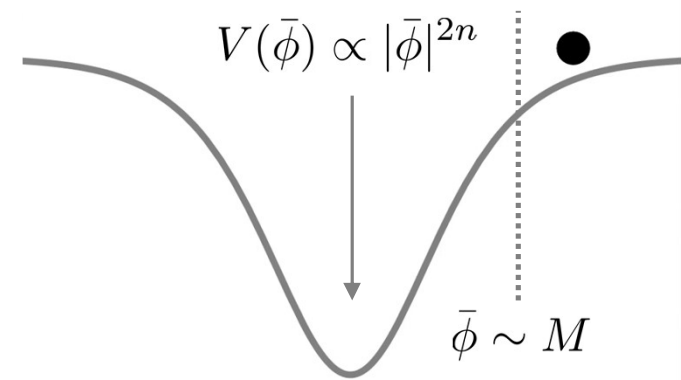
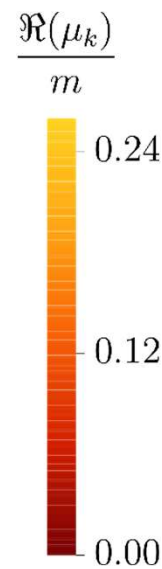
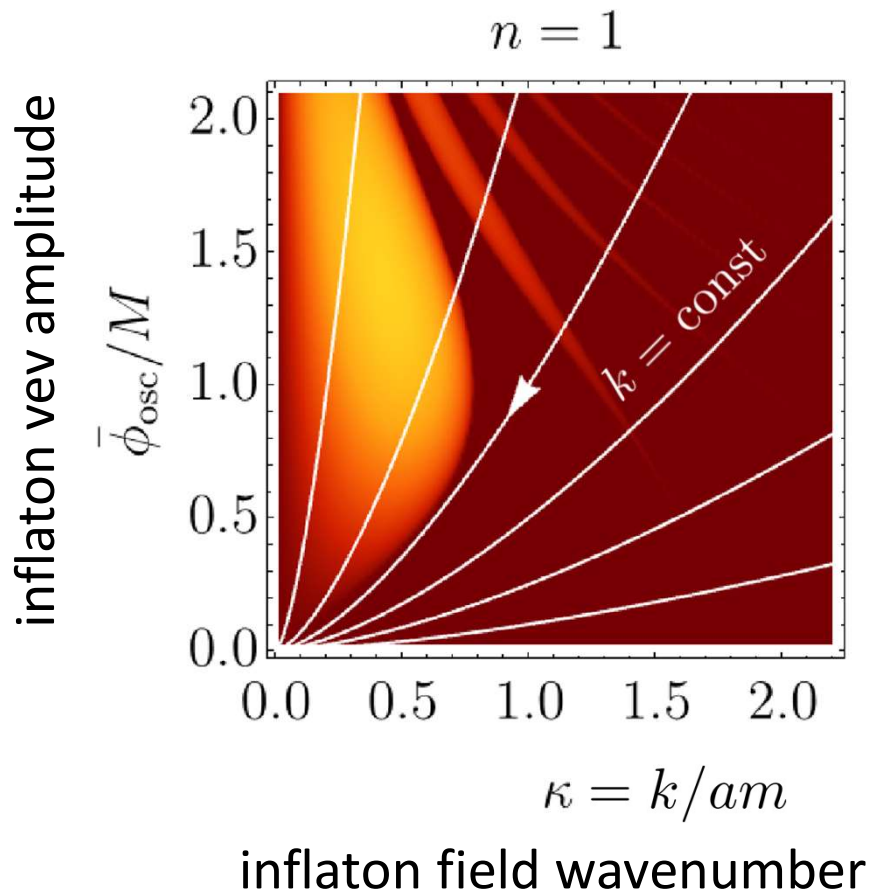
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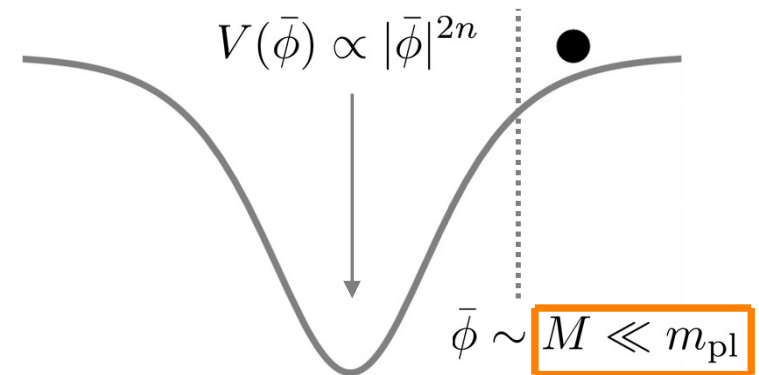
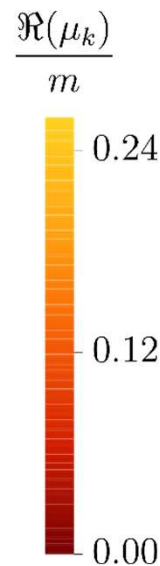
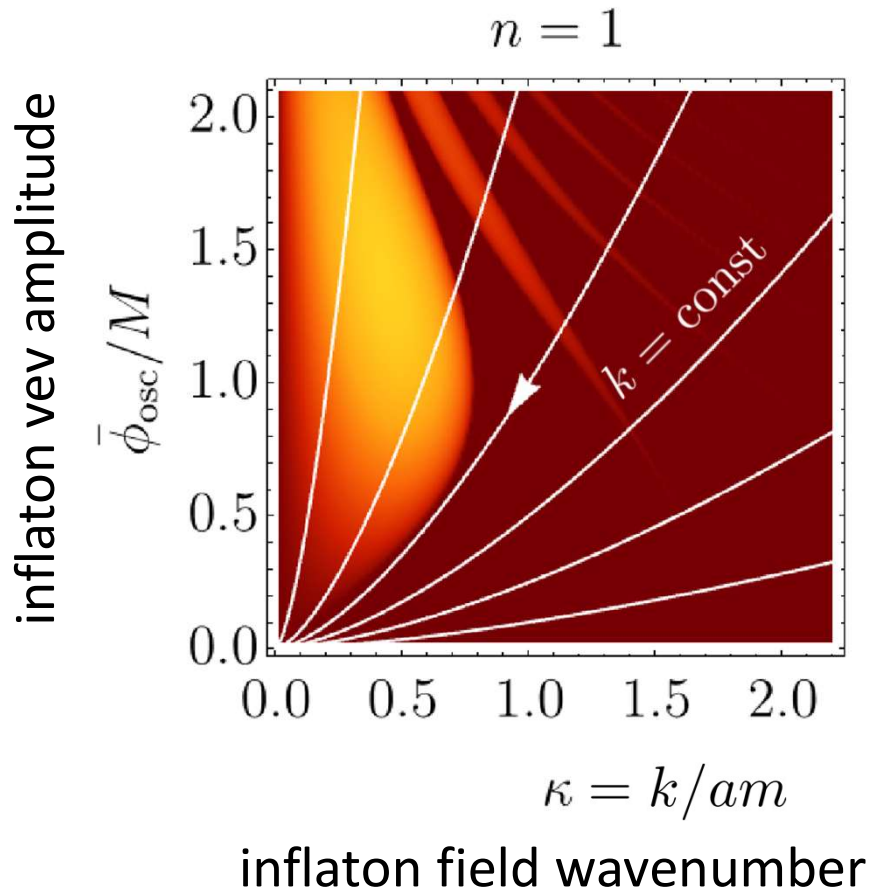
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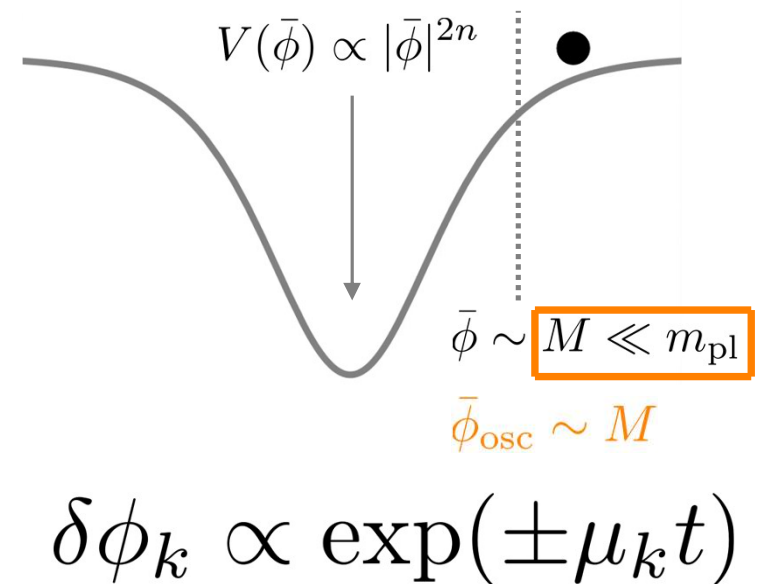
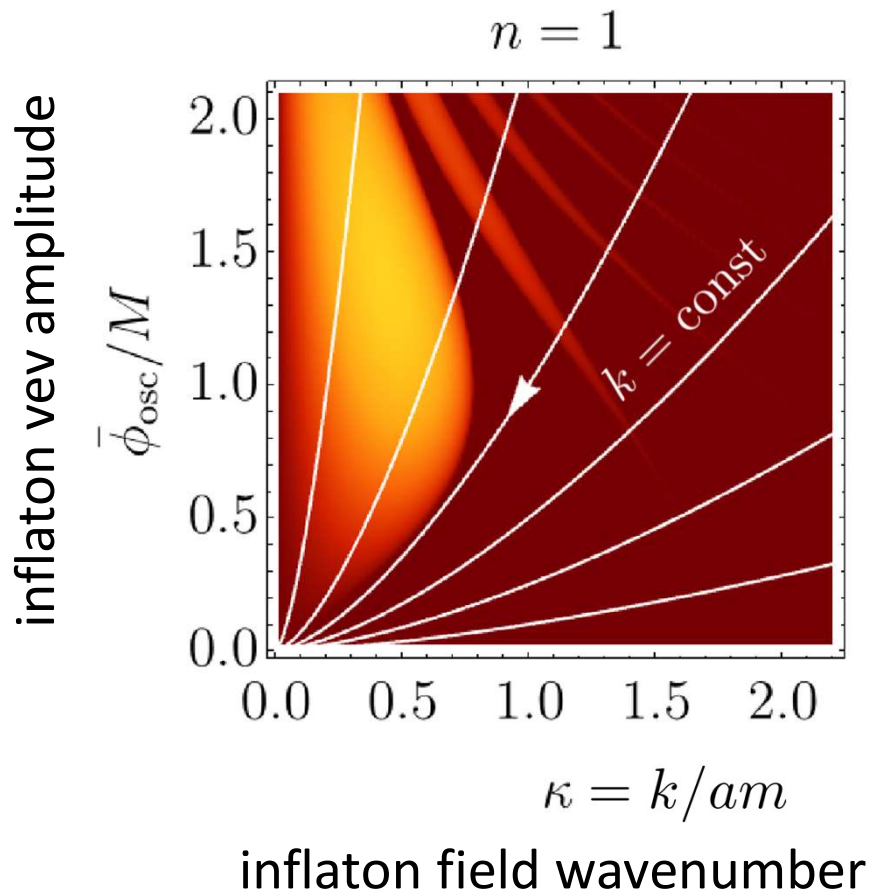
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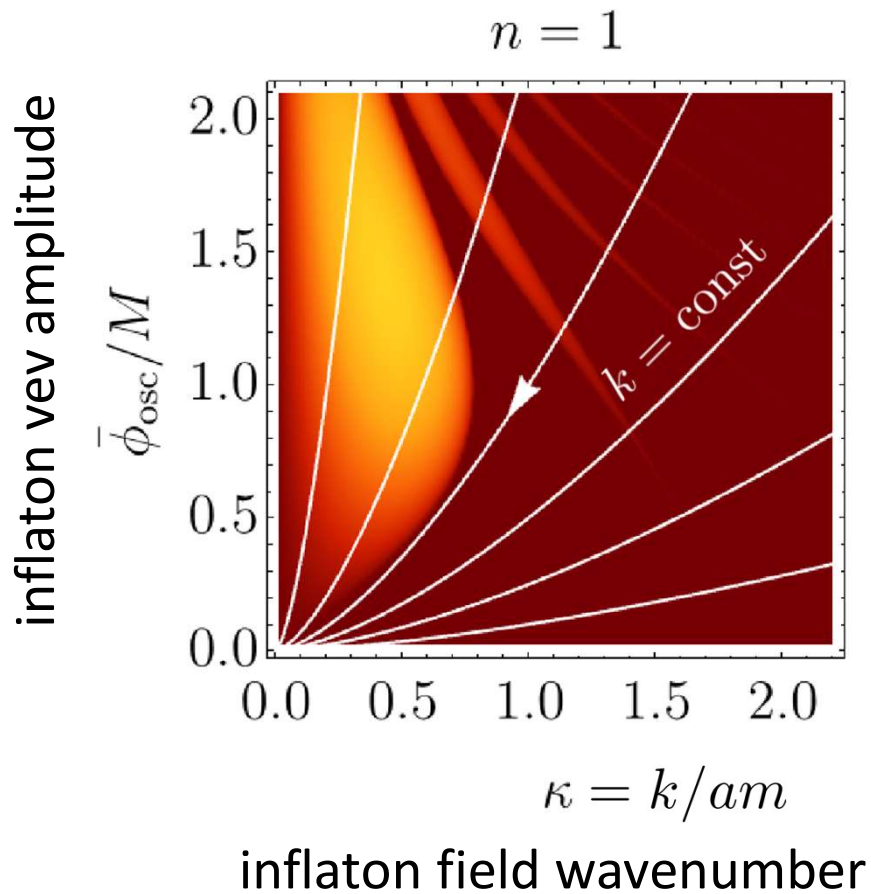


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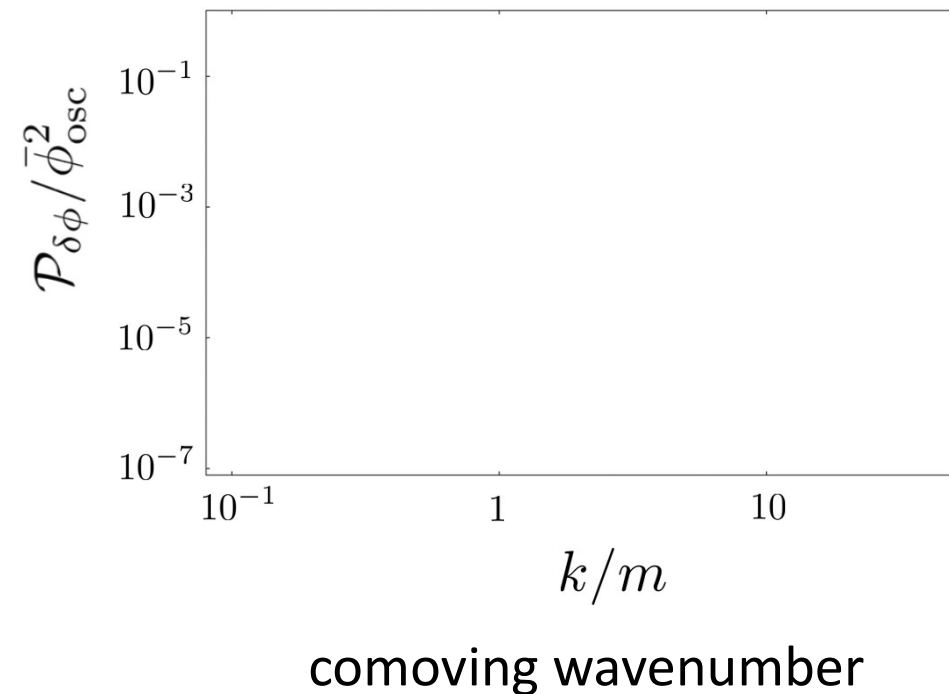
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Non-perturbative decay (parametric self-resonance)



Power spectrum:

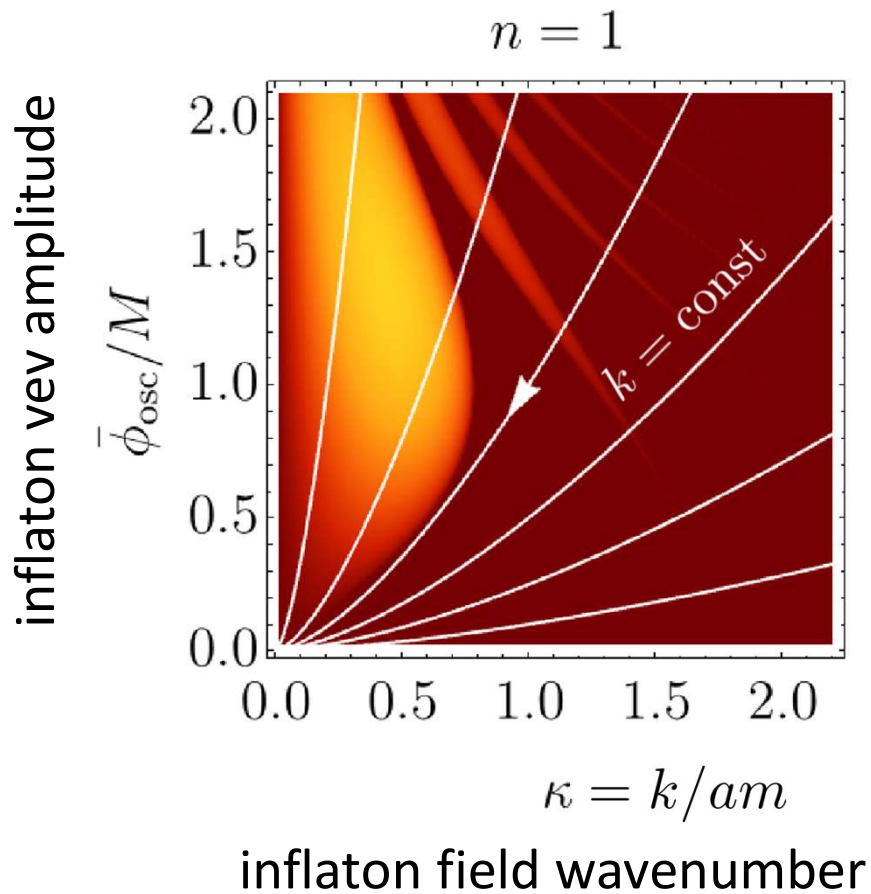
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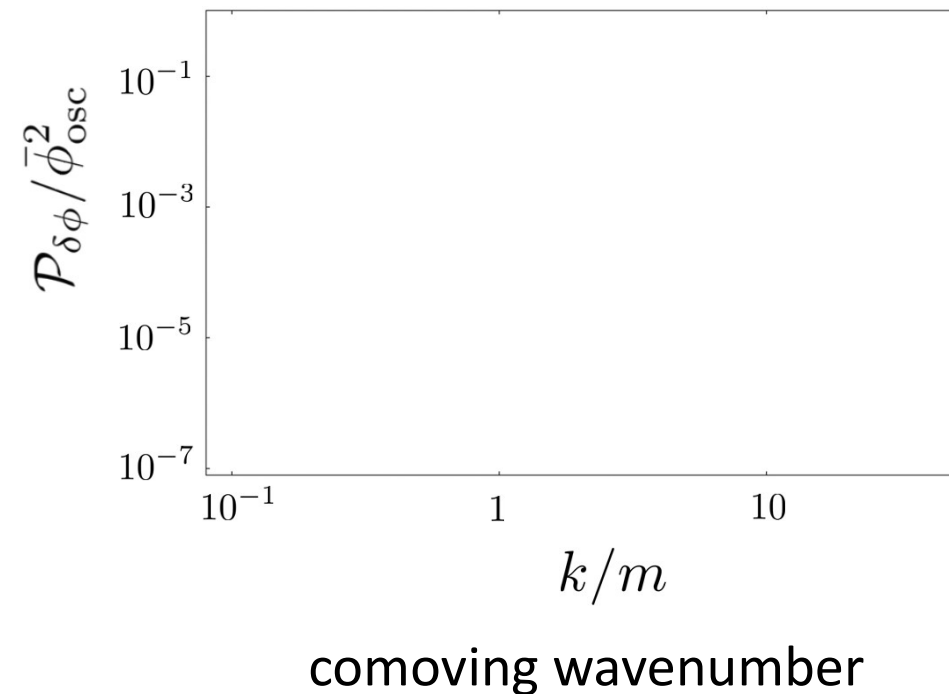
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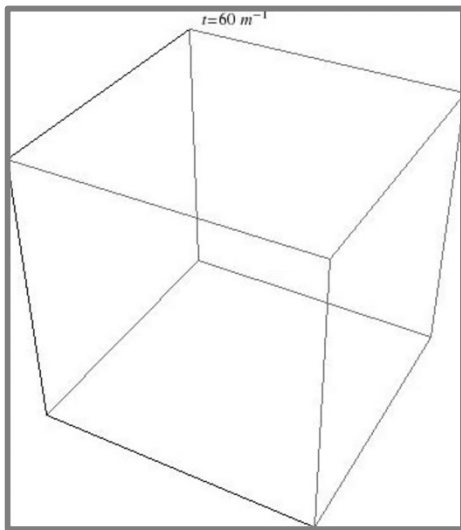


Matter domination?

$$n = 1$$

$$M \ll m_{\text{pl}}$$

$$M \sim m_{\text{pl}}$$



- $\delta\phi(t, \mathbf{x})$ production shut off
- $\bar{\phi}_{\text{osc}}(t) = \text{pressureless dust}$

- $\bar{\phi}$ forms oscillons (stable)

KL and Amin (2017, 18, 19)

See also Amin, Easter, Finkel, Flauger, Hertzberg (2011)

matter-like eos: $w = 0$

couplings to other fields?

Other connections ...

Other connections ...

- stochastic GWs from fragmentation

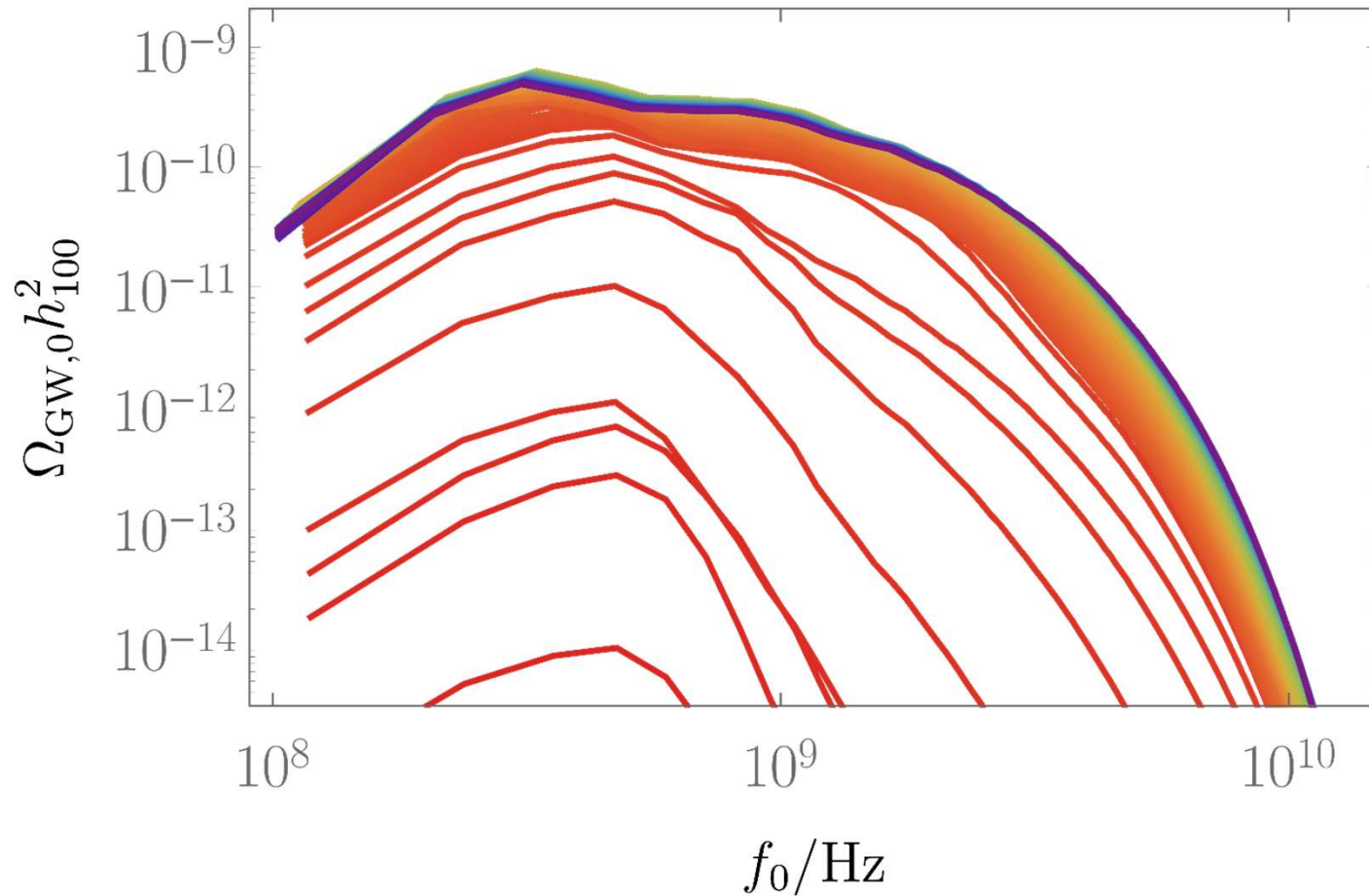
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Oscillons

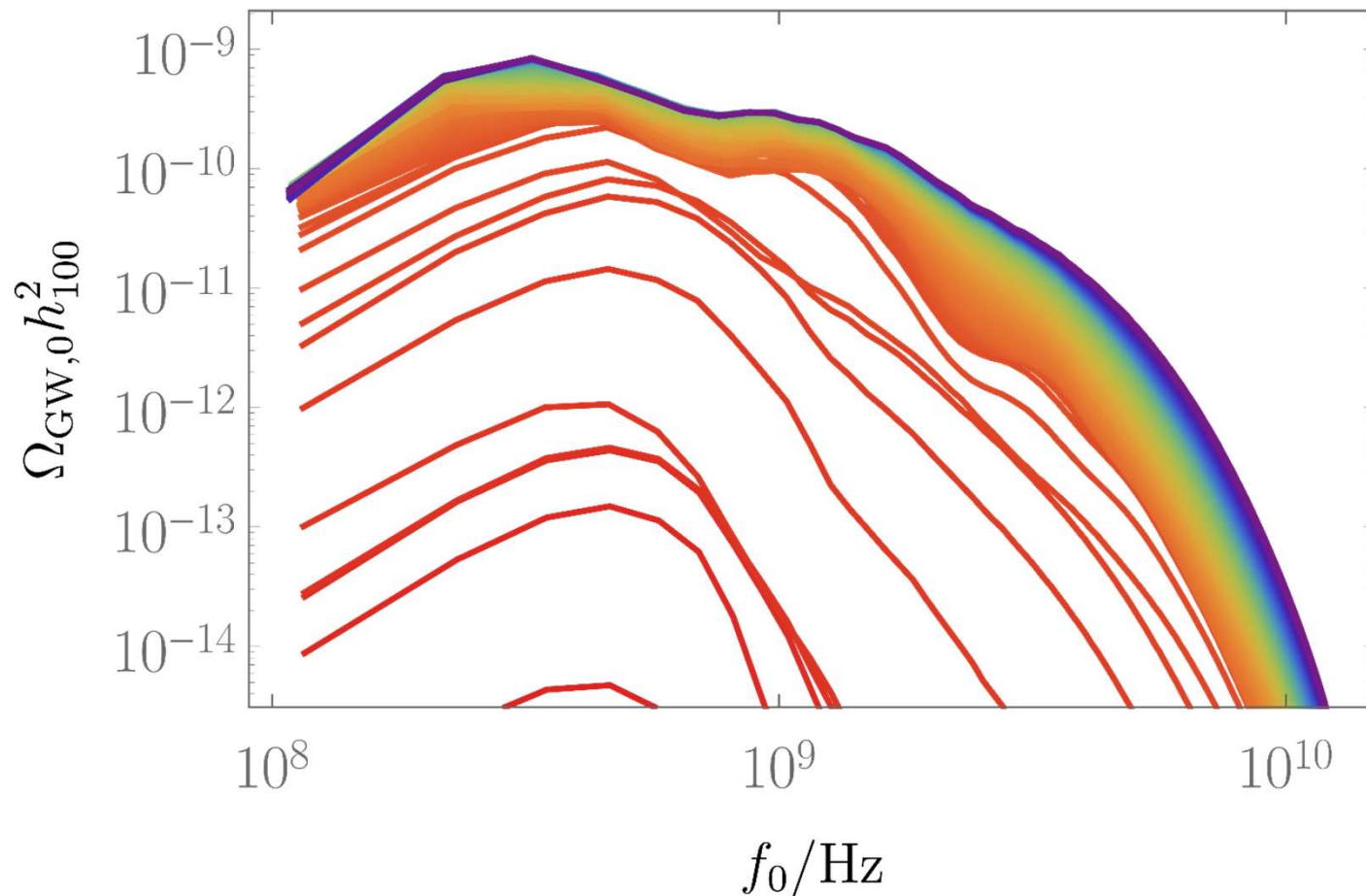


Other connections ...

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KL and M. Amin (2019)

Transients

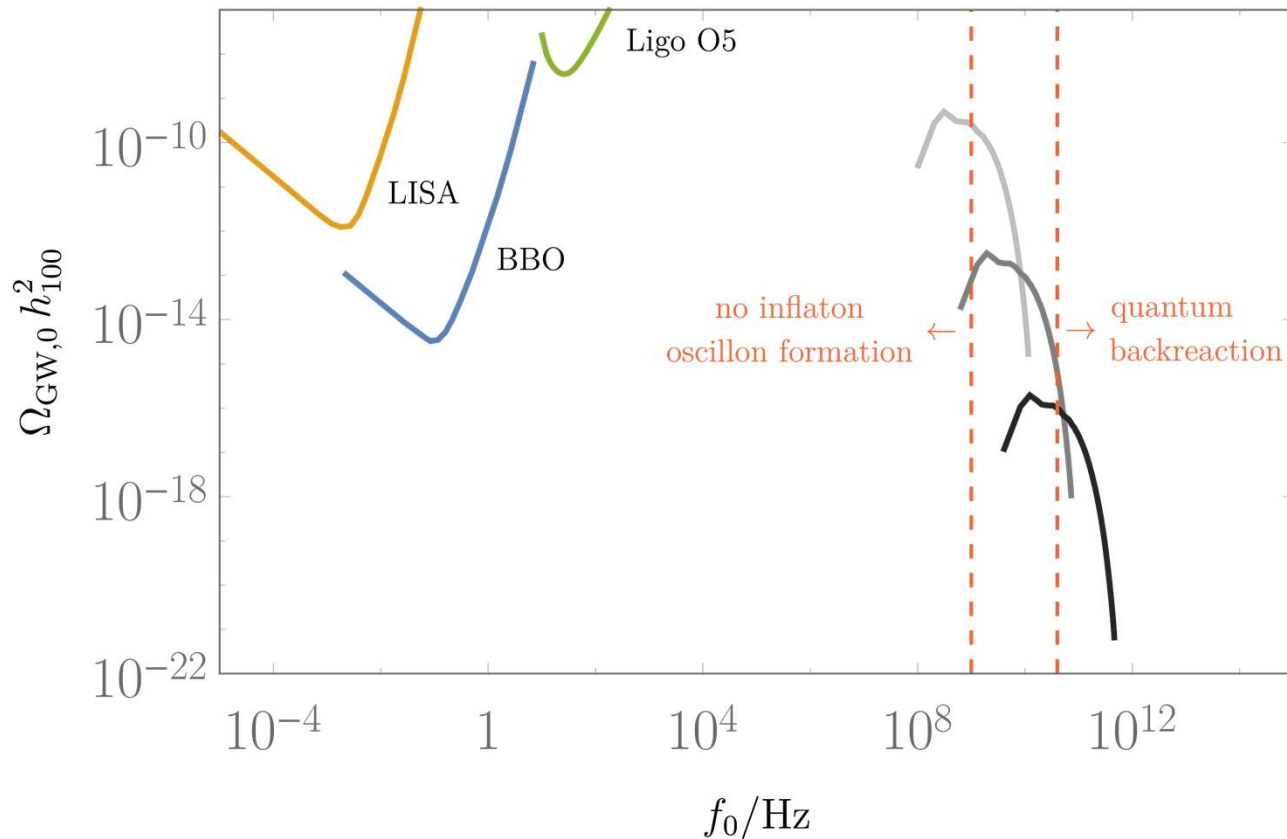


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INFLATON oscillons

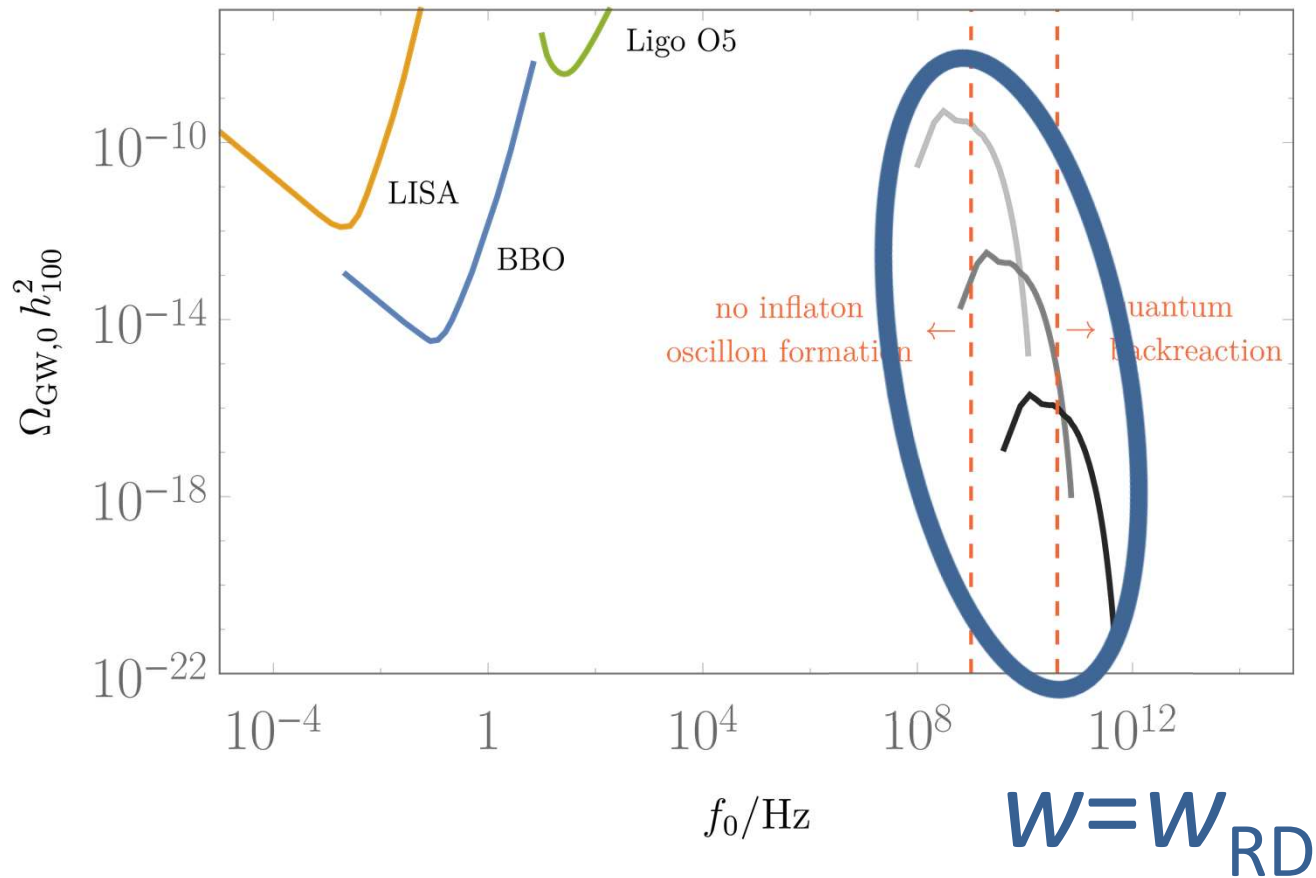


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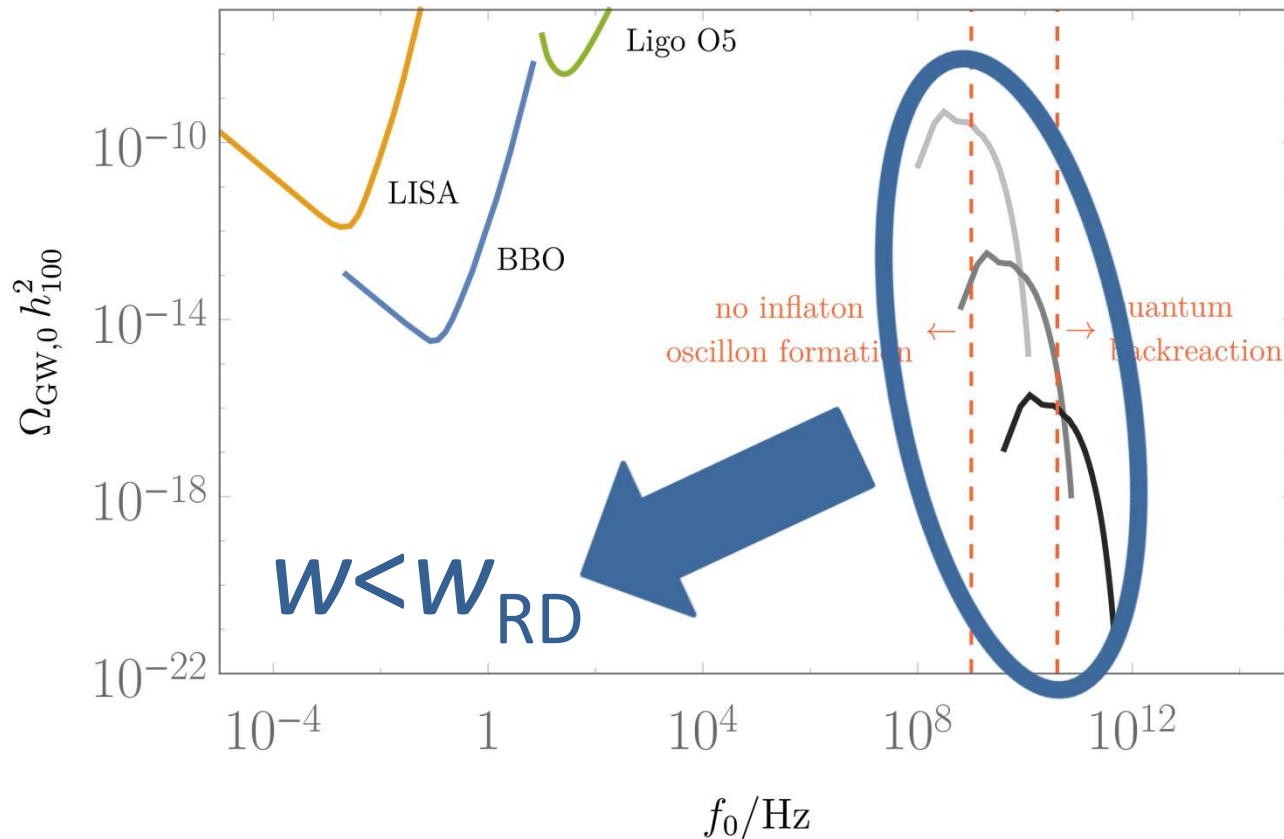


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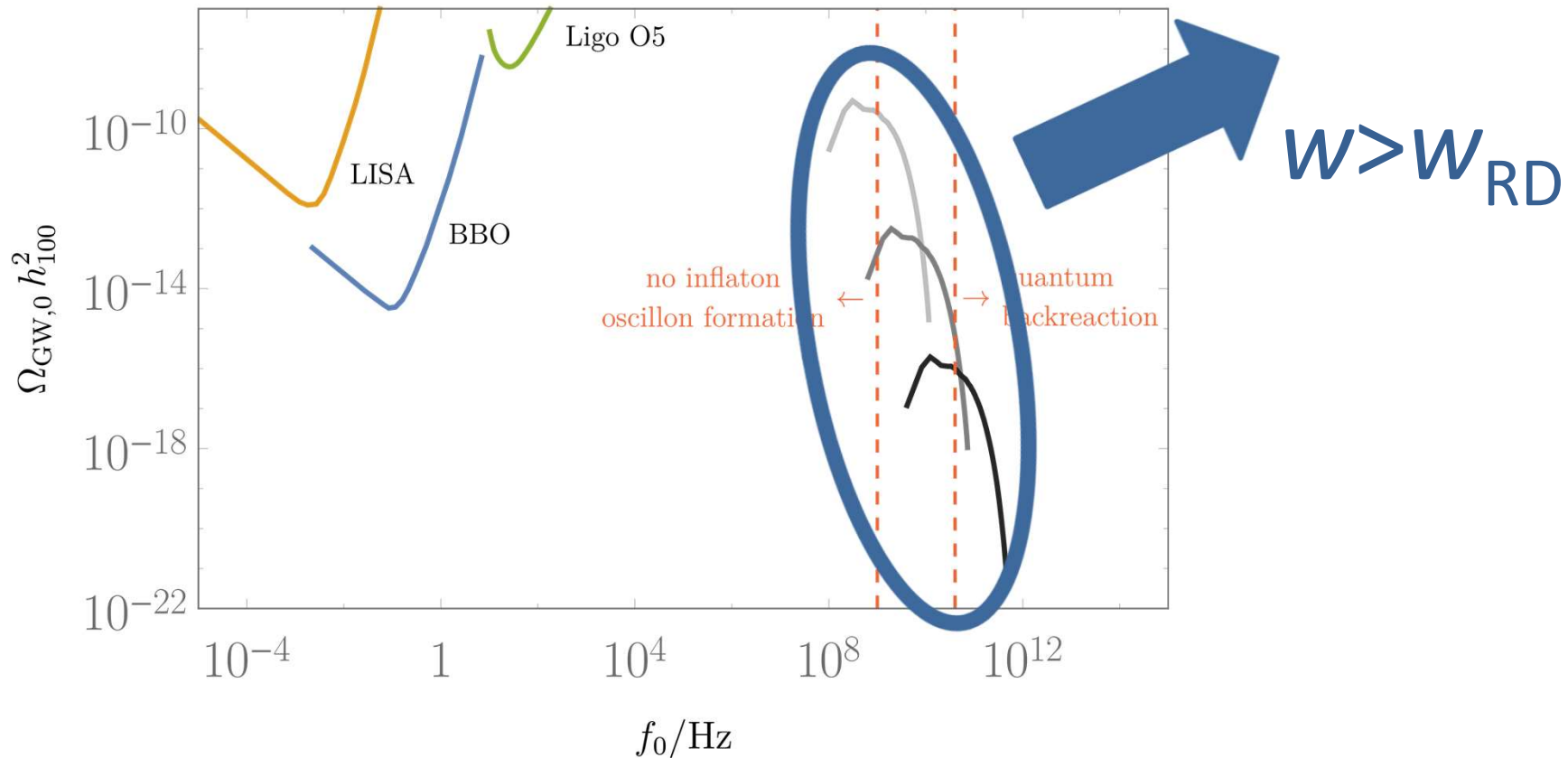


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M. Amin, KL and T. Smith (in progress)

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- dark matter ($n=1$ oscillons \rightarrow Q-balls)
- matter-antimatter asymmetry... KL and M. Amin (2014)

Oscillons & baryogenesis

KL and M. Amin, PRD 90, 083528 (2014)

complex
inflaton ϕ
+ ~~$U(1)$~~



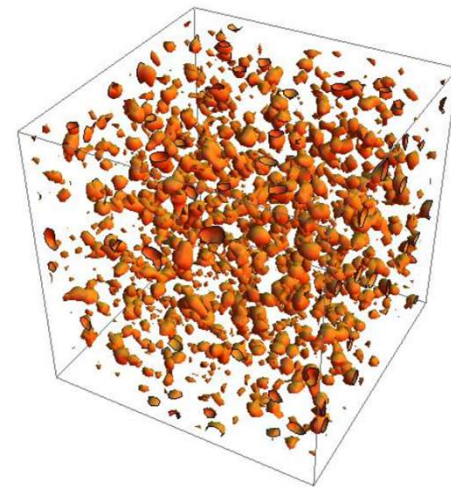
dynamics at the end of inflation

inflaton/anti-inflaton
asymmetry



decay

matter-antimatter
asymmetry
 $\eta \approx 6 \times 10^{-10}$



very different dynamics from homogeneous case!

Gauge fields, inflation & reheating

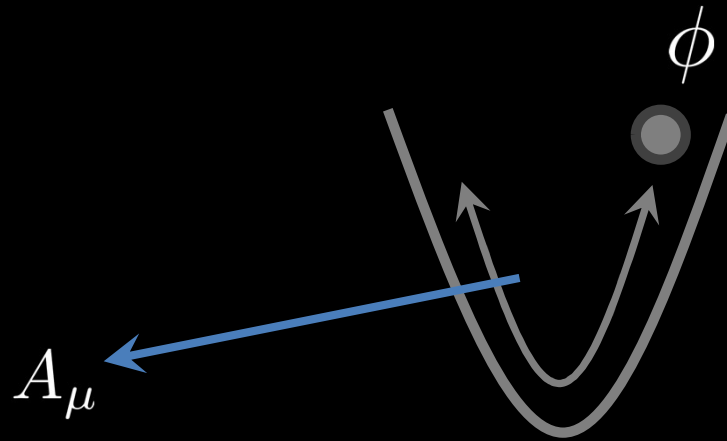
KL and M. Amin, JCAP 1606 032 (2016)

$$S = \int d^4x \sqrt{-g} \left[\frac{1}{16\pi G} \mathcal{R} - (D_\mu \phi)^\dagger D^\mu \phi - V(|\phi|) - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} \right]$$

Gauge fields, inflation & reheating

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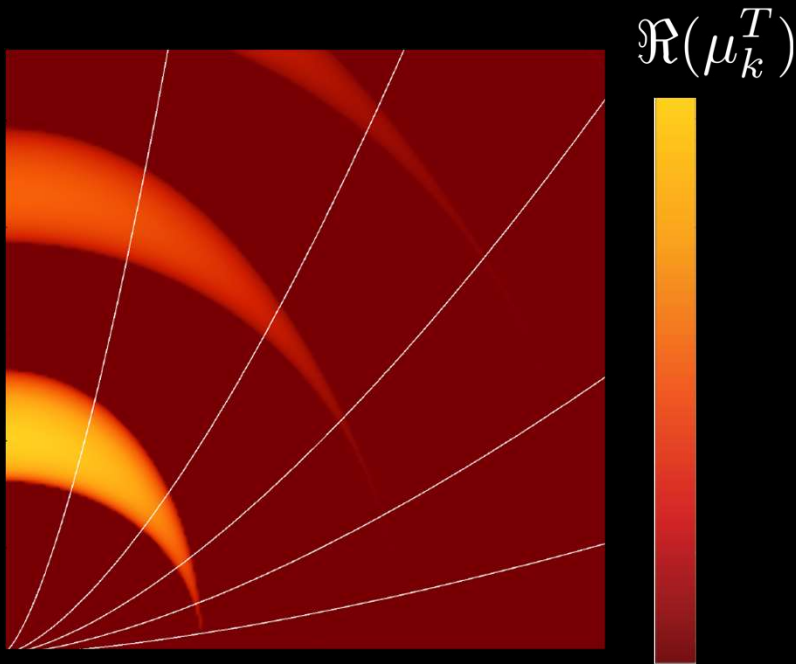
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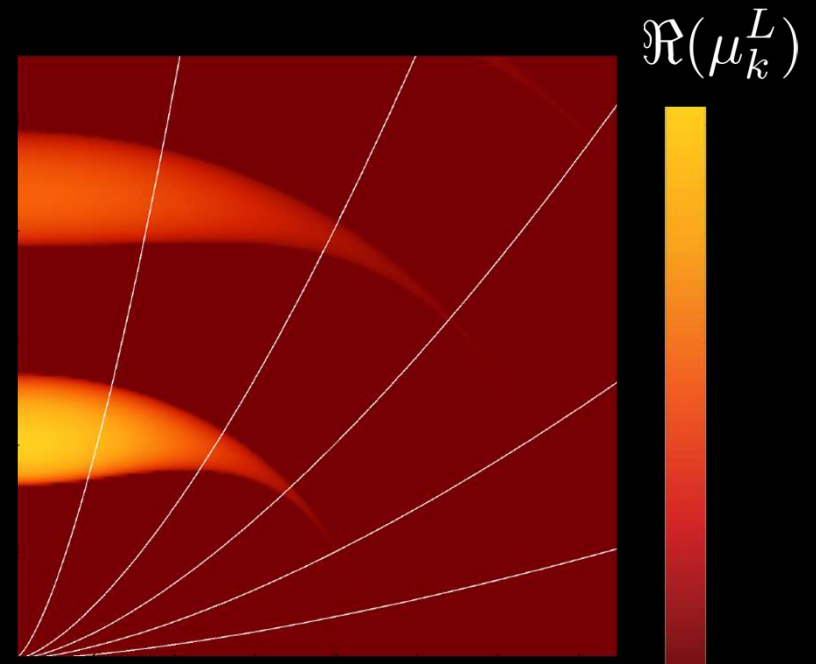
$$A_k^T \propto \exp(\pm \mu_k^T t)$$

$$A_k^L \propto \exp(\pm \mu_k^L t)$$

inflaton amplitude



wavenumber

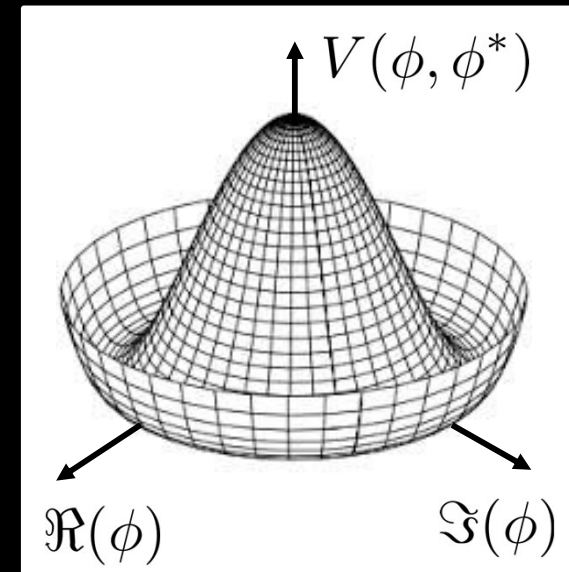
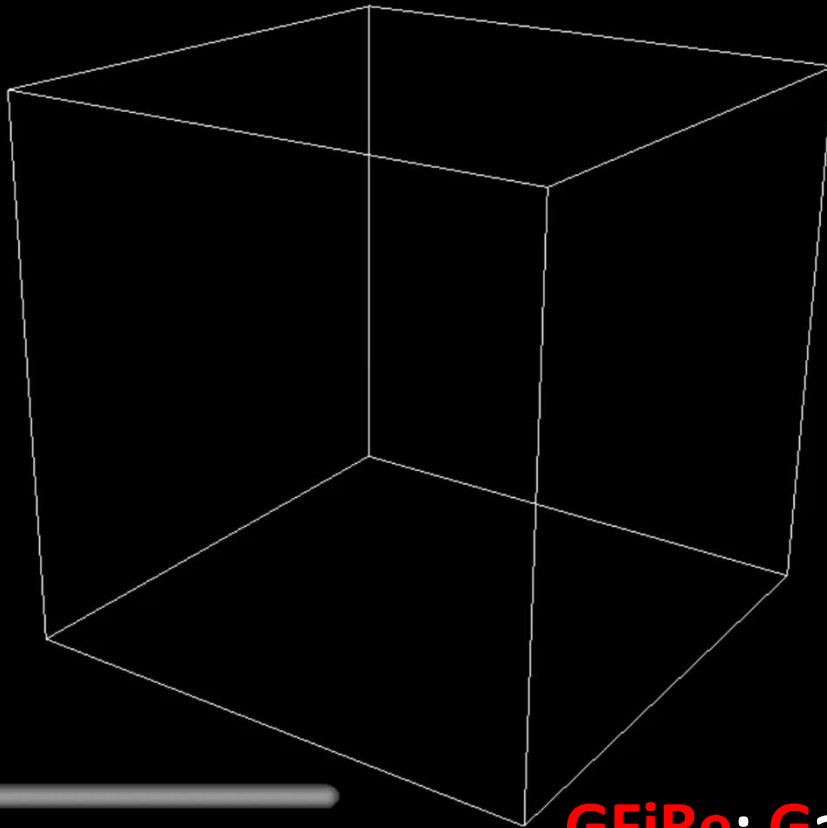


wavenumber

Gauge fields, inflation & reheating

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**Non-Abelian code in collaboration with
A. Caravano, KL, Eiichiro Komatsu, J. Weller
(2021a, 2021b, in progress)**



Abelian code: KL, M. Amin (2020)

**GFiRe: Gauge Field integrator for Reheating
NEW LATTICE CODE!!! JCAP 058 04 (2020)**

Gauge fields, inflation & reheating

Adshead and Wyman (2012)

$$S = \int d^4x \sqrt{-g} \left[\frac{1}{16\pi G} \mathcal{R} - \frac{1}{2} \partial_\mu \phi \partial^\mu \phi - V(\phi) - \frac{1}{4} F_{\mu\nu}^a F^{a\mu\nu} - \frac{\phi}{4f} F_{\mu\nu}^a \tilde{F}^{a\mu\nu} \right]$$

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Non-trivial vevs:

Gauge fields, inflation & reheating

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$$\bar{\phi}(t)$$

Gauge fields, inflation & reheating

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Drive inflation

Gauge fields, inflation & reheating

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(+extensions)

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Drive inflation
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Spectator sector

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Dimastrogiovanni et al (2016)

Maleknejad (2016, 2018)

Fujita, Sfakianakis, Shiraishi (2018)

A. Agrawal, Komatsu et al (2017, 2018)

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Gauge fields, inflation & reheating

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$$\delta A_i^b$$

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δA_i^b and

Gauge fields, inflation & reheating

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$$\delta A_i^b \text{ and } h_{ij}$$

Gauge fields, inflation & reheating

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Enhanced and chiral GWs!

Gauge fields, inflation & reheating

Adshead and Wyman (2012)

$$S = \int d^4x \sqrt{-g} \left[\frac{1}{16\pi G} \mathcal{R} - \frac{1}{2} \partial_\mu \phi \partial^\mu \phi - V(\phi) - \frac{1}{4} F_{\mu\nu}^a F^{a\mu\nu} - \frac{\phi}{4f} F_{\mu\nu}^a \tilde{F}^{a\mu\nu} \right]$$

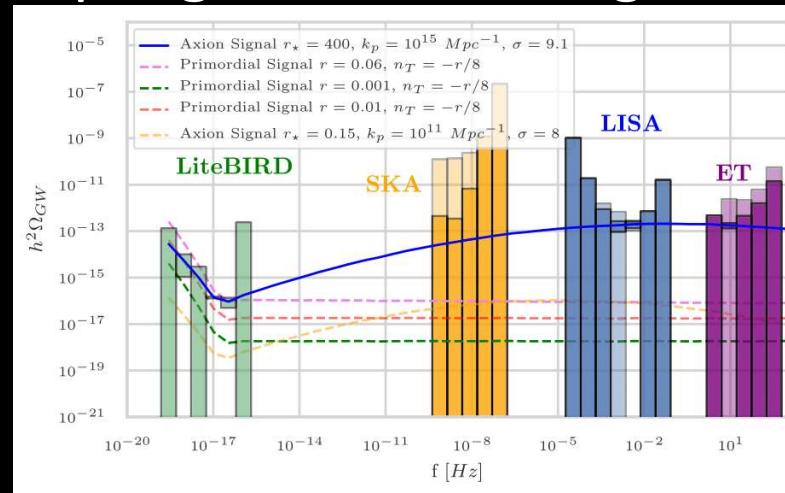
Non-trivial vevs:

$$\overline{\phi}(t)$$

$$\overline{A}_i^b(t) = a(t) Q(t) \delta_i^b$$

Linear coupling between Gauge Fields and GWs!

GW power sp.



log(frequency)

Figure from
Campeti, Komatsu et al.
arXiv:2007.04241

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See also Domcke et al (2018), Adshead, Sfakianakis (2015)

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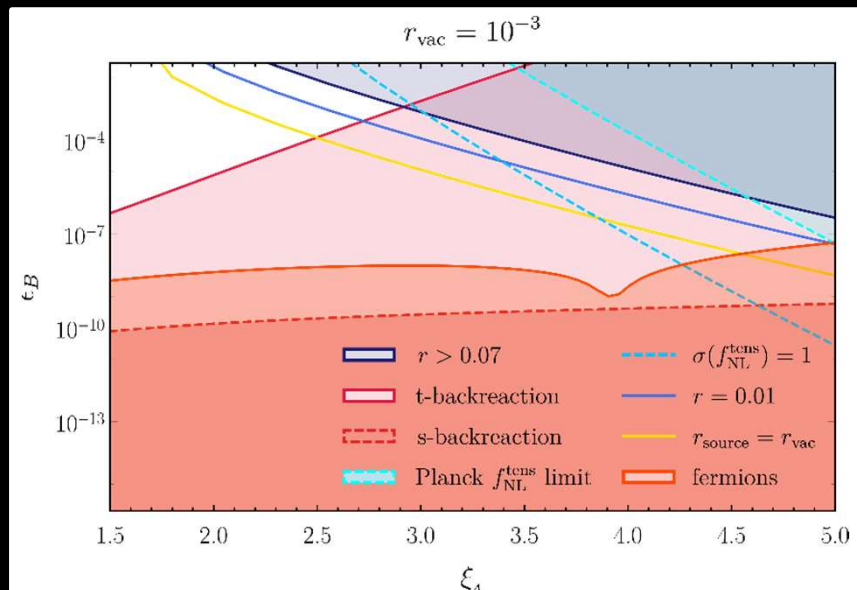
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Fraction of energy
in Gauge Fields



Gauge field coupling strength

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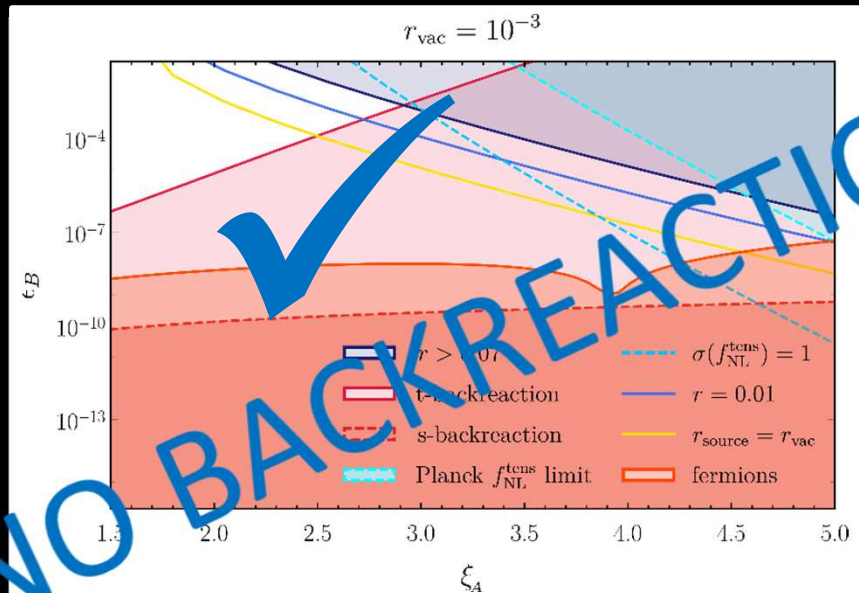
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See also

Mirzagoli, Komatsu, KL, Watanabe (2020)
(on gravitational CS)



Fuzzy Vector Dark Matter



Kaloian Lozanov
UIUC, Cosmology and HEPheo

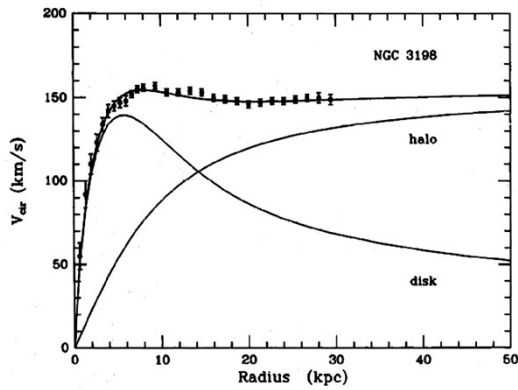


Based on: PRD 103, 103501 (2021) (Adshead and KL)

Dark Matter

Dark Matter

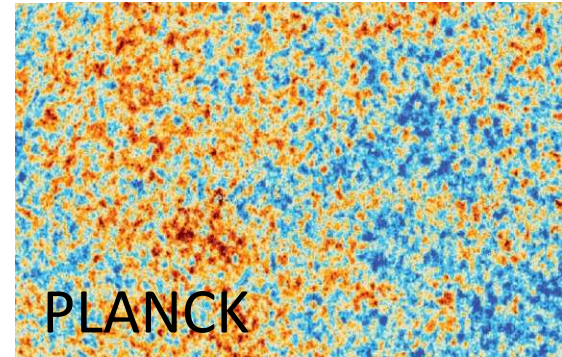
Rotation curves



LSS

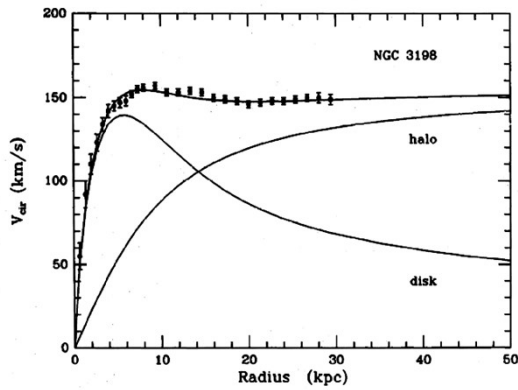


CMB



Dark Matter

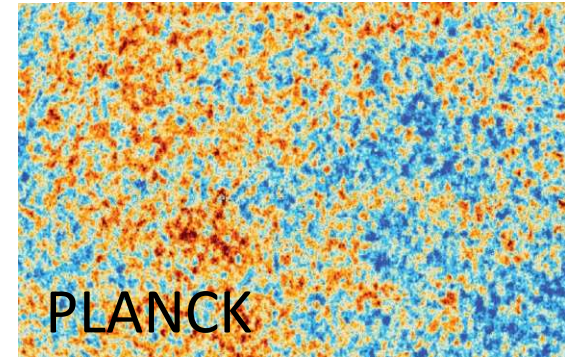
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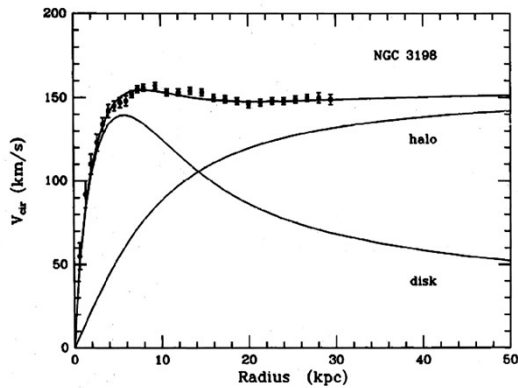
CMB



DM mass ??

Dark Matter

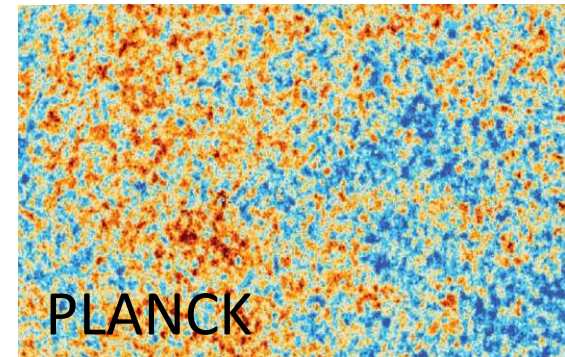
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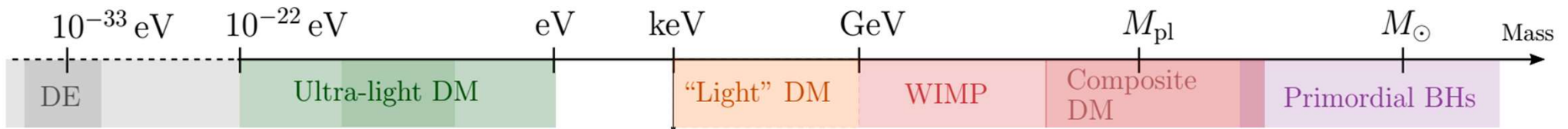
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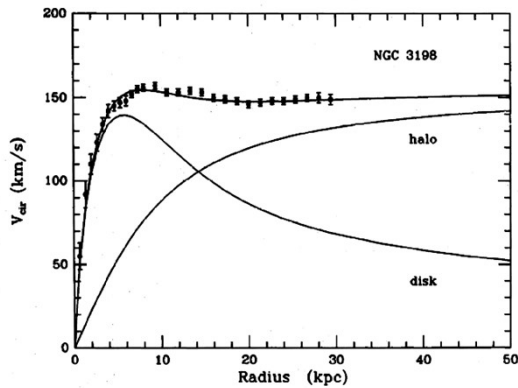


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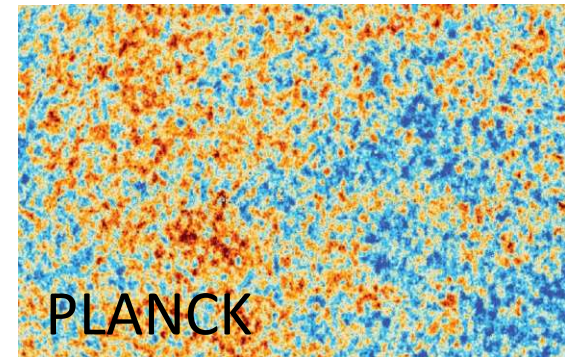
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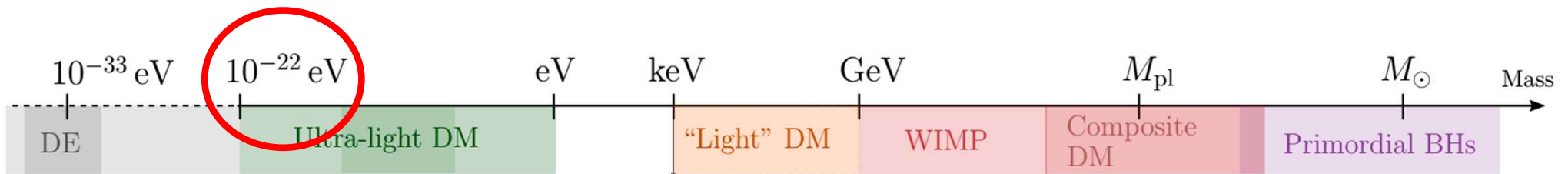


CMB



DM mass ??

Fuzzy DM



Fuzzy Dark Matter (FDM)

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- Cold and Bosonic: forms Bose-Einstein Condensate

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$$\lambda_{\text{dB}} \sim \left(\frac{M_{\text{halo}}}{10^8 M_{\odot}} \right)^{-1} \left(\frac{m}{10^{-22} \text{eV}} \right)^{-2} \text{kpc}$$

Motivation for FDM

Hu, Barkana, Gruzinov (2000)

Hui, Ostriker, Tremaine, Witten (review) (2016)

Motivation for FDM

Λ CDM small-scale challenges

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FDM solution

- No grav. collapse below λ_{dB}

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Scalar FDM

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$$\Delta \Phi = \frac{m |\psi|^2}{2m_{\text{pl}}^2}$$

$$\Phi = \Psi$$

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Scalar FDM

$$S = \int d^4x \sqrt{-g} \left[-\frac{m_{\text{pl}}^2 R}{2} + \frac{\partial_\mu \phi \partial^\mu \phi}{2} - \frac{m^2 \phi^2}{2} \right]$$



$$\frac{1}{\sqrt{-g}} \partial_\mu (\sqrt{-g} g^{\mu\nu} \partial_\nu \phi) + m^2 \phi = 0$$

$$g_{\mu\nu} = a^2(\tau) \begin{pmatrix} 1 + 2\Phi & 0 & 0 & 0 \\ 0 & 2\Psi - 1 & 0 & 0 \\ 0 & 0 & 2\Psi - 1 & 0 \\ 0 & 0 & 0 & 2\Psi - 1 \end{pmatrix}$$



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Vector FDM

Vector FDM

$$S = \int d^4x \sqrt{-g} \left[-\frac{m_{\text{pl}}^2 R}{2} - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \frac{1}{2} m^2 A^\mu A_\mu \right]$$

Vector FDM

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$$\nabla_\mu F^{\mu\nu} = -m^2 A^\nu$$

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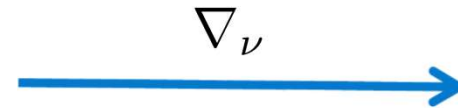
$$\nabla_\mu F^{\mu\nu} = -m^2 A^\nu \equiv J^\nu$$

Vector FDM

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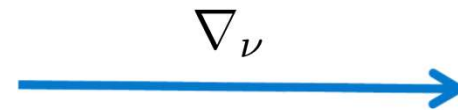


Vector FDM

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$$\square A^\nu + m^2 A^\nu + R^\nu{}_\mu A^\mu = 0$$

$$\nabla_\nu A^\nu = 0$$

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$g_{\mu\nu} = ?$



Vector FDM

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$$\begin{aligned} \Phi &= \Psi \\ V_i &= 0 \end{aligned}$$

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Vector FDM

Vector FDM

$$S = \int d^4x \sqrt{-g} \left[-\frac{m_{\text{pl}}^2 R}{2} - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \frac{1}{2} m^2 A^\mu A_\mu \right]$$

$$A^\mu = \sqrt{\frac{2}{m}} \Re[\mathcal{A}^\mu e^{-imt}]$$

$$\left[i\partial_t + \frac{3}{2}H + \frac{\Delta}{2a^2 m} - \Phi m \right] \mathcal{A}^i = 0$$

$$\Delta\Phi = \frac{m \sum_j |\mathcal{A}^j|^2}{2m_{\text{pl}}^2}$$

Vector FDM \longleftrightarrow

Vector FDM

$$S = \int d^4x \sqrt{-g} \left[-\frac{m_{\text{pl}}^2 R}{2} - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \frac{1}{2} m^2 A^\mu A_\mu \right]$$

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$$\Delta\Phi = \frac{m \sum_j |\mathcal{A}^j|^2}{2m_{\text{pl}}^2}$$

Vector FDM \iff 3 \times Scalar FDM

Vector FDM

$$S = \int d^4x \sqrt{-g} \left[-\frac{m_{\text{pl}}^2 R}{2} - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \frac{1}{2} m^2 A^\mu A_\mu \right]$$

$$A^\mu = \sqrt{\frac{2}{m}} \Re[\mathcal{A}^\mu e^{-imt}]$$

$$\left[i\partial_t + \frac{3}{2}H + \frac{\Delta}{2a^2m} - \Phi m \right] \mathcal{A}^i = 0$$

$$\Delta\Phi = \frac{m \sum_j |\mathcal{A}^j|^2}{2m_{\text{pl}}^2}$$

Dynamical eqns.

Use for:

1.) Solitonic cores

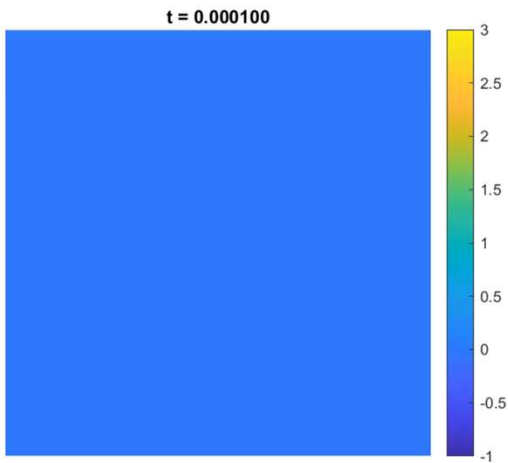
2.) FDM simulations

Vector FDM \iff 3 \times Scalar FDM

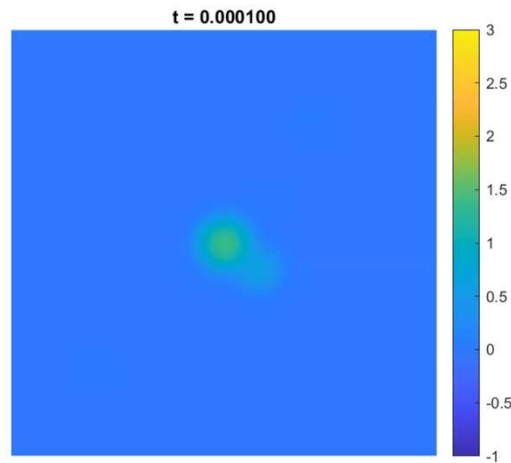
Vector FDM Simulations

$$\left[i\partial_t + \frac{3}{2}H + \frac{\Delta}{2a^2m} - \Phi m \right] \mathcal{A}^i = 0$$

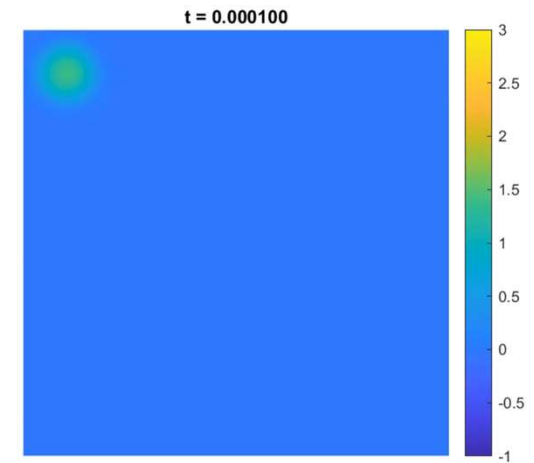
$$\Delta\Phi = \frac{m \sum_j |\mathcal{A}^j|^2}{2m_{\text{pl}}^2} \quad \text{set } H = 0$$



$\ln(|\mathcal{A}^x|^2)$



$\ln(|\mathcal{A}^y|^2)$



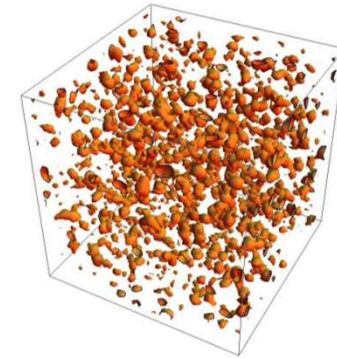
$\ln(|\mathcal{A}^z|^2)$

Work in progress

Conclusions

Reheating at the end of inflation:

- very rich dynamics



Future plans:

- more realistic models (e.g. gauge fields)
- observational signatures
 - expansion history, relics, GWs

